

Multi-Object and Long-Slit Spectroscopy of Very Low Mass Brown Dwarfs in the Orion Nebular Cluster

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Abstract

We present the results of a H - and K -band multi-object and long-slit spectroscopic survey of substellar mass candidates in the outer regions of the Orion Nebula Cluster. The spectra were obtained using MOIRCS on the 8.2-m Subaru telescope and ISLE on the 1.88-m telescope of Okayama Astronomical Observatory. Eight out of twelve spectra show strong water absorptions and we confirm that their effective temperatures are ≤ 3000 K (spectral type $\geq M6$) from a chi-square fit to synthetic spectra. We plot our sources on an HR diagram overlaid with theoretical isochrones of low-mass objects and identify three new young brown dwarf candidates. One of the three new candidates is a cool object near the brown dwarf and planetary mass boundary. Based on our observations and those of previous studies, we determine the stellar ($0.08 < M/M_{\odot} < 1$) to substellar ($0.03 < M/M_{\odot} < 0.08$) mass number ratio in the outer regions of the Orion nebular cluster to be 3.5 ± 0.8 . In combination with the number ratio reported for the central region ($3.3^{+0.8}_{-0.7}$), this result suggests the number ratio does not simply change with the distance from the center of the Orion nebular cluster.

Key words: Stars: low-mass, brown dwarfs — Galaxy: open clusters and associations: individual: Orion Nebular Cluster

1. Introduction

Brown dwarfs (BDs) are substellar mass objects between planetary and stellar mass regimes ($0.013 \leq M/M_{\odot} \leq 0.075$). Since their discovery in the 1990s, hundreds of examples have been reported, and NASA's WISE mission has recently added another hundred BDs (Kirkpatrick et al. 2011). However, since the typical thermal Jeans mass is $\sim 1 M_{\odot}$ in molecular cloud cores (Larson 1999, and references therein), the gravitational collapse might be expected to form stars, not BDs. Therefore the formation mechanism of BDs is still unclear. Several theories for the formation of BDs have been suggested, including turbulent fragmentation (Padoan & Nordlund 2004; Boyd & Whitworth 2005), magnetic field confinement (Boss 2001), stellar embryo ejection through dynamical interactions (Reipurth & Clarke 2001; Bate et al. 2002), and photo-evaporation of embryos by nearby hot stars (Whitworth & Zinnecker 2004), though none of these theories have yet proven conclusive. Star forming theories predict the existence of even planetary mass objects (PMOs) with a lowest mass limit of $0.001 - 0.010 M_{\odot}$ related to the opacity of the contracting object, known as opacity limited fragmentation (e.g., Low & Lynden-Bell 1976; Rees 1976; Boss 2001; Whitworth

& Stamatellos 2006). However, since they are not massive enough to sustain hydrogen burning and become fainter with time, substellar objects are very faint in the field. Therefore, only a few PMOs have been reported in the field (e.g., Cushing et al. 2011). In contrast, substellar objects in star forming regions have much higher luminosities and therefore, in addition to many BDs, even isolated PMOs or planetary mass companion candidates have been identified in such regions (Tamura et al. 1998; Oasa et al. 1999; Lucas & Roche 2000; Zapatero Osorio et al. 2000; Neuhäuser et al. 2008; Kuzuhara et al. 2011). In order to reveal their formation processes, it is important to uncover their properties, such as binarity, space distribution, and especially initial mass function (IMF) in star forming regions.

Recent surveys of low-mass BDs and PMOs have been conducted in nearby star forming regions. In σ Orionis, which is a well-studied field of PMOs (e.g., Caballero et al. 2007; Bihain et al. 2009), Peña Ramírez et al. (2012) have conducted a wide ($\sim 0.78 \text{ deg}^2$) and deep (translated masses down to $\sim 0.004 M_{\odot}$) optical/infrared imaging based on VISTA Orion survey data (Petr-Gotzens et al. 2011). They have reported that σ Orionis cluster harbors as many BDs (69 sources, $0.012 - 0.072 M_{\odot}$) and PMOs (37 sources, $0.004 - 0.012 M_{\odot}$) as low-

mass stars (104 sources, $0.072 - 0.25 M_{\odot}$). In the NGC1333, ρ Ophiuchi and Chamaeleon-I cores, a systematic survey, called the Substellar Objects in Nearby Young Clusters (SONYC) project, was conducted (Scholz et al. 2009; Geers et al. 2011; Mužić et al. 2011, 2012). Scholz et al. (2012) performed multi-object spectroscopy to extend their previous work over more targets by using Subaru/FMOS in NGC1333. They identified an L-type cluster member and estimated its mass to be as low as $0.006 M_{\odot}$. However, since the construction of an IMF needs a large number of photometric and spectroscopic observations, current surveys in star forming regions are incomplete in the substellar regime, particularly in the planetary-mass domain. In order to understand the formation scenario of low-mass BDs and PMOs, we need to conduct deeper surveys for more targets in these star forming regions.

The Orion nebular cluster (ONC) is the most suitable star forming region for IMF studies. It is nearby, ~ 450 pc (Muench et al. 2008), and young, ≤ 1 Myr (Hillenbrand 1997; Riddick et al. 2007). Previous studies of the ONC have concentrated on the central part of the ONC region. Hillenbrand (1997) conducted an optical imaging survey and spectroscopic follow-up, including spectra for ~ 900 stars, and measured the stellar masses down to $\sim 0.1 M_{\odot}$ within 4.5 pc \times 4.8 pc ($34' \times 36'$) of the Trapezium cluster. Luhman et al. (2000) conducted infrared imaging at the central region of the Trapezium cluster ($140'' \times 140''$; 0.3 pc \times 0.3 pc) and K -band spectroscopy for ~ 100 sources. Slesnick et al. (2004) performed 97 J and K -band spectroscopy measurements of the inner $5.1' \times 5.1'$ region of the ONC (Hillenbrand & Carpenter 2000). Although they confirmed many low-mass cluster members and even BDs of masses down to $0.02 M_{\odot}$, the bright nebulosities make it difficult to extend the measurements to lower masses. In order to cover a larger area of the ONC, deep and wide imaging surveys have been conducted (Robberto et al. 2010; Andersen et al. 2011; Da Rio et al. 2012). However, no spectroscopic observations of the targets have been made yet. Lucas et al. (2005) have made observations focused on the outer regions of the ONC, where fainter nebulosities enable observations of fainter objects. They detected 33 faint planetary mass candidates. To confirm their cluster membership, spectroscopic observations were conducted and about 10 PMO candidates were identified as young stellar objects (Lucas et al. 2006; Weights et al. 2009). However, since many of the photometric candidates have not been spectroscopically examined yet, it is important to observe the candidates for the characterization of BDs and PMOs in this region. We therefore have conducted spectroscopic follow-up observations of the Lucas et al. (2005) candidates, and found two new young BDs and a BD/Planetary-mass boundary objects.

In section 2, we explain our observations and data reduction. In section 3, we show the analysis and results. In section 4, we present the HR-diagram and the number ratio of the stellar to substellar mass objects and discuss our conclusions.

2. Observations and data reduction

2.1. Near-infrared spectroscopy

Near infrared spectra of BD candidates in the ONC were obtained through two observing programs. The first data set

was acquired on November 30, 2007, using the Multi-Object InfraRed Camera and Spectrograph (MOIRCS; Suzuki et al. 2008) mounted on the 8.2-m Subaru telescope (f/12.2 at the cassegrain focus). In this instrument, a $4' \times 7'$ field of view is covered by two HAWAII 2048 \times 2048 HgCdTe arrays with a pixel scale of 0.117 arcsec/pixel. The second data set was obtained during December 3–7, 2010, using ISLE (Yanagisawa et al. 2006, 2008), a near-infrared imager and spectrograph for the cassegrain focus (f/18) of the 1.88-m telescope at Okayama Astronomical Observatory (OAO) of the National Astronomical Observatory of Japan (NAOJ). The detector is a HAWAII 1024 \times 1024 HgCdTe array that covers a $4.3' \times 4.3'$ field of view with a pixel scale of 0.25 arcsec/pixel.

Target sources for the spectroscopy were selected from the catalogue of Lucas et al. (2005), which presents the JHK photometry of the BD candidates obtained by Gemini-South/Flamingos. The selection criteria were as follows. (1) All sources have prospective masses ranging from the hydrogen-burning limit ($< 0.075 M_{\odot}$) down to the deuterium-burning limit of $0.013 M_{\odot}$. (2) Sources that were close to or embedded in the bright nebulosity were avoided. For optimum sensitivity, objects located in areas of faint nebulosities were chosen (Figure 1 in Lucas et al. (2005)). (3) Sources were only chosen if they possessed fairly low visual absorption ($A_V < 7.5$). Although most of these objects have uncertainties (~ 0.1 mag) in their magnitude, two of them (208-736 and 215-652) have large errors (> 0.2 mag) in their J and H magnitudes. Robberto et al. (2010) have also observed these objects with ISPI, the facility infrared camera at the CTIO Blanco 4-m telescope, and have achieved better quality for 208-736 in H -band, and 215-652 in J - and H -band. Therefore, in our study we used the magnitudes of Robberto et al. (2010) instead of Lucas et al. (2005) for these objects: $H = 17.37 \pm 0.10$ for 208-736 and $J = 17.87 \pm 0.17$, $H = 17.23 \pm 0.09$ for 215-652.

Unfortunately, we could only observe relatively bright candidates because auto-guider problems and high-humidity conditions prevented a full observing program with MOIRCS. Therefore, we focus on the good signal-to-noise ratio (S/N) candidates for MOIRCS in this paper. In total, 12 BD candidates were obtained, with eight of these objects obtained by MOIRCS and four obtained by ISLE (Table 1).

We took simultaneous spectra of eight of the BD candidates using the multi-object spectroscopy (MOS) mode of MOIRCS. We used a 1-arcsec slit width and the $HK500$ grism, optimized for 1.4 – $2.5 \mu\text{m}$ and low-resolution ($R \sim 500$) spectra. To subtract the background noise, the objects were observed with slit nodding. Dome flats were obtained before a run. In order to calibrate telluric absorptions, we observed a standard star with an F-type spectrum (HD7386) in a similar air mass as the main observations for BDs.

The ISLE observations were carried out in the long-slit mode. The slit width was 1 arcsec and nod dithering was performed. A low resolution mode was selected with $R \sim 350$ in the H -band and $R \sim 450$ in the K -band. Observations on Dec 3 and Dec 4 were photometric. The observation of 065-207 on Dec 7 was taken under a condition of thin clouds and we carefully analyzed this source to reduce the effect of this condition. We employed a 1.5-arcsec slit width to acquire sufficient S/N. We obtained the dome flat before and after the observa-

tion. We selected an F-type star (HD24635) and a B-type star (HD34748) to measure the telluric absorption. When we use a B-type star possessing primarily strong hydrogen recombination lines as a telluric standard star, we should carefully deal with the intrinsic stellar lines since such lines cause artificial features in the science spectra. Therefore, we removed the intrinsic lines in the B-type stellar spectra by interpolating it with the neighboring flux.

2.2. Data reduction

Both the multiple and long slit data were reduced using IRAF software. Standard techniques were used to reduce the MOIRCS multiple slit data. In order to subtract the sky background, exposures for each BD candidate pair were separated into their relative nod positions (A or B). We subtract the adjacent B frame from the A frame to decrease the residual. Flat fielding was corrected by using the dome flat frames. Note that channel 1 of MOIRCS was the engineering detector from Oct 2007 to Jun 2008 and had prominent, large ring-like high dark noise. Therefore we masked and did not use these regions in this data reduction. Cosmic ray cleaning was performed by L.A.Cosmic (Laplacian Cosmic Ray Identification; van Dokkum 2001). Array distortion was corrected by the MCSGEOCORR task in the MCSRED¹. To calibrate the wavelength of the data, OH lines were identified as a function of wavelength. After performing 5-pixel binning along the wavelength directions of the data to improve the S/N ratio, we extracted the spectra with the APALL task in IRAF. We correct spectral contamination due to telluric absorption using the standard star measurements. Extracted pairs from each nod position were combined. Finally, a de-reddening procedure was conducted with an adopted reddening parameter of 3.1. We used the $J-H$ vs H color and magnitude of the objects and synthetic color-magnitude diagram (Allard et al. 2010) with age ~ 1 Myr and distance ~ 450 pc to derive the A_V values using the reddening law of Cardelli et al. (1989).

For the ISLE spectra, the reduction process was similar to that of MOIRCS. However, we did not correct the array distortion because the influence of the distortion was negligible. We obtained the H -band and K -band spectra separately. Therefore, the H -band and K -band spectra were scaled to the 2MASS magnitude of the objects whose Vega magnitude and zero point magnitude were based on Cohen et al. (2003). The advantages of using 2MASS magnitudes are that the 2MASS magnitudes of our objects were obtained with low errors (< 0.1 mag), and 2MASS K_s -band filter reduced the noise contribution from the thermal background beyond $2.3 \mu\text{m}$.

3. Analysis and results

To derive the effective temperature and surface gravity, we compare our observed spectra to a grid of synthetic spectra and find the best-fitting model in the χ^2 sense (e.g., Mohanty et al. 2007). First, we obtain the best-fitting parameters (§3.2.1). We then conduct a Monte Carlo simulation in order to constrain the error range of parameter estimates (§3.2.2) (e.g., Cushing et al.

2008; Bowler et al. 2009; Aller et al. 2013). We restrict the range of stellar parameters in the fit to effective temperatures $1800 \text{ K} \leq T_{\text{eff}} \leq 4900 \text{ K}$ and surface gravity $3.0 \leq \log(g) \leq 5.5$. Also, we use a solar metallicity as the average one of low-mass members in the ONC.

3.1. BT-Settl model

We use the BT-Settl model atmosphere (Allard et al. 2010), which is available via the Phoenix web simulator. Allard et al. (2010) updated the BT-Settl model applying the revised solar abundances by Asplund et al. (2009), the latest H_2O line list (Barber et al. 2008), a cloud model based on condensation and sedimentation timescales by Rossow (1978), the supersaturation computed from pre-tabulated chemical equilibrium, and mixing from 2D radiation hydrodynamic simulations by Freytag et al. (2010).

BDs have an atmosphere including water vapor which causes strong absorption in the infrared regime. Because young BDs found in Trapezium are still gravitationally contracting, they have lower surface gravities than field dwarfs. Collision induced H_2 absorption ($\text{H}_2\text{-CIA}$) over the near infrared wavelength regime is pronounced at environments of strong surface gravity. Therefore, young BDs have a triangular H -band shape due to deep H_2O and weak $\text{H}_2\text{-CIA}$ absorptions (Kirkpatrick et al. 2006).

3.2. χ^2 - fitting

3.2.1. Procedure 1 - Seek best fitting parameters

In order to determine best fitting parameters, we conduct χ^2 -fitting between the observed and BT-Settl model spectra. The results are shown in the column 3 or 4 of Table 2. The parameter ranges of the effective temperature and surface gravity are $1800 \text{ K} \leq T_{\text{eff}} \leq 4900 \text{ K}$ and $3.0 \leq \log(g) \leq 5.5$. The range of the surface gravity is predicted from BT-Settl evolutionary tracks for an age ranging from 1 Myr to 1 Gyr. Although the determinations of the metal content in the Orion region reach controversial conclusions, we use a solar metallicity as the average one of low-mass members in the ONC based on D’Orazi et al. (2009). We seek the best T_{eff} and $\log(g)$ to minimize the reduced χ^2 value:

$$\chi^2 = \frac{1}{N-2} \sum \left(\frac{C f_{\text{mod}} - f_{\text{obs}}}{\sigma_{\text{obs}}} \right)^2 \quad (1)$$

where N is the bin number of the wavelength, f_{mod} is the flux of the model spectra, f_{obs} is the flux of the observed spectra, and σ_{obs} the error of the observed spectra. The error is derived from the standard deviation of 4 combined spectra. C is the scaling coefficient of f_{mod} with f_{obs} and is derived from $d\chi^2/dC = 0$ for Equation 1:

$$C = \frac{(\sum f_{\text{obs}} f_{\text{mod}} / \sigma_{\text{obs}})^2}{\sum (f_{\text{mod}} / \sigma_{\text{obs}})^2} \quad (2)$$

In order to minimize telluric contamination from water vapor, we restrict the regions for the fit $1.50\text{--}1.78 \mu\text{m}$ and $2.10\text{--}2.29 \mu\text{m}$ in the case of MOIRCS, and $1.5\text{--}1.78 \mu\text{m}$ and $2.10\text{--}2.38 \mu\text{m}$ for the ISLE data. 061-400, 215-652 and 216-540 have emission lines in the K -band spectra (H_2 : $2.12, 2.22 \mu\text{m}$; Br γ : $2.16 \mu\text{m}$), which make the χ^2 value worse when we include the K -band data in the fit. Therefore, we exclude the emission

¹ MCSRED was provided by Dr. Ichi Tanaka. See <http://www.naoj.org/staff/ichi/MCSRED/mcsred.html> software for MOIRCS imaging.

Table 1. Summary of observed objects

Object	Data Set	<i>J</i>	<i>H</i>	<i>K</i>	A_V	Obs time	Reference ^a	ID(LRT05) ^b
		[mag]	[mag]	[mag]		<i>H,K</i> [min]		
030-524	MOIRCS	17.88	17.26	16.77	1.8 ± 1.5	60,60	1,2	178
037-628	MOIRCS	18.52	17.66	16.99	4.4 ± 1.5	60,60	2	94
061-400	MOIRCS	18.31	17.33	16.48	5.7 ± 1.5	60,60	2	270
065-207	ISLE	13.47	12.81	12.24	$0.8^{+1.5}_{-0.8}$	10,60	-	382
072-638	ISLE	15.39	14.80	14.29	$1.0^{+1.5}_{-1.0}$	10,-	1	70
099-411	MOIRCS	17.08	16.40	15.99	2.3 ± 1.5	60,60	-	256
104-451	ISLE	15.10	13.99	13.13	5.8 ± 1.5	5,30	-	230
183-729	MOIRCS	18.05	17.38	17.24	2.3 ± 1.5	60,60	1,2	18
208-736	MOIRCS	18.28 ^c	17.37 ^d	16.56	5.0 ± 2.4	60,60	-	8
215-652	MOIRCS	17.87 ^d	17.23 ^d	17.15	2.0 ± 2.0	60,60	-	56
216-540	MOIRCS	16.95	15.89	15.47	3.7 ± 1.5	60,60	-	155
217-653	ISLE	14.99	14.26	13.32	2.4 ± 1.5	5,30	1	55

Note. Object names are coordinate based, following O'dell & Wong (1996). 030-524 means that the coordinate of the object is (R.A., Decl.) = (05^h35^d03^s.0, -05°25'24"). *JHK* magnitudes are taken from Lucas et al. (2005) with the exception of 208-736 and 215-652. Visual absorption (A_V) is estimated from a comparison between observed and synthetic *J-H* vs *H* color-magnitude diagrams.

^a The previous spectroscopic studies. 1. Riddick et al. (2007); 2. Weights et al. (2009)

^b The running number is used in Lucas et al. (2005)

^c The magnitude has a large uncertainty (> 0.2 mag)

^d The magnitude is taken from Robberto et al. (2010) instead of Lucas et al. (2005)

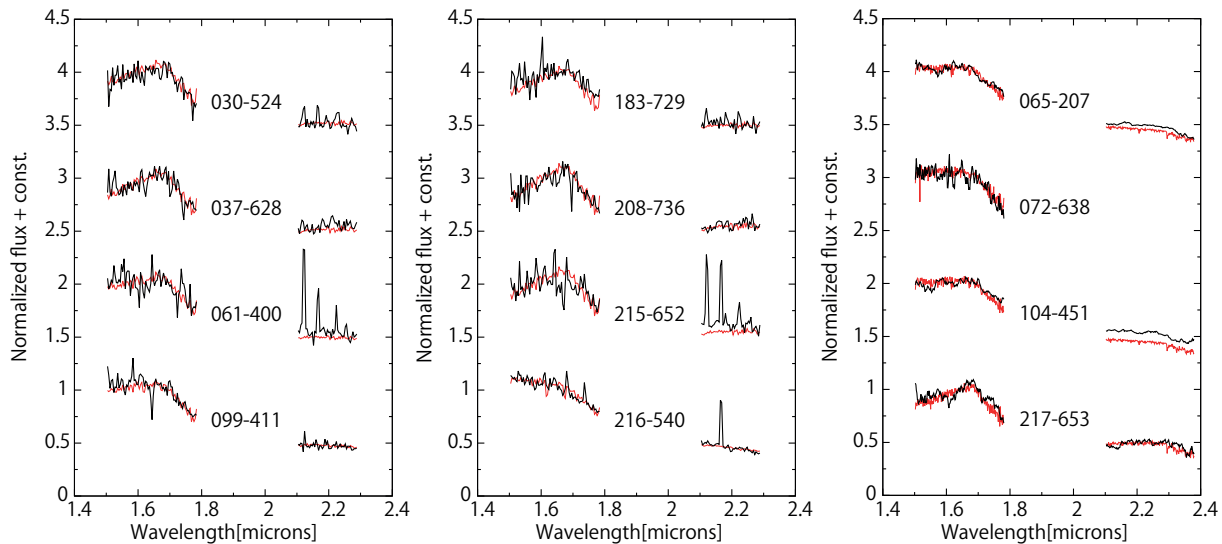


Fig. 1. Near infrared observed spectra. The left two panels show spectra obtained by MOIRCS and the right panel shows spectra obtained by ISLE. The observed (dereddened) spectra are shown as black solid lines while the best-fit model spectra are shown as red lines. The spectra of 061-400, 215-652 and 216-540 exhibit emission lines (H_2 : 2.12, 2.22 μ m; Br γ : 2.16 μ m).

lines in the fitting process. Also, we can see spiky substructures in some H -band spectra. For example, there is the spiky bump around $1.7 \mu\text{m}$ in the spectrum of 208-736, which could be caused by partly very low-S/N ($\sim 1 \sigma$). Such substructure affected by poor S/N make χ^2 value worse. When both H and K spectra exist, we perform χ^2 -fitting in two cases: using only the H -band and using both the H - and K -bands.

After the fit, we adopt the best-fitting parameters as those that minimize the χ^2 value. We show the selected wavelength (H or HK) resulting in the best-fitting in column 5 of Table 2. In many cases we use both the H - and K -bands to increase the sampling points, except for 065-207, 072-638 and 104-451. We adopt the result of fitting only using H -band for 072-638 due to a lack of K -band spectrum. Also, we employ the results of H -band for 065-207 and 104-451 because, in the fitting by only the H -band, we find a systematic offset between the observed and synthetic K -band spectra. The offset may be explained by the K -band excess from a circumstellar structure. Since the excess due to the circumstellar structure is more apparent in the longer wavelength, we check L' -band data of the objects. Lada et al. (2004) presented the L' -band data, $L' = 11.69$ for 065-207 and $L' = 12.27$ for 104-451, by using VLT/ISAAC, which enable us to derive the dereddened photometric colors $H - L' = 1.02$ and $H - L' = 0.99$ respectively. Then, we estimate the synthetic colors $H - L'$ with the theoretical spectra fitting with H -band for 065-207 and 104-451, and with the H and L' filter response curves, which results in $H - L' = 0.71$ (065-207) and $H - L' = 0.70$ (104-451). The photometric colors of these objects are adequately redder than the synthetic colors ($\Delta(H - L') > 0.3$), which is strong evidence of a circumstellar structure.

3.2.2. Procedure 2 - Estimate the uncertainty

We perform a Monte Carlo simulation, the parametric bootstrapping, to find the T_{eff} distribution, allowing us to determine the error range of T_{eff} . In the simulation, we conduct a lot of χ^2 -fittings between model spectra and randomly generated mock spectrum. There is an uncertainty in the parameter estimate from the χ^2 -fitting due to both the S/N of the observed spectra and the uncertainty of visual absorption (A_V) estimate. Therefore, we separately evaluate the impact of these uncertainties on the derived T_{eff} . We show the values of T_{eff} and errors in column 6 of Table 2. The procedures for deriving these values are as follows:

- (1) In order to conduct a χ^2 -fitting, we simulate a mock spectrum by adding an artificial noise to the theoretical spectrum. The parameters of the theoretical spectrum are determined based on the values derived in §3.2.1 and described in column 3 or 4 of Table 2. For example, in the case of 030-524, we use the theoretical spectrum with $T_{\text{eff}} = 2600 \text{ K}$ and $\log(g)=3.0$. The artificial noise is randomly generated based on the error of observed spectrum, which means that we model the artificial noise at each wavelength as a drawn from a Gaussian distribution with a standard deviation adjusted to the 1σ error of the observed spectrum at the corresponding wavelength. Note that this procedure to simulate the mock spectra differs from the previous studies (e.g., Cushing et al. 2008) which simulated their mock spectra by adding artificial noise to the observed spectra.
- (2) We determine the best-fitting parameters for the mock

spectrum using χ^2 fitting.

- (3) We repeat 10,000 iterations of 1 – 2 steps, making the distribution of T_{eff} . We fit a Gaussian function to the distribution of T_{eff} measured in step 2. The error range of T_{eff} should be treated carefully since the error range is affected by the S/N of the spectrum and the uncertainty of the visual absorption A_V . By taking into account of the S/N of the spectrum, we adopt the 1σ confidence range of the Gaussian distribution as the error range $\Delta T_{\text{eff},S/N}$.
- (4) In order to evaluate the impact of the uncertainty ΔA_V on the error range of T_{eff} , we conduct the χ^2 -fitting using three different A_V ; the average value A_V , the upper limit $A_V + \Delta A_V$ and the lower limit $A_V - \Delta A_V$. For example, in the case of 030-524, we use three spectra dereddened with $(A_V - \Delta A_V, A_V, A_V + \Delta A_V) = (0.3, 1.8, 3.3)$. For the three spectra, we conduct the χ^2 -fitting as described in §3.2.1 to seek the best fitting parameters.
- (5) We obtain three effective temperatures: $T_{\text{eff}}(A_V)$, $T_{\text{eff}}(A_V + \Delta A_V)$ and $T_{\text{eff}}(A_V - \Delta A_V)$. In order to describe the error affected by ΔA_V , we define the upper error value as $\Delta T_{\text{eff},A_V,+} \equiv T_{\text{eff}}(A_V + \Delta A_V) - T_{\text{eff}}(A_V)$, and the lower error value as $\Delta T_{\text{eff},A_V,-} \equiv T_{\text{eff}}(A_V) - T_{\text{eff}}(A_V - \Delta A_V)$.
- (6) We choose the larger of $\Delta T_{\text{eff},S/N}$ and $\Delta T_{\text{eff},A_V,\pm}$ as our T_{eff} uncertainty. In most cases, our error is dominated by the uncertainty in A_V .

3.2.3. Results

The obtained spectra and best-fitting theoretical spectra from §3.2.2 are shown in Figure.1, and the estimated physical parameters are shown in Table 2. Most objects are well reproduced by the BT-Settl models. However, 216-540 required a change to the extinction value $\Delta A_V \sim 2$ to obtain better χ^2 values, which results in $T_{\text{eff}} = 3320_{-270}^{+760} \text{ K}$.

Although some objects have slightly higher T_{eff} than those estimated by the previous works (037-628 and 188-739), our result is largely consistent with those of previous works (Riddick et al. 2007; Weights et al. 2009). Note however that $T_{\text{eff}} = 2750_{-170}^{+250} \text{ K}$ of 061-400 is much higher than the previous result $T_{\text{eff}} \leq 2400 \text{ K}$ (Weights et al. 2009). In Weights et al. (2009), the photometry and spectroscopy for this object were originally based on Lucas et al. (2001), in which they adopted $J = 18.250$ and $H = 17.706$. If we adopt the same magnitudes, we obtain $A_V = 0$, which is lower than the adopted value $A_V = 5.7$ in this work, and the lower A_V results in a lower temperature. However, Robberto et al. (2010) provided a magnitude of $J = 18.35 \pm 0.20 \text{ mag}$ and $H = 17.46 \pm 0.10 \text{ mag}$. Our adopted magnitudes are consistent with these more recent values. Thus, we use our result for 061-400 in the following discussions.

4. Discussion

Our observations identify new very-low-mass objects in the Orion nebular cluster. Here we plot T_{eff} and absolute H magnitude on a synthetic HR diagram. As a result, we obtain three new young BD candidates: 099-411 is an object with boundary mass between stellar and substellar mass, 215-652 is a BD candidate with emission lines associated with youth, and 208-736 is a very low mass object which is even a BD/Planetary-mass

object. In this section we will discuss the properties of each object based on the HR diagram, and then derive a number ratio of the stellar to substellar mass objects and discuss the star formation process in the outer region of the ONC.

4.1. HR diagram

We estimate the masses of the BD candidates by plotting the effective temperature and the absolute H magnitude on an HR diagram overlaid with theoretical isochrones (Figure 2). The error range of the mass corresponds to the error of the effective temperature in the HR diagram. We show the derived masses and the errors in Table 2. In order to convert the apparent magnitudes to the absolute magnitudes, we employ 450 pc ($DM = 8.26$) as the ONC distance. This value is the average distance in the previous investigation (Muench et al. 2008). The error of the distance is at most 50 pc ($DM = 0.25$) but the error is not significant compared with the error of the effective temperature in deriving the mass.

In Figure 2, it appears that there are several old age objects (~ 10 Myr) although the ONC is a very young cluster, ≤ 1 Myr (Hillenbrand 1997). In order to explain the old population, we consider several scenarios. The first is that we are seeing contamination from another foreground population. It is known that the ONC is neighbored by several group of stars. Alves & Bouy (2012) have investigated the surrounding population of the ONC, and reported the contiguous population, NGC 1980, overlaps significantly with the ONC. Since they have shown that NGC 1980 can make up for more than 10–20 % of the ONC population, the projection of the foreground population may generate the age dispersion. However, its members are thought to be only $\sim 4–5$ Myr and therefore cannot account for the 10 Myr population on the HR diagram. Second we consider the case that the older objects are foreground galactic field stars. In particular, it is standing that 030-524 and 216-540 are older than 10 Myr. Since the local number density peaks at M3-M4 (Cruz 2008), such objects are often contaminated. In contrast, it seems difficult to be contaminated by fainter objects since the local number density steeply decreases for the later spectral type ($\geq M5$), and the dense background molecular clouds prohibit penetration of the light of the very cool stars. In order to confirm how frequently field stars can contaminate our sample, we calculate the probability of the contamination in the spectral type range of our interest.

Among 12 objects observed in our work, eight of those have $2400 \text{ K} < T < 3000 \text{ K}$, corresponding from M9 to M6 (Luhman et al. 2003). First, we estimate the number of contaminating field stars with the spectral type M6 to M9. For the purpose, we adopt the same method as Caballero et al. (2008) with the FOV of MOIRCS $\sim 4' \times 7'^2$. We conservatively adopt a magnitude range of $12 < H < 18$ to include the brightest magnitude $H = 12.8$ mag (065-207) and the faintest magnitude $H = 17.7$ mag (037-628) in our work. For the ONC, we use $E_{B-V} \sim 0.32$ from Lee (1968). As a result, the number of contaminating field stars is 0.23.

We estimate the probability of contamination by using the Poisson distribution with a mean of 0.23. The probability of de-

tecting at least one contaminant is 20.3%. This means that one field BD having the similar temperature to 030-524 may be observed in our small field, and thus the number of contaminants is not significant for M6–M9 dwarfs. However, when M3–M5 dwarfs are included, the expected number of contaminants increases to 11.5, so the probability of early M contaminant with a similar temperature to 216-540 is quite high. Therefore, we cannot make a strong statement about the membership of the M3–M5 objects from the contaminant probability.

Next, we verify if the objects which seem to be older in photometry have any evidence of youth. Since protostellar objects are still experiencing gravitational contraction, they have a low surface gravity. 030-524 has a sharp triangular H -band spectrum, which is evidence of low gravity and therefore a protostar undergoing gravitational contraction. 216-540 shows a Bry emission in the K -band, which is associated with the mass accretion process and known as a common mass accretion tracer (Muzerolle et al. 1998). 061-400 and 215-652 also have an emission line in the K -band spectra, which means that these objects are mass accreting protostars. As has been noted, several objects have spectral features consistent with being young objects although they seem to be old on our HR diagram. One explanation for this discrepancy is that these objects are protostars with edge-on disks which scatter light from the central star and decrease its apparent magnitude. In fact, the HST/ACS study of Ricci et al. (2008) reported that 037-628, 061-400, 215-652 and 216-540 have ionized disks seen in emission. The existence of such objects has been indicated by previous surveys not only in the ONC (Slesnick et al. 2004) but also in other regions (e.g., Briceño et al. 2002; Luhman et al. 2003). For example, in two T Tauri systems HH30 and HK Tau C with edge-on circumstellar disks, the K -band magnitude is more than 3 mag lower than expected for their age, distance and spectral type (Burrows et al. 1996; Stapelfeldt et al. 1998). Consequently, the apparently older objects can be explained as young stellar objects with edge-on disks. Also, the episodic accretion process may explain the low luminosity of the apparently older objects, even without edge-on disks (Baraffe et al. 2009). However, in order to reinforce the validity of their membership, we need a further evidence of youth. For example, we should investigate the infrared excess due to the circumstellar disks at longer wavelengths.

4.2. Stellar to substellar mass number ratio

Here, we derive the stellar to substellar mass number ratio R , which can be a clue to discuss a possibility of difference in the formation process between stars and substellar objects (e.g., Luhman et al. 2003; Andersen et al. 2008). In previous works, there are two definitions of R , which are given as,

$$R_1 = \frac{N(0.08 < M/M_\odot < 1)}{N(0.03 < M/M_\odot < 0.08)} \quad (3)$$

$$R_2 = \frac{N(0.08 < M/M_\odot < 10)}{N(0.02 < M/M_\odot < 0.08)} \quad (4)$$

In this paper, we derive R following the above definitions. We measure the R_1 and R_2 according to the following steps. (1) We choose objects with $A_V < 7.5$ out of Lucas et al. (2005) whose observation was conducted for the $1.7'–5.7'$ regions from the ONC center. The completeness limit for A_V is

² Since Caballero et al. (2008) did not provide information for the H -band, we quote the $J-H$ color from West et al. (2008).

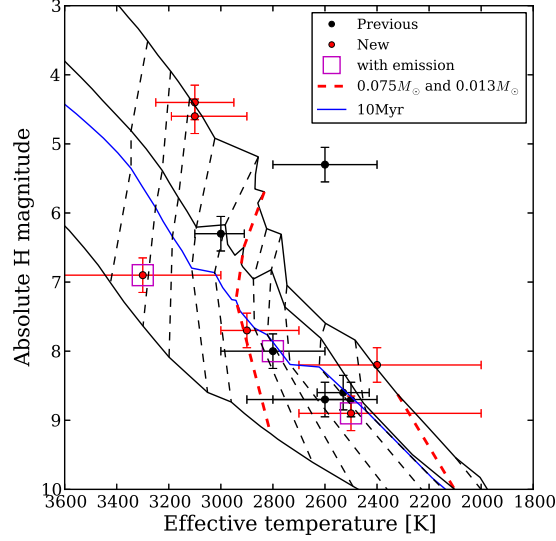


Fig. 2. T_{eff} vs absolute H -band HR diagram for 12 objects in our spectroscopic sample. The lines indicate the synthetic HR-diagram (BT-Settl; Allard et al. 2010). The solid lines show the age isochrones 1, 5, 10, 100 Myr from top to bottom. The dashed lines show the mass isochrones 0.007, 0.013, 0.02, 0.03, 0.04, 0.05, 0.06, 0.075, 0.1, 0.11, 0.15, 0.2, 0.3 M_{\odot} from right to left (the red dashed lines indicate 0.013 and 0.075 M_{\odot}). The red dots mean newly observed objects and the black dots mean follow-up objects from previous works (Riddick et al. 2007; Weights et al. 2009). The magenta squares represent the objects with emission lines.

Table 2. The result of χ^2 -fitting

Name	M_H^a	H^b		HK^b		Fit ^c	T_{eff} [K] ^d	Mass [M_{\odot}] ^e	Reference	Reference		
		T_{eff}	$\log(g)$	χ^2	T_{eff}				$\log(g)$	χ^2	SpT ^f	T_{eff} [K] ^g
030-524	8.7	3000	3.0	0.9	2600	3.0	2.1	HK	2600^{+300}_{-100}	$0.037^{+0.049}_{-0.015}$	$M8.0 \pm 0.75$	2710^{+127}_{-233}
037-628	8.6	2700	3.5	1.7	2500	3.0	1.7	HK	2500^{+100}_{-100}	$0.022^{+0.006}_{-0.005}$	$>M9.5$	< 2400
061-400	8.0	3400	3.0	1.8	2800	3.5	3.0	HK	2800^{+200}_{-200}	$0.046^{+0.056}_{-0.018}$	$>M9.5$	< 2400
065-207	4.4	3100	4.5	0.3	2900	4.0	1.0	H	3100^{+150}_{-150}	$0.145^{+0.97}_{-0.039}$	-	-
072-638	6.3	3000	5.5	0.9	- - -	-	-	H	3000^{+100}_{-90}	$0.097^{+0.018}_{-0.019}$	$M6.5 \pm 0.5$	2935 ± 55
099-411	7.7	2900	4.5	4.8	2900	4.5	3.1	HK	2900^{+100}_{-200}	$0.062^{+0.030}_{-0.026}$	-	-
104-451	4.6	3100	4.0	1.0	2500	3.0	2.8	H	3100^{+90}_{-200}	$0.134^{+0.078}_{-0.031}$	-	-
183-729	8.7	2800	3.0	1.9	2500	3.0	2.0	HK	2500^{+300}_{-100}	$0.022^{+0.035}_{-0.006}$	$\geq M9.5$	< 2400
208-736	8.2	2500	3.0	1.4	2400	3.0	1.7	HK	2400^{+300}_{-400}	$0.016^{+0.022}_{-0.010}$	-	-
215-652	8.9	2900	3.5	2.8	2500	3.0	4.2	HK	2500^{+200}_{-500}	$0.022^{+0.027}_{-0.016}$	-	-
216-540	6.9	3300	5.0	3.8	3300	5.0	2.6	HK	3300^{+700}_{-300}	> 0.100	-	-
217-653	5.3	2900	3.0	2.3	2600	3.0	1.7	HK	2600^{+200}_{-200}	$0.031^{+0.028}_{-0.015}$	$M7.75 \pm 0.5$	2752^{+80}_{-120}

^a Absolute M_H magnitudes.

^b The values are derived from χ^2 -fitting for seeking best fitting parameters (see §3.2.1). The T_{eff} and $\log(g)$ are used for simulating the mock spectrum (see §3.2.2).

^c This column shows the wavelength used for determine best fitting parameters (see §3.2.1).

^d The best T_{eff} and the uncertainty which is derived from the Monte Carlo simulation (see §3.2.2).

^e Masses derived from the HR diagram (see §4.1 and Figure 2).

^f The reference of spectral type (SpT) is shown in Table 1. When the object has been observed both in Riddick et al. (2007) and Weights et al. (2009), we use the SpT of Weights et al. (2009).

^g These quantities were derived using the reference spectral type to effective temperature scale of Luhman et al. (2003)

Table 3. Star to substellar number ratio in the orion nebular cluster

reference	R_1	R_2	r
This work	3.5 ^a (2.7-4.3)	3.2 ^a (2.6-3.9)	1.7' < r < 5.7'
Luhman et al. (2003)	-	3.8	r < 1.5'
Slesnick et al. (2004)	3.3 ^b (2.6-4.1)	5.0	r < 2.5'
Andersen et al. (2011) inner	7.2 (1.6-12.8)	-	2.9' < r < 5.8'
Andersen et al. (2011) cluster	2.4 (2.2-2.6)	-	2.9' < r < 12.5'

^aThe error range is estimated in the same manner as Scholz et al. (2012).

^b Andersen et al. (2008) have compiled the study of Slesnick et al. (2004) and shown this value.

determined to be $A_V = 7.5$. (2) We consider that objects with spectroscopically known properties are fiducial sample to investigate R . However, the sample which we spectroscopically observed is still insufficient for a statistical discussion. In order to increase the sample size, we have compiled information of spectroscopic observations of the corresponding $A_V < 7.5$ sample from previous studies (Hillenbrand 1997; Lucas et al. 2006; Riddick et al. 2007; Weights et al. 2009). Also, non-member stars are excluded from the sample based on the previous studies. Hillenbrand (1997) conducted an optical imaging and spectroscopic survey for stellar mass objects in a large region of ONC, and almost 80% complete down to $0.1 M_\odot$. Lucas et al. (2006, infrared), Riddick et al. (2007, optical) and Weights et al. (2009, infrared) are spectroscopic studies for the substellar mass candidates ($< 0.075 M_\odot$) found by Lucas et al. (2001) and Lucas et al. (2005). (3) In order to avoid inconsistencies from different mass determinations, we recalculate the masses of individual samples using the HR-diagram as explained in §4.1, instead of the masses derived in the previous papers. (4) We correct for the bias due to sample incompleteness of the spectroscopic survey. As described above, Hillenbrand (1997) observed stellar mass objects, but the other studies have concentrated on substellar mass objects. As a result, the spectroscopic completeness is 68% in the stellar mass regime and 58% in the substellar mass regimes as defined for the calculation of R_1 . Also, the completeness is 72% in the stellar mass regime and 55% in the substellar mass regime used for R_2 . When we estimate the number ratio of the stellar to substellar mass, the bias of the completeness causes the number ratio to be overestimated. Therefore, in order to correct the bias, we add three hypothetical substellar objects to calculate R_1 and five hypothetical substellar objects to calculate R_2 according to the completeness of the stellar mass regime. Consequently, we determine $R_1 = 90/26 = 3.5 \pm 0.8$ and $R_2 = 110/34 = 3.2^{+0.7}_{-0.6}$ with 1.7'–5.7' regions from the ONC center, and the errors are estimated using the method of Scholz et al. (2012) who use the confidence interval provided by Cameron (2011).

Scholz et al. (2012) listed R values for various star forming regions. In comparison with the other star forming regions, our derived R_1 and R_2 are second lowest values, which means that the number ratio of substellar mass objects in the ONC is relatively high compared with the other star forming regions. In the ONC, the R has been investigated from its center to outer edge. Andersen et al. (2011) have discontinuously covered the large area ($26' \times 33'$) using HST/NICMOS Camera 3. In the total of the observed area, they have shown $R_1 = 2.4 \pm 0.2$ whose small value suggests a flat IMF in the substellar mass

regime. R_1 (Andersen, *cluster*) is lower than our derived R_1 (Suenaga) = 3.5 ± 0.8 , which means that in a wide field of view the number of substellar objects is significantly high in the ONC. Meanwhile, Da Rio et al. (2012) have also covered a large field of view ($32' \times 33'$) with WFI camera mounted on the 2.2-m MPG/ESO telescope. They have constructed IMF down to $0.02 M_\odot$ which declines steeply with decreasing mass, and they have found no evidence of IMF flattening toward the substellar mass regime. The inconsistency between Andersen et al. (2011) and Da Rio et al. (2012) may be explained by a different estimate of background contamination. In order to confirm the validity of these studies, we should extend the coverage area of our spectroscopic study with the multi-object spectrographs.

Considering the inner regions ($r < 2.5'$) in the ONC, Luhman et al. (2003) have measured $R_2 = 3.8$. The properties of brighter objects ($K < 12$) in their sample are taken from K -band spectroscopy, however the fainter substellar objects are mostly based on photometric data with HST/NICMOS camera 3 (Luhman et al. 2000). On the other hand, the inner regions was spectroscopically investigated by Slesnick et al. (2004), who has indicated R_2 (Slesnick) = 5.0. Andersen et al. (2008) have compiled the study of Slesnick et al. (2004), and shown R_1 (Slesnick) = $3.3^{+0.8}_{-0.7}$ in this region. Our derived R_1 (Suenaga) = 3.5 ± 0.8 is consistent with the value R_1 (Slesnick), but R_2 (Suenaga) = $3.2^{+0.7}_{-0.6}$ is lower than R_2 (Slesnick). Since our observations focus on the outer regions of the ONC, in contrast to Slesnick et al. (2004), the central concentration of the high mass population ($> 1M_\odot$) found by Hillenbrand & Hartmann (1998) may explain the difference in measure R_2 values.

The outer regions ($2.9' < r < 12.5'$) were observed by Andersen et al. (2011). They reported that R_1 (Andersen, *inner*) = 7.2 ± 5.6 at the 2.9'–5.8' area is, although having the large error, consistent with the value R_1 (Slesnick) of the inner part ($r < 2.5'$). In addition, Andersen et al. (2011) found that the number ratio R decreases as a function of radius. Our result suggests that R_1 (Suenaga) = 3.5 ± 0.8 of the outer part ($1.7' < r < 5.7'$) is consistent with the R_1 (Slesnick) of the inner area of the ONC ($r < 2.5'$). This result suggests that the cluster distribution does not decrease simply as a function of radius.

To summarize, we obtained 12 spectra with the multi-object slit and long-slit spectrograph. We confirm eight of those show strong water absorptions and have very low temperatures (< 3000 K). Two of cold objects are new young BDs and one of the objects is a BD/Planetary-mass boundary

object. We estimate the stellar to substellar mass number ratio to be $R_1 = 3.5 \pm 0.8$ and $R_2 = 3.2 \pm 0.6$, the former is consistent with the value of the inner part R_1 , but the latter is lower than the inner part R_2 . Although we add 12 results of BD candidates to the spectroscopic sample, the follow-up observations of additional targets are ongoing. When the survey is completed, by the multi-object slit spectrograph, we can reveal the low mass end of the IMF.

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