

Suppression of Excited-State Contributions to Stellar Reaction Rates

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(Dated: August 28, 2013)

It has been shown in previous work [Phys. Rev. Lett. **101**, 191101 (2008); Phys. Rev. C **80**, 035801 (2009)] that a suppression of the stellar enhancement factor (SEF) occurs in some endothermic reactions at and far from stability. This effect is re-evaluated using the ground-state contributions to the stellar reaction rates, which were shown to be better suited to judge the importance of excited state contributions than the previously applied SEFs. An update of the tables shown in Phys. Rev. C **80**, 035801 (2009) is given. The new evaluation finds 2350 cases (out of a full set of 57513 reactions) for which the ground-state contribution is larger in the reaction direction with negative reaction Q value than in the exothermic direction, thus providing exceptions to the commonly applied Q value rule. The results confirm the Coulomb suppression effect but lead to a larger number of exceptions than previously found. This is due to the fact that often a large variation in the g.s. contribution does not lead to a sizeable change in the SEF. On the other hand, several previously identified cases do not appear anymore because it is found that their g.s. contribution is smaller than inferred from the SEF.

PACS numbers: 26.50.+x

I. INTRODUCTION

Astrophysical reaction rates are central to any investigation in nucleosynthesis as they provide the quantitative description of temporal abundance changes in a hot plasma. They are related to reaction cross sections which, in turn, may be predicted in theoretical models or extracted from experiments. Compared to standard nuclear physics investigations of reactions on the ground state (g.s.) of nuclei, a treatment of reaction rates is complicated by the thermal population of excited states in the nuclei which are in thermal equilibrium with their hot environment. On the other hand, as long as thermal equilibrium is upheld, there is a mathematically rigorous reciprocity relation connecting the rates of forward and reverse reaction [1–3]. This implies that only the *stellar rate* for one direction, say, $A(a,b)B$ has to be determined by theoretical or experimental means. The *stellar rate* for its inverse reaction $B(b,a)A$ automatically follows by applying the simple reciprocity relations.

Determining contributions of reactions proceeding on excited states of nuclei to the stellar rate remains an experimental challenge. Therefore it is interesting to theoretically investigate not only the magnitude of the contributions for each reaction but also which reaction direction is impacted less by such contributions. A determination of the rate in that direction would then be closer to the actually required stellar value. It is easy to show that fewer excited state transitions are involved in a stellar rate of a reaction with positive reaction Q value than in one with negative Q value. This is especially important for intermediate mass and heavy nuclei, whereas

in light nuclei excited state contributions are often small in either reaction direction, anyway. The well-known “ Q value rule” follows directly, stating that for astrophysical purposes it is preferable to use the direction of positive Q value [2–5]. It has been shown recently, however, that an effect termed “Coulomb suppression of excited state contributions” can lead to a violation of this rule when the Coulomb barriers in entrance and exit channels are very different [4, 5]. Since the relevant contributions of reactions on excited states involve different relative interaction energies, depending on the excitation energies of the participating excited states and the reaction Q value, Coulomb barriers in entrance and exit channel act differently.

Ref. [5] provided an in-depth investigation of this effect across the nuclear chart and also discussed the relevance in nucleosynthesis studies. It focused, however, on a comparison of the stellar enhancement factors (SEFs) in the forward and reverse reaction direction. Later it was shown that the SEF is not always a good measure of how reactions on excited states contribute to the stellar rate [6]. In this work, a similar investigation of the suppression effect is performed but using the better suited g.s. and excited state contributions to the stellar rate to check whether the conclusions of the previous investigation still hold and to provide updated tables.

II. DEFINITIONS

The SEF f is defined as the ratio of the stellar rate r^* and reactivity R^* relative to the ground state (g.s.) rate

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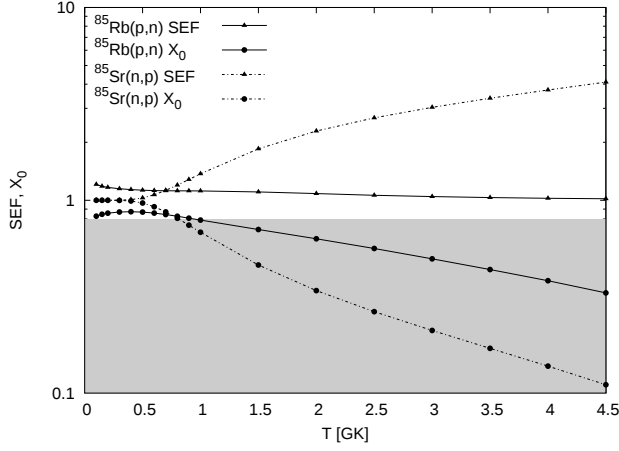


FIG. 1. Ground-state contributions X_0 and SEFs f as function of plasma temperature T for $^{85}\text{Sr}(n,p)^{85}\text{Rb}$ and its reverse reaction. The exclusion limit $X_0 < 0.8$ is marked by the shaded area.

$r^{\text{g.s.}}$ and reactivity R_0 , respectively, [3, 7]

$$f(T) = \frac{r^*(T)}{r^{\text{g.s.}}(T)} = \frac{R^*(T)}{R_0(T)} \quad (1)$$

at a plasma temperature T . While Coulomb suppression certainly drives the SEF value towards unity, a value close to 1.0 cannot be interpreted *a priori* as that excited state contributions are negligible. Later it has been shown that the more useful quantity to consider is the g.s. contribution to the stellar rate, [6, 8]

$$X_0(T) = (2J_0 + 1) \frac{R_0}{G(T)R^*} \quad , \quad (2)$$

where J_0 is the g.s. spin of the target nucleus and G is the temperature-dependent partition function, defined as

$$G(T) = \sum_i (2J_i + 1) e^{-\frac{E_i}{kT}} \quad . \quad (3)$$

The sum runs over g.s. ($i = 0$) and excited states ($i > 0$) at excitation energy E_i .

The SEF can be close to unity even when excited state contributions are sizeable whereas X_0 directly shows the impact of such contributions, being $X_0 \approx 1$ when they are negligible and $X_0 \ll 1$ when they dominate the rate [6]. Laboratory measurements only provide cross sections for nuclei in their g.s., and thus $r^{\text{g.s.}}$, except for $^{180\text{m}}\text{Ta}$ which naturally occurs in its long-lived isomer. Therefore, X_0 is the quantity of interest also to experimentalists because it shows what fraction of the stellar rate can be constrained by measurements.

III. SUPPRESSION OF X_0

Here, an investigation of the suppression effect is performed similar to the one shown in [5] but focussing

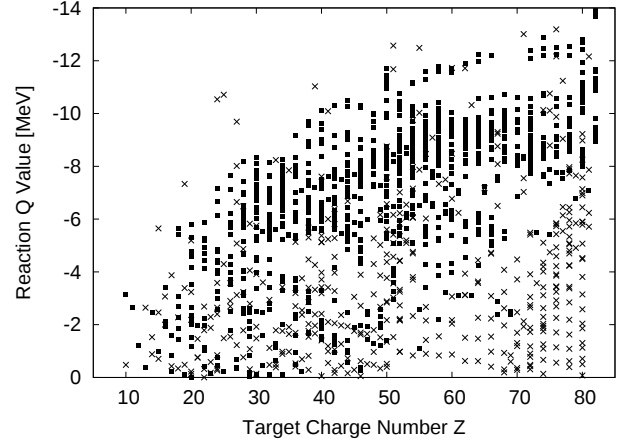


FIG. 2. Reaction Q values for (α,n) reactions (crosses) and (α,p) reactions (squares) with $X_0^{\text{rev}} > X_0^{\text{forw}}$. To be compared to Fig. 2 of [5].

on X_0 , comparing the g.s. contribution of the forward rate X_0^{forw} to that for its reverse rate X_0^{rev} and identifying cases with $X_0^{\text{rev}} > X_0^{\text{forw}}$. The calculations were performed with the Hauser-Feshbach code SMARAGD [9] using updated excited state information [10, 11] and masses [12]. The code also uses an improved barrier penetration routine which is appropriate to treat low-energy α transmissions through high Coulomb barriers [3]. As in the previous work, the forward reaction direction is defined as the one with positive reaction Q value, and reactions involving light projectiles (nucleons, α) are studied, with target nuclei ranging from Ne to Bi between the proton and neutron driplines. The selection criteria for the displayed cases had to be modified with respect to the ones used in [5] because of the different properties of X_0 . To avoid trivial cases and find sizeable differences, only reactions with $X_0^{\text{rev}}/X_0^{\text{forw}} > 1.3$ and $X_0^{\text{rev}} > 0.8$ were considered. This choice ensures that the stellar rate of the reaction with negative Q value is dominated by the g.s. contribution. It also implies that the g.s. contributions differ by at least 30% between forward and reverse direction. Only plasma temperatures $T \leq 4.5$ GK were included to identify cases important in most nucleosynthesis environments and to eliminate those only occurring at very high temperature.

Using these restrictions, 2350 reactions (out of 57513) exhibiting a strong suppression effect of X_0^{rev} remained, more than in the previous investigation. The reason for a larger number of cases being found is that a large variation in the g.s. contribution does not cause an equally big change in the SEF in many cases. On the other hand, some cases do not appear anymore because the g.s. contribution decreases although the SEF is remaining constant or becoming smaller. Figure 1 shows an example of this and also how the selection criteria work for the reaction $^{85}\text{Sr}(n,p)^{85}\text{Rb}$ and its reverse reaction. Although the SEF of $^{85}\text{Rb}(p,n)^{85}\text{Sr}$ ($Q = -1.847$ MeV) is close to unity for all temperatures, the actual g.s. contribution

TABLE I. Targets for (α, γ) reactions with negative Q value but larger X_0 than their reverse reaction. Stable or long-lived targets are in italics. Underlined targets were also found in [5].

⁸⁰ Mo	¹⁰⁶ Sb	¹¹⁶ Xe	¹²³ Ba	¹³¹ Ce	¹³⁶ Nd	¹⁴³ Sm	¹⁴⁷ Gd	¹⁴⁶ Dy	¹⁶⁹ Ho	¹⁵⁴ Yb	¹⁷⁹ Hf	¹⁸⁵ Re	¹⁹¹ Ir	¹⁹³ Au	²¹⁰ Hg
⁹⁸ Cd	¹⁰⁷ Sb	¹¹⁷ Xe	¹²⁴ Ba	¹³² Ce	¹³⁷ Nd	¹⁴⁴ Sm	¹⁴⁸ Gd	¹⁴⁸ Dy	¹⁴² Er	¹⁵⁶ Yb	¹⁸⁰ Hf	¹⁸⁷ Re	¹⁹² Ir	¹⁹⁴ Au	²¹¹ Hg
⁹⁹ Cd	¹⁰⁸ Sb	¹¹⁸ Xe	¹²⁵ Ba	¹³³ Ce	¹³⁸ Nd	¹⁴⁵ Sm	¹⁴⁹ Gd	¹⁴⁹ Dy	¹⁴⁴ Er	¹⁵⁸ Yb	¹⁸¹ Hf	¹⁸⁸ Re	¹⁹³ Ir	¹⁹⁵ Au	²¹² Hg
¹⁰⁰ Cd	¹⁰⁹ Sb	¹¹⁹ Xe	¹²⁶ Ba	¹³⁴ Ce	¹³⁹ Nd	¹⁴⁶ Sm	¹⁵⁰ Gd	¹⁵⁰ Dy	¹⁴⁶ Er	¹⁶⁰ Yb	¹⁸² Hf	¹⁸⁹ Re	¹⁹⁴ Ir	¹⁹⁷ Au	²¹³ Hg
¹⁰¹ Cd	¹¹¹ Sb	¹²⁰ Xe	¹²⁷ Ba	¹³⁵ Ce	¹⁴¹ Nd	¹⁴⁷ Sm	¹⁵¹ Gd	¹⁵¹ Dy	¹⁴⁸ Er	¹⁶² Yb	¹⁸³ Hf	¹⁹⁰ Re	¹⁹⁵ Ir	¹⁹⁸ Au	²¹⁴ Hg
⁹⁹ In	¹¹² Sb	¹²¹ Xe	¹²⁸ Ba	¹³⁹ Ce	¹⁴² Nd	¹⁴⁸ Sm	¹⁵² Gd	¹⁵² Dy	¹⁵⁰ Er	¹⁷³ Yb	¹⁸⁴ Hf	²⁰¹ Re	²⁰³ Ir	¹⁹⁹ Au	²¹⁵ Hg
¹⁰¹ In	¹¹³ Sb	¹²² Xe	¹³⁰ Ba	¹⁴⁰ Ce	¹⁴³ Nd	¹⁵⁰ Sm	¹⁵³ Gd	¹⁵⁴ Dy	¹⁵¹ Er	¹⁷⁴ Yb	¹⁵⁵ Ta	¹⁶⁷ Os	²⁰⁵ Ir	²⁰¹ Au	²¹⁶ Hg
¹⁰² In	¹¹⁴ Sb	¹²³ Xe	¹³⁸ Ba	¹⁴¹ Ce	¹⁴⁴ Nd	¹³⁷ Eu	¹⁵⁴ Gd	¹⁵⁶ Dy	¹⁵² Er	¹⁷⁵ Yb	¹⁶⁰ Ta	¹⁶⁸ Os	²⁰⁷ Ir	²⁰⁵ Au	²¹⁸ Hg
¹⁰³ In	¹¹⁵ Sb	¹³⁶ Xe	¹³⁹ Ba	¹⁴² Ce	¹⁴⁵ Nd	¹³⁹ Eu	¹⁵⁵ Gd	¹⁵⁷ Dy	¹⁵³ Er	¹⁷⁶ Yb	¹⁷⁹ Ta	¹⁶⁹ Os	¹⁶⁸ Pt	²⁰⁷ Au	²²⁰ Hg
¹⁰⁴ In	¹¹⁶ Sb	¹³⁷ Xe	¹⁴⁰ Ba	¹⁴³ Ce	¹⁴⁶ Nd	¹⁴¹ Eu	¹⁵⁶ Gd	¹⁵⁸ Dy	¹⁵⁴ Er	¹⁷⁷ Yb	¹⁸³ Ta	¹⁸⁸ Os	¹⁷⁰ Pt	²⁰⁹ Au	¹⁷⁷ Tl
¹⁰⁵ In	¹³³ Sb	¹¹¹ Cs	¹⁴¹ Ba	¹⁴⁴ Ce	¹⁴⁷ Nd	¹⁴² Eu	¹⁵⁷ Gd	¹⁵⁹ Dy	¹⁵⁶ Er	¹⁷⁸ Yb	¹⁸⁴ Ta	¹⁸⁹ Os	¹⁸⁸ Pt	²¹³ Au	¹⁷⁹ Tl
¹⁰⁶ In	¹⁰³ Te	¹¹⁷ Cs	¹²⁷ La	¹²⁹ Pr	¹⁴⁸ Nd	¹⁴³ Eu	¹⁵⁸ Gd	¹⁶⁰ Dy	¹⁵⁷ Er	¹⁸⁰ Yb	¹⁸⁵ Ta	¹⁹⁰ Os	¹⁹⁰ Pt	²¹⁵ Au	¹⁸¹ Tl
¹⁰⁷ In	¹⁰⁴ Te	¹¹⁹ Cs	¹²⁸ La	¹³¹ Pr	¹³¹ Pr	¹⁴⁴ Eu	¹⁴² Tb	¹⁶¹ Dy	¹⁵⁸ Er	¹⁸² Yb	¹⁵⁷ W	¹⁹¹ Os	¹⁹¹ Pt	¹⁷⁴ Hg	¹⁹⁰ Tl
¹⁰⁸ In	¹⁰⁶ Te	¹²⁰ Cs	¹²⁹ La	¹³² Pr	¹³³ Pm	¹⁴⁵ Eu	¹⁴⁴ Tb	¹⁶² Dy	¹⁶⁴ Er	¹⁸⁴ Yb	¹⁵⁸ W	¹⁹² Os	¹⁹² Pt	¹⁷⁷ Hg	¹⁹⁵ Tl
⁹⁹ Sn	¹⁰⁸ Te	¹²¹ Cs	¹³⁰ La	¹³³ Pr	¹³⁷ Pm	¹⁴⁷ Eu	¹⁴⁶ Tb	¹⁶³ Dy	¹⁶⁵ Er	¹⁵¹ Lu	¹⁵⁹ W	¹⁹³ Os	¹⁹⁴ Pt	¹⁷⁸ Hg	¹⁹⁶ Tl
¹⁰² Sn	¹¹⁰ Te	¹²² Cs	¹³¹ La	¹³⁵ Pr	¹³⁹ Pm	¹⁴⁸ Eu	¹⁵⁰ Tb	¹⁶⁴ Dy	¹⁶⁶ Er	¹⁵³ Lu	¹⁶⁰ W	¹⁹⁴ Os	¹⁹⁵ Pt	¹⁸⁸ Hg	¹⁹⁷ Tl
¹⁰⁴ Sn	¹¹¹ Te	¹²³ Cs	¹³⁸ La	¹³⁶ Pr	¹⁴⁰ Pm	¹⁴⁹ Eu	¹⁵² Tb	¹⁶⁵ Dy	¹⁶⁷ Er	¹⁵⁵ Lu	¹⁶³ W	²⁰¹ Os	¹⁹⁶ Pt	¹⁹⁰ Hg	¹⁹⁸ Tl
¹⁰⁵ Sn	¹¹² Te	¹²⁴ Cs	¹³⁹ La	¹³⁷ Pr	¹⁴¹ Pm	¹⁵⁰ Eu	¹⁵³ Tb	¹⁶⁶ Dy	¹⁶⁸ Er	¹⁶⁷ Lu	¹⁶⁵ W	²⁰² Os	¹⁹⁷ Pt	¹⁹² Hg	¹⁹⁹ Tl
¹⁰⁶ Sn	¹¹⁴ Te	¹²⁵ Cs	¹⁴⁰ La	¹⁴⁰ Pr	¹⁴² Pm	¹⁵¹ Eu	¹⁵⁵ Tb	¹⁴⁵ Ho	¹⁶⁹ Er	¹⁷⁵ Lu	¹⁸¹ W	²⁰³ Os	¹⁹⁸ Pt	¹⁹⁴ Hg	²⁰⁰ Tl
¹⁰⁷ Sn	¹¹⁵ Te	¹²⁶ Cs	¹⁴¹ La	¹⁴¹ Pr	¹⁴³ Pm	¹⁵² Eu	¹⁵⁶ Tb	¹⁴⁷ Ho	¹⁷⁰ Er	¹⁷⁶ Lu	¹⁸⁴ W	²⁰⁴ Os	²⁰³ Pt	¹⁹⁶ Hg	²⁰¹ Tl
¹⁰⁸ Sn	¹¹⁶ Te	¹²⁷ Cs	¹⁴² La	¹⁴³ Pr	¹⁴⁷ Pm	¹⁵³ Eu	¹⁵⁷ Tb	¹⁵⁰ Ho	¹⁷¹ Er	¹⁷⁷ Lu	¹⁸⁵ W	²⁰⁵ Os	²⁰⁴ Pt	¹⁹⁸ Hg	²⁰³ Tl
¹¹⁰ Sn	¹¹⁷ Te	¹³⁷ Cs	¹⁴³ La	¹⁴⁵ Pr	¹⁴⁸ Pm	¹³² Gd	¹⁵⁸ Tb	¹⁵¹ Ho	¹⁷² Er	¹⁷⁸ Lu	¹⁸⁶ W	²⁰⁶ Os	²⁰⁵ Pt	¹⁹⁹ Hg	²⁰⁴ Tl
¹¹¹ Sn	¹¹⁷ I	¹³⁸ Cs	¹¹⁶ Ce	¹⁴⁶ Pr	¹⁴⁹ Pm	¹³⁶ Gd	¹⁵⁹ Tb	¹⁵³ Ho	¹⁴⁵ Tm	¹⁷⁹ Lu	¹⁸⁷ W	²⁰⁷ Os	²⁰⁶ Pt	²⁰⁰ Hg	²⁰⁵ Tl
¹¹² Sn	¹¹⁹ I	¹³⁹ Cs	¹¹⁸ Ce	¹²² Nd	¹²⁶ Sm	¹³⁸ Gd	¹⁶⁰ Tb	¹⁵⁹ Ho	¹⁴⁶ Tm	¹⁵² Hf	¹⁸⁸ W	²⁰⁸ Os	²⁰⁷ Pt	²⁰² Hg	²¹³ Tl
¹¹³ Sn	¹³⁵ I	¹¹² Ba	¹²⁰ Ce	¹²⁴ Nd	¹²⁸ Sm	¹³⁹ Gd	¹⁶¹ Tb	¹⁶¹ Ho	¹⁴⁷ Tm	¹⁵⁴ Hf	²⁰⁰ W	¹⁶⁸ Ir	²⁰⁸ Pt	²⁰³ Hg	²¹⁵ Tl
¹¹⁴ Sn	¹⁰⁸ Xe	¹¹⁴ Ba	¹²⁴ Ce	¹²⁶ Nd	¹³⁶ Sm	¹⁴⁰ Gd	¹⁶² Tb	¹⁶² Ho	¹⁵² Tm	¹⁵⁵ Hf	²⁰¹ W	¹⁷¹ Ir	²⁰⁹ Pt	²⁰⁴ Hg	²¹⁹ Tl
¹¹⁵ Sn	¹¹⁰ Xe	¹¹⁶ Ba	¹²⁵ Ce	¹³¹ Nd	¹³⁷ Sm	¹⁴¹ Gd	¹³⁶ Dy	¹⁶³ Ho	¹⁷⁰ Tm	¹⁵⁶ Hf	²⁰² W	¹⁷² Ir	²¹⁰ Pt	²⁰⁵ Hg	
¹⁰¹ Sb	¹¹² Xe	¹¹⁸ Ba	¹²⁶ Ce	¹³² Nd	¹³⁸ Sm	¹⁴² Gd	¹³⁸ Dy	¹⁶⁴ Ho	¹⁷² Tm	¹⁵⁷ Hf	²⁰³ W	¹⁷³ Ir	²¹¹ Pt	²⁰⁶ Hg	
¹⁰² Sb	¹¹³ Xe	¹²⁰ Ba	¹²⁸ Ce	¹³³ Nd	¹³⁹ Sm	¹⁴³ Gd	¹⁴⁰ Dy	¹⁶⁵ Ho	¹⁴⁸ Yb	¹⁵⁸ Hf	¹⁶⁴ Re	¹⁷⁵ Ir	²¹² Pt	²⁰⁷ Hg	
¹⁰³ Sb	¹¹⁴ Xe	¹²¹ Ba	¹²⁹ Ce	¹³⁴ Nd	¹⁴⁰ Sm	¹⁴⁴ Gd	¹⁴² Dy	¹⁶⁷ Ho	¹⁵² Yb	¹⁷⁷ Hf	¹⁶⁷ Re	¹⁷⁷ Ir	²¹⁴ Pt	²⁰⁸ Hg	
¹⁰⁵ Sb	¹¹⁵ Xe	¹²² Ba	¹³⁰ Ce	¹³⁵ Nd	¹⁴² Sm	¹⁴⁶ Gd	¹⁴⁴ Dy	¹⁶⁸ Ho	¹⁵³ Yb	¹⁷⁸ Hf	¹⁸² Re	¹⁸⁹ Ir	¹⁷² Au	²⁰⁹ Hg	

TABLE II. Same as Table I but for (p, γ) .

³³ Ar	⁶² Ge	⁹⁴ Cd	¹¹³ Xe	¹³⁸ Gd	¹⁵² Hf	¹⁶⁸ Os	¹⁷⁹ Pb
³⁷ Ca	⁶³ Ge	⁹⁵ Cd	¹¹⁴ Xe	¹³⁸ Dy	¹⁵⁴ Hf	¹⁷⁰ Os	¹⁸⁰ Pb
⁴⁰ Ti	⁶⁴ Ge	⁹⁶ Cd	¹¹¹ Cs	¹⁴⁰ Dy	¹⁵⁵ Hf	¹⁶⁸ Pt	¹⁸¹ Pb
⁴¹ Ti	⁶⁸ Se	⁹⁴ In	¹¹² Cs	¹⁴² Dy	¹⁵⁶ Hf	¹⁷⁰ Pt	¹⁸² Pb
⁴⁵ Fe	⁷² Kr	⁹⁸ Sn	¹¹² Ba	¹⁴⁴ Dy	¹⁵⁷ Hf	¹⁷² Pt	¹⁸³ Pb
⁴⁸ Fe	⁷⁴ Sr	¹⁰⁰ Sn	¹¹⁶ Ba	¹⁴² Er	¹⁵⁸ Hf	¹⁷⁴ Pt	¹⁸⁴ Pb
⁴⁹ Fe	⁷⁵ Sr	¹⁰¹ Sn	¹¹⁸ Ba	¹⁴⁶ Er	¹⁵⁶ W	¹⁷⁶ Pt	¹⁸⁵ Pb
⁵² Ni	⁷⁶ Sr	¹⁰² Sn	¹¹⁸ Ce	¹⁴⁸ Er	¹⁵⁷ W	¹⁷³ Hg	¹⁸⁶ Pb
⁵³ Ni	⁷³ Y	¹⁰³ Sn	¹²² Ce	¹⁴⁵ Tm	¹⁵⁸ W	¹⁷⁴ Hg	¹⁸⁷ Pb
⁵⁴ Ni	⁸¹ Nb	¹⁰⁴ Te	¹²⁴ Nd	¹⁴⁶ Tm	¹⁵⁹ W	¹⁷⁶ Hg	¹⁸⁸ Pb
⁵⁴ Zn	⁸⁴ Mo	¹⁰⁶ Te	¹²⁶ Nd	¹⁴⁸ Yb	¹⁶⁰ W	¹⁷⁷ Hg	
⁵⁶ Zn	⁸³ Tc	¹⁰⁷ Te	¹²³ Pm	¹⁵⁰ Yb	¹⁶² W	¹⁷⁸ Hg	
⁵⁷ Zn	⁸⁴ Tc	¹⁰⁸ Te	¹³² Sm	¹⁵¹ Yb	¹⁶⁴ W	¹⁸⁰ Hg	
⁵⁸ Zn	⁸⁶ Ru	¹⁰⁹ Te	¹³⁴ Sm	¹⁵² Yb	¹⁶² Os	¹⁸¹ Hg	
⁶⁰ Ge	⁹⁰ Pd	¹⁰⁶ I	¹³⁴ Gd	¹⁵³ Yb	¹⁶⁴ Os	¹⁷⁵ Tl	
⁶¹ Ge	⁹² Pd	¹¹² Xe	¹³⁶ Gd	¹⁵⁴ Yb	¹⁶⁶ Os	¹⁷⁸ Pb	

decreases with increasing temperature. Moreover, comparing the evolution of X_0 as function of temperature for forward and reverse reaction, it becomes apparent that X_0^{rev} is larger than $1.3X_0^{\text{forw}}$ only when both are smaller than 0.8. At lower T , both g.s. contributions are above the cutoff at 0.8 but there $X_0^{\text{forw}} > X_0^{\text{rev}}$. Therefore, the reaction $^{85}\text{Rb}(p, n)^{85}\text{Sr}$ does not show up anymore in the tables given below.

Figure 2 can be directly compared to Fig. 2 of [5], illustrating the dependence of the suppression of the Coulomb barrier. Shown is the obtained range of Q val-

TABLE III. Same as Table I but for (γ, n) .

²⁹ Ne	⁴⁴ Mg	⁶¹ Ca	¹⁰⁵ Se	¹²² Sr	¹⁶⁹ Te	²³⁰ Yb	²⁴³ Au
³¹ Ne	²² Al	⁷² Ti	¹¹⁵ Se	¹²⁴ Zr	¹⁷³ Ba	²³⁶ Hf	²⁴³ Hg
²⁸ Na	⁴⁰ Al	⁷⁷ V	¹¹⁶ Se	¹²⁸ Mo	¹⁸⁶ Ce	²³⁹ Os	²⁴⁵ Hg
³⁵ Mg	⁴⁵ Al	⁸¹ Mn	¹⁰⁹ Kr	¹⁴⁸ Ru	¹⁸⁸ Ce	²⁴⁴ Os	²⁶⁰ Hg
³⁷ Mg	⁴⁸ Si	¹⁰¹ Ge	¹¹⁸ Kr	¹⁴⁹ Rh	¹⁸⁸ Nd	²⁵¹ Pt	²⁶⁴ Hg
⁴² Mg	⁵⁵ Ar	¹⁰⁸ Ge	¹²¹ Rb	¹⁵¹ Rh	¹⁹⁰ Sm	²⁵² Pt	

ues for (α, n) and (α, p) reactions with negative Q values fulfilling the above criteria. As explained in [5], larger maximal $|Q|$ is allowed with increasing charge Z . This documents that the explanation given in the earlier paper [5] can also be applied when using g.s. contributions instead of SEF.

IV. RESULTS AND DISCUSSION

Tables I–IX list the reactions found to have $X_0^{\text{rev}} > X_0^{\text{forw}}$ according to the criteria discussed above. Shown are the target nuclides for the endothermic direction. Stable or long-lived target nuclei are printed in italics. The present Tables I–VIII are in the same sequence as and formatted similarly to tables I–VIII in [5] and supersede those. Underlined target nuclei appeared already in the previous tables. Most nuclei are new entries, they are not underlined. Target nuclei which do not appear

TABLE IV. Same as Table I but for (γ, p) , (n, α) , and (n, γ) reactions.

(γ, p) :						
^{38}Ti	^{46}Fe	^{53}Co	^{64}Se	^{99}Sn	^{137}Eu	^{165}Ta
^{46}Cr	^{52}Co	^{55}Zn	^{79}Sr	^{131}Pm	^{150}Tb	^{169}Ta
(n, α) :						
^{65}Cu						
(n, γ) :						
^{38}Mg	^{79}Cr	^{130}Ru	^{133}Ag	^{191}Ce	^{194}Gd	^{246}Pt
^{44}Mg	^{80}Fe	^{132}Ru	^{136}Cd	^{192}Sm	^{196}Gd	^{243}Hg
^{40}Si	^{82}Ni	^{132}Pd	^{158}Sn	^{189}Eu	^{222}Yb	^{263}Tl
^{74}Cr	^{91}Cu	^{144}Pd	^{173}I	^{192}Gd	^{231}Lu	

TABLE V. Same as Table I but for (p, n) .

^{32}Si	^{77}As	^{104}Ru	^{126}Te	^{140}Ce	^{163}Dy	^{180}Ta	^{202}Hg
^{36}Cl	^{79}Se	^{106}Ru	^{127}Te	^{141}Ce	^{166}Dy	^{174}W	^{204}Hg
^{40}Ar	^{82}Se	^{105}Rh	^{128}Te	^{142}Ce	^{162}Ho	^{188}W	^{201}Tl
^{42}Ar	^{81}Br	^{107}Pd	^{130}Te	^{144}Ce	^{165}Ho	^{182}Re	^{202}Tl
^{46}Ca	^{85}Kr	^{110}Pd	^{132}Te	^{143}Nd	^{167}Er	^{184}Re	^{203}Tl
^{48}Ca	^{86}Kr	^{112}Pd	^{124}I	^{144}Nd	^{169}Er	^{185}Re	^{204}Tl
^{47}Sc	^{87}Rb	^{113}Cd	^{131}Xe	^{145}Nd	^{164}Yb	^{187}Re	^{205}Tl
^{49}Ti	^{87}Sr	^{116}Cd	^{133}Xe	^{146}Nd	^{175}Yb	^{191}Os	^{200}Pb
^{53}Cr	^{88}Sr	^{118}Cd	^{134}Xe	^{148}Nd	^{176}Yb	^{194}Os	^{206}Pb
^{60}Fe	^{90}Sr	^{111}In	^{136}Xe	^{150}Nd	^{178}Yb	^{186}Ir	^{207}Pb
^{59}Co	^{91}Zr	^{113}In	^{128}Cs	^{147}Pm	^{163}Lu	^{188}Ir	^{208}Pb
^{64}Ni	^{92}Zr	^{115}In	^{130}Cs	^{145}Sm	^{174}Lu	^{193}Ir	^{209}Pb
^{66}Ni	^{93}Zr	^{118}Sn	^{135}Cs	^{156}Sm	^{175}Lu	^{195}Pt	^{210}Pb
^{65}Cu	^{94}Zr	^{120}Sn	^{135}Ba	^{148}Eu	^{177}Lu	^{196}Pt	^{212}Pb
^{67}Cu	^{96}Zr	^{122}Sn	^{136}Ba	^{155}Eu	^{160}Hf	^{197}Pt	
^{70}Zn	^{97}Mo	^{124}Sn	^{137}Ba	^{150}Tb	^{179}Hf	^{198}Pt	
^{72}Zn	^{98}Mo	^{126}Sn	^{138}Ba	^{152}Tb	^{182}Hf	^{200}Pt	
^{71}Ga	^{100}Mo	^{123}Sb	^{139}La	^{156}Tb	^{167}Ta	^{197}Au	
^{76}Ge	^{99}Tc	^{125}Sb	^{139}Ce	^{161}Tb	^{174}Ta	^{199}Au	

anymore can be found by direct comparison with the previous tables. Notable are the many more cases found for endothermic (α, p) reactions (Table VIII) compared to the previous investigation. Table IX is new because endothermic (p, α) reactions did not show any SEF suppression previously but a suppression of the g.s. contribution was found here. A further difference to the previous investigation are (γ, α) reactions on the following target nuclei: ^{27}S , ^{29}Ar , ^{38}Sc , ^{48}Mn , ^{83}Tc , ^{29}S , ^{36}K , ^{40}Sc , ^{57}Ge , ^{97}Sn . No SEF suppression was found for endothermic (γ, α) but a suppression of X_0 was found for the above few cases.

The majority of the shown cases are caused by Coulomb suppression of the excited state contributions in one reaction channel. As in [5], however, a suppression is also found for a few (n, γ) , (γ, n) , (γ, p) , (γ, α) , (n, p) , (n, α) , and (p, α) cases (Tables III, IV, VI, and the (γ, α) reactions mentioned above). Obviously, this suppression is not caused by the Coulomb barrier but rather transitions from excited states are suppressed by selection rules and centrifugal barriers. It has to be noted that the results for most of these nuclei – as well as others

far off stability – are purely based on theory, as neither spectroscopic information nor masses are available from experiment. In these cases, recommended mass values were taken [12, 13] or a theoretical mass model [14] was used to compute the separation energies. In the absence of knowledge about excited states, a theoretical nuclear level density was employed [15, 16]. Therefore new mass measurements could change the reaction Q values and also further experimental information on excited states may impact the results presented here, especially close to the driplines.

As already discussed in [5], to current knowledge most of the shown reactions are not of direct astrophysical importance. They are merely interesting cases to demonstrate the suppression effect and to study where the standard Q value rule can be applied and where this is not possible. Future developments in astrophysical models and experimental techniques should not be precluded, however, and therefore the full tables are given. For further details on the astrophysical relevance and also on the limitations of the applied reaction model, see [5].

As has been pointed out previously, most remarkable in the current astrophysics context are the (α, γ) reactions on proton-rich targets listed in Table I, especially the ones on stable targets with neutron number $N \geq 82$. These are important to study (γ, α) rates and the optical α -nucleus potential required for predictions relevant to the synthesis of p nuclei [17]. Also important in the astrophysical γ - and in the νp process are (p, n) reactions with negative Q value (Table V) because (n, p) reactions on proton-rich isotopes play a role in these nucleosynthesis processes (see, e.g., [4, 5, 17–21]). Endothermic (α, p) reactions (Table VIII) on proton-rich nuclei are also of interest for νp process studies [22]. Further of interest are the (p, γ) reactions along the proton dripline shown in Table II because they appear in the rp-process [23]. Moreover, the large number of (α, γ) and (p, γ) reactions (and the comparatively small number of (γ, p) and (γ, α) in Table IV and given in the text above) found in this study underlines the fact that it is almost always preferable to measure in the capture direction, even for endothermic captures. This is consistent with the findings of [3, 24].

V. SUMMARY

A previous investigation was improved by studying the g.s. contributions in forward and reverse reaction directions instead of comparing the SEFs. Again, it was found that – contrary to common wisdom – some endothermic reactions exhibit smaller g.s. contributions to the stellar reaction rate than their exothermic counterparts and are thus preferable for experimental and theoretical studies, if one of the reaction directions is found to be important in a hot plasma. The main cause of suppression of the excited state contributions in an endothermic reaction is the Coulomb suppression of transitions with low relative

TABLE VI. Same as Table I but for (n, p) reactions.

^{32}S	^{40}Ca
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TABLE VII. Same as Table I but for (α, n) .

²² Ne	⁵⁰ Ti	⁶² Ni	⁸⁴ Se	⁹³ Y	⁹⁷ Mo	¹¹⁶ Ag	¹¹³ Sb	¹³⁷ I	¹⁴⁷ Pr	¹⁷⁴ Dy	¹⁹⁶ Yb	¹⁹⁴ W	²⁰⁰ Os	²⁰⁶ Pt	²¹² Hg
²⁷ Al	⁵² Ti	⁶⁴ Ni	⁷⁵ Br	⁹⁴ Y	⁹⁸ Mo	¹¹⁴ Cd	¹¹⁵ Sb	¹²⁶ Xe	¹³² Nd	¹⁷⁶ Dy	¹⁹⁸ Yb	¹⁹⁶ W	²⁰¹ Os	²⁰⁷ Pt	²¹⁴ Hg
²⁹ Si	⁴⁸ V	⁶⁵ Ni	⁸² Br	⁹⁵ Y	¹⁰⁰ Mo	¹¹⁶ Cd	¹²⁵ Sb	¹³⁰ Xe	¹³⁸ Nd	¹⁷⁸ Dy	²⁰⁰ Yb	¹⁹⁷ W	²⁰² Os	²⁰⁸ Pt	²¹⁵ Hg
³² Si	⁵⁰ V	⁶⁶ Ni	⁸³ Br	⁹⁶ Y	¹⁰² Mo	¹¹⁷ Cd	¹²⁷ Sb	¹³⁴ Xe	¹⁴⁸ Nd	¹⁸⁰ Dy	¹⁶⁷ Lu	¹⁹⁸ W	²⁰³ Os	²⁰⁹ Pt	²¹⁶ Hg
³¹ P	⁵¹ V	⁶⁸ Ni	⁸⁵ Br	⁹⁷ Y	¹⁰³ Mo	¹¹⁸ Cd	¹²⁹ Sb	¹³⁵ Xe	¹⁵⁰ Nd	¹⁸² Dy	¹⁷⁸ Lu	¹⁹⁹ W	²⁰⁴ Os	²¹⁰ Pt	²¹⁸ Hg
³³ P	⁵³ V	⁶⁵ Cu	⁸² Kr	⁸⁹ Zr	¹⁰⁴ Mo	¹²⁰ Cd	¹³¹ Sb	¹³⁶ Xe	¹⁵⁸ Nd	¹⁸⁴ Dy	¹⁸⁵ Lu	²⁰⁰ W	²⁰⁵ Os	²¹¹ Pt	²²⁰ Hg
³⁵ P	⁴⁶ Cr	⁶⁷ Cu	⁸⁴ Kr	⁹⁰ Zr	¹⁰⁶ Mo	¹²² Cd	¹³² Sb	¹³⁷ Xe	¹⁶⁰ Nd	¹⁸⁰ Ho	¹⁹⁷ Lu	²⁰¹ W	²⁰⁶ Os	²¹² Pt	²²² Hg
³³ S	⁴⁹ Cr	⁶⁸ Cu	⁸⁶ Kr	⁹¹ Zr	¹⁰⁸ Mo	¹¹⁷ In	¹³³ Sb	¹³⁸ Xe	¹⁶² Nd	¹⁶² Ho	¹⁹⁹ Lu	²⁰² W	²⁰⁷ Os	²¹⁴ Pt	²²⁴ Hg
³⁶ S	⁵¹ Cr	⁶⁹ Cu	⁸⁷ Kr	⁹³ Zr	¹⁰² Tc	¹¹⁹ In	¹³⁵ Sb	¹⁴⁰ Xe	¹⁶⁴ Nd	¹⁶⁴ Ho	¹⁸⁸ Hf	²⁰³ W	²⁰⁸ Os	²¹⁶ Pt	²²⁶ Hg
³⁷ Cl	⁵² Cr	⁷¹ Cu	⁸⁸ Kr	⁹⁴ Zr	¹⁰⁴ Tc	¹²¹ In	¹²⁶ Te	¹⁴² Xe	¹³¹ Pm	¹⁶⁵ Ho	¹⁹⁰ Hf	²⁰⁴ W	²¹⁰ Os	²¹⁸ Pt	²²⁸ Hg
³⁹ Cl	⁵³ Cr	⁶⁸ Zn	⁹⁰ Kr	⁹⁵ Zr	¹⁰⁷ Tc	¹²³ In	¹²⁸ Te	¹⁴⁴ Xe	¹⁵³ Pm	¹⁶⁸ Ho	¹⁹² Hf	²⁰⁵ W	²¹² Os	²²⁰ Pt	²³⁰ Hg
³⁹ Ar	⁵⁴ Cr	⁶⁹ Zn	⁸⁴ Rb	⁹⁶ Zr	¹⁰⁹ Tc	¹²⁵ In	¹³⁰ Te	¹⁴⁶ Xe	¹⁴² Sm	¹⁸⁰ Er	¹⁹⁴ Hf	²⁰⁶ W	²¹⁴ Os	²²² Pt	²³² Hg
⁴⁰ Ar	⁵⁶ Cr	⁷⁰ Zn	⁸⁶ Rb	⁹⁷ Zr	⁹³ Ru	¹²⁷ In	¹³¹ Te	¹²¹ Cs	¹⁶² Sm	¹⁸² Er	¹⁹⁵ Hf	²⁰⁸ W	²¹⁶ Os	²²⁴ Pt	²³⁴ Hg
³⁹ K	⁴⁷ Mn	⁷² Zn	⁸⁷ Rb	⁹⁸ Zr	¹⁰⁹ Ru	¹¹⁸ Sn	¹³² Te	¹²² Cs	¹⁶⁴ Sm	¹⁸⁴ Er	¹⁹⁶ Hf	²¹⁰ W	²¹⁸ Os	²²⁶ Pt	²³⁶ Hg
⁴¹ K	⁵³ Mn	⁷⁴ Zn	⁸⁹ Rb	¹⁰⁰ Zr	¹¹⁰ Ru	¹²⁰ Sn	¹³³ Te	¹³⁷ Cs	¹⁶⁶ Sm	¹⁸⁶ Er	¹⁹⁷ Hf	²¹² W	²²⁰ Os	²²⁸ Pt	²³⁸ Hg
⁴³ K	⁵⁷ Mn	⁷⁰ Ga	⁸⁶ Sr	¹⁰² Zr	¹¹² Ru	¹²² Sn	¹³⁴ Te	¹³⁸ Ba	¹⁶⁸ Sm	¹⁸⁸ Er	¹⁹⁸ Hf	²¹⁴ W	¹⁸⁶ Ir	²³⁰ Pt	²⁴⁰ Hg
⁴¹ Ca	⁵³ Fe	⁷⁵ Ge	⁸⁷ Sr	⁸⁵ Nb	¹¹¹ Rh	¹²⁴ Sn	¹³⁵ Te	¹³⁹ Ba	¹⁷⁰ Sm	¹⁹⁴ Er	¹⁹⁹ Hf	¹⁸² Re	¹⁹³ Ir	²⁰⁵ Au	²⁴² Hg
⁴² Ca	⁵⁵ Fe	⁷⁶ Ge	⁸⁸ Sr	⁹⁵ Nb	¹¹³ Rh	¹²⁶ Sn	¹³⁶ Te	¹⁴⁵ Ba	¹⁴⁸ Eu	¹⁹⁶ Er	²⁰⁰ Hf	¹⁸⁷ Re	²⁰³ Ir	²¹³ Au	¹⁹⁵ Tl
⁴⁴ Ca	⁵⁸ Fe	⁷⁸ Ge	⁸⁹ Sr	⁹⁷ Nb	¹¹⁵ Rh	¹²⁸ Sn	¹³⁷ Te	¹⁴⁷ Ba	¹⁶⁶ Gd	¹⁹⁵ Tm	²⁰¹ Hf	¹⁸⁹ Re	²⁰⁵ Ir	²¹⁵ Au	²⁰¹ Tl
⁴⁶ Ca	⁶⁰ Fe	⁷¹ As	⁹⁰ Sr	⁹⁹ Nb	¹¹¹ Pd	¹³⁰ Sn	¹³⁸ Te	¹⁴⁸ Ba	¹⁶⁸ Gd	¹⁹⁷ Tm	²⁰² Hf	²⁰¹ Re	²⁰⁷ Ir	²⁰² Hg	²¹⁷ Tl
⁴⁸ Ca	⁵³ Co	⁷⁸ As	⁹¹ Sr	¹⁰⁰ Nb	¹¹² Pd	¹³¹ Sn	¹³⁹ Te	¹⁵⁰ Ba	¹⁷⁰ Gd	¹⁶⁰ Yb	²⁰⁴ Hf	¹⁸⁴ Os	¹⁹⁸ Pt	²⁰⁶ Hg	²¹⁸ Tl
⁴⁴ Sc	⁵⁵ Co	⁷⁹ As	⁹² Sr	¹⁰¹ Nb	¹¹³ Pd	¹³² Sn	¹⁴⁰ Te	¹⁵⁹ La	¹⁷² Gd	¹⁸⁶ Yb	²⁰⁶ Hf	¹⁸⁶ Os	²⁰¹ Pt	²⁰⁷ Hg	²²¹ Tl
⁴⁷ Sc	⁵⁷ Co	⁸¹ As	⁹³ Sr	¹⁰³ Nb	¹¹⁴ Pd	¹³³ Sn	¹⁴² Te	¹²⁴ Ce	¹⁷⁴ Gd	¹⁸⁸ Yb	¹⁶⁵ Ta	¹⁹⁰ Os	²⁰² Pt	²⁰⁸ Hg	
⁴⁹ Sc	⁵⁹ Co	⁷⁹ Se	⁹⁴ Sr	¹⁰⁵ Nb	¹¹⁶ Pd	¹³⁴ Sn	¹²⁶ I	¹⁴¹ Ce	¹⁷⁶ Gd	¹⁹⁰ Yb	¹⁹⁹ Ta	¹⁹⁵ Os	²⁰³ Pt	²⁰⁹ Hg	
⁴⁷ Ti	⁶¹ Co	⁸⁰ Se	⁹⁶ Sr	⁹¹ Mo	¹¹⁸ Pd	¹³⁵ Sn	¹³³ I	¹⁵⁴ Ce	¹⁶³ Tb	¹⁹² Yb	¹⁶⁶ W	¹⁹⁸ Os	²⁰⁴ Pt	²¹⁰ Hg	
⁴⁸ Ti	⁶³ Co	⁸² Se	⁷⁹ Y	⁹³ Mo	¹¹⁴ Ag	¹³⁶ Sn	¹³⁵ I	¹⁵⁶ Ce	¹⁷² Dy	¹⁹⁴ Yb	¹⁸⁴ W	¹⁹⁹ Os	²⁰⁵ Pt	²¹¹ Hg	

interaction energy [5]. The larger number of cases is due to the different temperature dependence of the g.s. contributions compared to the SEFs.

For experimental application, it is advisable to first search for a reaction with astrophysical importance and then to check in the above tables whether it belongs to the exceptions for which the experimentally accessible g.s. contribution to the stellar rate is larger in the endothermic direction. Otherwise, a determination of the cross sections of the exothermic reaction would be preferable for astrophysical application, if feasible.

Suppression of thermally excited state contributions is

not limited to reactions treated in the Hauser-Feshbach model but also appears in resonant reactions and in direct reactions and similar considerations apply as those discussed above.

ACKNOWLEDGMENTS

This work was supported by the Hungarian Academy of Sciences, the ESF EUROCORES programme EuroGENESIS, the THEXO collaboration within the 7th Framework Programme of the EU, the European Research Council, and the Swiss NSF.

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TABLE VIII. Same as Table I but for (α ,p).

²³ Ne	⁵⁵ Fe	⁸⁹ Zn	¹⁰² Se	¹¹³ Sr	¹¹¹ Mo	¹³⁰ Rh	¹⁵³ Cd	¹⁰⁴ <u>Te</u>	¹⁵¹ Xe	¹⁷⁷ Ba	¹⁵⁴ Pr	¹⁸⁸ Sm	²⁰⁰ Dy	¹⁶⁵ Hf	¹⁹⁹ Pt
²⁸ Na	⁶² Fe	⁹¹ Zn	¹⁰³ Se	¹¹⁴ Sr	¹¹² Mo	¹³² Rh	¹⁵⁴ Cd	¹⁰⁶ <u>Te</u>	¹⁵² Xe	¹⁷⁹ Ba	¹⁵⁶ Pr	²⁰⁰ Sm	²¹⁵ Dy	¹⁶⁷ Hf	²⁰¹ Pt
²⁵ Mg	⁶³ Fe	⁹² Zn	¹⁰⁴ Se	¹¹⁵ Sr	¹¹³ Mo	¹⁴⁰ Rh	¹⁵⁵ Cd	¹⁰⁷ <u>Te</u>	¹⁵⁴ Xe	¹⁸¹ Ba	¹⁹⁴ Pr	²⁰¹ Sm	²¹⁶ Dy	¹⁷⁹ Hf	²⁰³ Pt
³⁷ Al	⁶⁹ Fe	⁹³ Zn	¹⁰⁵ Se	¹¹⁶ Sr	¹¹⁴ Mo	¹⁴⁶ Rh	¹⁵⁶ Cd	¹⁰⁹ Te	¹⁵⁵ Xe	¹⁸² Ba	¹²² Nd	²⁰⁷ Sm	²¹⁷ Dy	¹⁸³ Hf	²⁰⁵ Pt
²⁹ Si	⁷⁰ Fe	⁹⁴ Zn	¹⁰⁷ Se	¹²⁰ Sr	¹¹⁵ Mo	¹⁴⁸ Rh	⁹⁹ In	¹¹³ Te	¹⁵⁶ Xe	¹⁸⁴ Ba	¹²⁴ <u>Nd</u>	²⁰⁸ Sm	²²¹ Dy	¹⁸⁵ Hf	²⁰⁸ Pt
³⁷ P	⁷¹ Fe	⁹⁵ Zn	¹⁰⁸ Se	¹²¹ Sr	¹¹⁷ Mo	⁹² Pd	¹⁰¹ In	¹¹⁵ Te	¹⁵⁷ Xe	¹⁸⁵ Ba	¹²⁹ Nd	¹³⁷ Eu	¹⁴⁵ Ho	¹⁸⁶ Hf	²⁴⁷ Pt
³³ S	⁷² Fe	⁹⁶ Zn	¹¹⁰ Se	¹²³ Sr	¹¹⁸ Mo	⁹⁴ Pd	¹⁰³ In	¹³³ Te	¹⁵⁸ Xe	¹⁸⁶ Ba	¹³² Nd	¹³² Gd	¹⁴⁷ Ho	¹⁸⁷ Hf	²⁵³ Pt
⁴³ S	⁷³ Fe	⁹⁷ Zn	¹¹¹ Se	¹²⁴ Sr	¹¹⁹ Mo	⁹⁵ Pd	¹⁰⁵ In	¹³⁵ Te	¹⁵⁹ Xe	¹⁸⁷ Ba	¹³³ Nd	¹³³ Gd	¹⁷⁰ Ho	¹⁸⁸ Hf	²⁵⁵ Pt
⁴⁵ S	⁷⁴ Fe	⁹⁹ Zn	¹¹² Se	¹²⁵ Sr	¹²⁰ Mo	⁹⁶ Pd	¹⁰⁷ In	¹³⁶ Te	¹⁶⁰ Xe	¹⁸⁸ Ba	¹³⁵ Nd	¹³⁴ <u>Gd</u>	¹⁷² Ho	¹⁹⁰ Hf	²⁵⁶ Pt
³⁸ Cl	⁷⁹ Fe	¹⁰⁰ Zn	¹¹⁵ Se	¹²⁶ Sr	¹²¹ Mo	⁹⁷ Pd	¹⁴⁰ In	¹³⁷ Te	¹⁶¹ Xe	¹⁸⁹ Ba	¹⁴⁷ Nd	¹³⁵ Gd	¹⁴⁵ Er	¹⁹² Hf	²⁵⁷ Pt
³⁹ Cl	⁸² Fe	¹⁰¹ Zn	¹¹⁷ Se	¹²⁸ Sr	¹²³ Mo	⁹⁸ Pd	¹⁴² In	¹³⁹ Te	¹⁶² Xe	¹⁴⁰ La	¹⁵¹ Nd	¹³⁷ Gd	¹⁴⁶ Er	¹⁹³ Hf	²⁵⁸ Pt
⁶¹ Cl	⁸⁴ Fe	¹⁰³ Zn	¹¹⁸ Se	¹²⁹ Sr	¹²⁵ Mo	¹⁰¹ Pd	¹⁴⁸ In	¹⁴⁰ Te	¹⁶⁴ Xe	¹⁴² La	¹⁵³ Nd	¹³⁸ Gd	¹⁵⁴ Er	¹⁹⁴ Hf	²⁵⁹ Pt
³⁷ Ar	⁸⁵ Fe	⁶² Ga	⁹² Br	¹¹⁰ Y	¹²⁷ Mo	¹⁰⁷ Pd	⁹⁹ <u>Sn</u>	¹⁴² Te	¹⁶⁶ Xe	¹⁴⁴ La	¹⁵⁵ Nd	¹⁴⁰ Gd	¹⁵⁸ Er	¹⁹⁵ Hf	¹⁹⁸ Au
³⁹ Ar	⁸⁷ Fe	⁶⁹ Ga	¹⁰⁸ Br	¹²⁶ Y	¹²⁸ Mo	¹⁰⁹ Pd	¹⁰⁰ <u>Sn</u>	¹⁴³ Te	¹⁶⁷ Xe	¹⁴⁸ La	¹⁵⁷ Nd	¹⁴² Gd	¹⁶⁰ Er	¹⁹⁷ Hf	¹⁷³ Hg
⁴² Ar	⁸⁸ Fe	⁷⁰ Ga	⁸⁰ Kr	⁸⁰ Zr	¹³⁰ Mo	¹¹⁶ Pd	¹⁰¹ Sn	¹⁴⁴ Te	¹⁶⁸ Xe	¹⁵⁰ La	¹⁵⁸ Nd	¹⁴⁴ Gd	¹⁶² Er	¹⁹⁸ Hf	¹⁷⁵ Hg
⁴³ Ar	⁸⁹ Fe	⁸⁶ Ga	⁷² Kr	⁸¹ Zr	¹³² Mo	¹²⁴ Pd	¹⁰⁷ Sn	¹⁴⁵ Te	¹⁶⁹ Xe	¹⁵² La	¹⁶⁰ Nd	¹⁴⁶ Gd	¹⁶⁴ Er	¹⁹⁹ Hf	¹⁷⁷ Hg
⁴⁴ Ar	⁵⁹ Co	⁶² Ge	⁷³ Kr	⁸⁴ Zr	¹³³ Mo	¹²⁵ Pd	¹¹⁷ Sn	¹⁴⁶ Te	¹⁷⁰ Xe	¹⁵⁴ La	¹⁶¹ Nd	¹⁵¹ Gd	¹⁶⁶ Er	²³¹ Hf	¹⁸⁷ Hg
⁴⁷ Ar	⁶¹ Co	⁶⁴ Ge	⁷⁴ Kr	⁸⁵ Zr	¹³⁴ Mo	¹²⁷ Pd	¹¹⁸ Sn	¹⁴⁷ Te	¹⁷¹ Xe	¹⁹⁰ La	¹⁶² Nd	¹⁶¹ Gd	¹⁶⁸ Er	²³³ Hf	¹⁹⁹ Hg
⁵⁷ Ar	⁶³ Co	⁶⁶ Ge	⁷⁵ Kr	⁸⁶ Zr	¹³⁵ Mo	¹²⁹ Pd	¹²¹ Sn	¹⁴⁸ Te	¹⁷² Xe	¹¹⁶ Ce	¹⁶³ Nd	¹⁶² Gd	¹⁷⁰ Er	²³⁵ Hf	²⁰⁹ Hg
³⁹ K	⁷¹ Co	⁶⁸ Ge	⁸⁰ Kr	⁸⁷ Zr	¹³⁶ Mo	¹³¹ Pd	¹²⁵ Sn	¹⁴⁹ Te	¹⁷³ Xe	¹¹⁸ Ce	¹⁶⁴ Nd	¹⁶⁴ Gd	¹⁷¹ Er	¹⁵⁸ Ta	²¹⁴ Hg
³⁹ Ca	⁵⁶ Ni	⁷⁰ Ge	⁸² Kr	⁸⁸ Zr	¹³⁷ Mo	¹³² Pd	¹³¹ Sn	¹⁵⁰ Te	¹⁷⁴ Xe	¹²⁰ Ce	¹⁶⁵ Nd	¹⁶⁵ Gd	¹⁷² Er	¹⁵⁷ W	²⁴³ Hg
⁴⁰ Ca	⁵⁷ Ni	⁸² Ge	⁸⁵ Kr	⁸⁹ Zr	¹⁴³ Mo	¹³⁴ Pd	¹³² Sn	¹⁵² Te	¹⁷⁵ Xe	¹²⁴ Ce	¹⁶⁶ Nd	¹⁶⁶ Gd	¹⁷³ Er	¹⁵⁸ W	²⁴⁶ Hg
⁴¹ Ca	⁵⁸ Ni	⁸⁵ Ge	⁹¹ Kr	⁹⁰ Zr	¹⁴⁴ Mo	¹³⁵ Pd	¹³³ Sn	¹⁵³ Te	¹⁷⁶ Xe	¹²⁵ Ce	¹⁶⁷ Nd	¹⁶⁷ Gd	¹⁷⁴ Er	¹⁶¹ W	²⁴⁷ Hg
⁴² Ca	⁵⁹ Ni	⁸⁷ Ge	⁹² Kr	⁹¹ Zr	¹²⁸ Tc	¹³⁶ Pd	¹³⁴ Sn	¹⁵⁴ Te	¹⁷⁹ Xe	¹²⁷ Ce	¹⁶⁸ Nd	¹⁷⁰ Gd	¹⁷⁵ Er	¹⁶⁶ W	²⁴⁸ Hg
⁴³ Ca	⁷¹ Ni	⁸⁸ Ge	⁹³ Kr	⁹⁹ Zr	¹³⁴ Tc	¹³⁷ Pd	¹³⁵ Sn	¹⁵⁵ Te	¹⁸⁰ Xe	¹³¹ Ce	¹⁷⁰ Nd	¹⁷¹ Gd	¹⁷⁶ Er	¹⁷¹ W	²⁴⁹ Hg
⁴⁴ Ca	⁷² Ni	⁸⁹ Ge	⁹⁴ Kr	¹⁰³ Zr	⁸⁸ Ru	¹³⁸ Pd	¹³⁶ Sn	¹⁵⁶ Te	¹⁸¹ Xe	¹³⁹ Ce	¹⁷² Nd	¹⁷² Gd	¹⁷⁸ Er	¹⁹⁰ W	²⁵³ Hg
⁴⁵ Ca	⁷³ Ni	⁹¹ Ge	⁹⁵ Kr	¹⁰⁴ Zr	⁸⁹ Ru	¹³⁹ Pd	¹³⁷ Sn	¹⁵⁷ Te	¹¹¹ Cs	¹⁴⁴ Ce	¹⁷³ Nd	¹⁷⁶ Gd	¹⁸⁰ Er	¹⁹¹ W	²⁵⁵ Hg
⁴⁷ Ca	⁷⁴ Ni	⁹² Ge	⁹⁶ Kr	¹⁰⁵ Zr	⁹⁰ Ru	¹⁴⁰ Pd	¹³⁸ Sn	¹⁵⁸ Te	¹¹⁹ Cs	¹⁴⁸ Ce	¹⁷⁵ Nd	¹⁷⁷ Gd	¹⁸¹ Er	¹⁹² W	²⁵⁶ Hg
⁴⁹ Ca	⁷⁵ Ni	⁹³ Ge	⁹⁷ Kr	¹⁰⁶ Zr	⁹¹ Ru	¹⁴¹ Pd	¹³⁹ Sn	¹⁵⁹ Te	¹⁵⁶ Cs	¹⁵⁰ Ce	¹⁷⁶ Nd	¹⁷⁸ Gd	¹⁸² Er	¹⁹⁴ W	²⁵⁷ Hg
⁶⁴ Ca	⁷⁷ Ni	⁹⁵ Ge	⁹⁹ Kr	¹⁰⁷ Zr	⁹² Ru	¹⁴² Pd	¹⁴⁰ Sn	¹⁶⁰ Te	¹¹² Ba	¹⁵¹ Ce	¹⁷⁷ Nd	¹⁷⁹ Gd	¹⁸³ Er	¹⁹⁵ W	²⁵⁹ Hg
⁶⁵ Ca	⁷⁸ Ni	⁹⁷ Ge	¹⁰⁰ Kr	¹⁰⁹ Zr	⁹³ Ru	¹⁴³ Pd	¹⁴¹ Sn	¹⁶² Te	¹¹⁴ Ba	¹⁵³ Ce	¹⁷⁸ Nd	¹⁸¹ Gd	¹⁸⁴ Er	¹⁹⁶ W	²⁶¹ Hg
⁶⁶ Ca	⁸⁴ Ni	⁹⁹ Ge	¹⁰¹ Kr	¹¹⁰ Zr	⁹⁴ Ru	¹⁴⁴ Pd	¹⁴² Sn	¹⁶³ Te	¹²⁰ Ba	¹⁵⁴ Ce	¹⁸⁰ Nd	¹⁸² Gd	¹⁸⁵ Er	¹⁹⁷ W	²⁶³ Hg
⁶⁷ Ca	⁸⁵ Ni	¹⁰⁰ Ge	¹⁰² Kr	¹¹² Zr	⁹⁵ Ru	¹⁴⁶ Pd	¹⁴⁴ Sn	¹⁶⁴ Te	¹²¹ Ba	¹⁵⁵ Ce	¹⁸¹ Nd	¹⁸⁴ Gd	¹⁸⁸ Er	¹⁹⁹ W	²⁶⁵ Hg
⁶⁸ Ca	⁸⁶ Ni	¹⁰¹ Ge	¹⁰³ Kr	¹¹³ Zr	⁹⁶ Ru	¹⁴⁷ Pd	¹⁴⁵ Sn	¹⁶⁶ Te	¹²² Ba	¹⁵⁶ Ce	¹⁸² Nd	¹⁸⁶ Gd	¹⁹⁰ Er	²⁰⁰ W	²⁰² Tl
⁶⁹ Ca	⁸⁷ Ni	¹⁰³ Ge	¹⁰⁴ Kr	¹¹⁴ Zr	⁹⁹ Ru	¹⁴⁸ Pd	¹⁴⁶ Sn	¹⁶⁷ Te	¹²³ Ba	¹⁵⁸ Ce	¹⁸⁵ Nd	¹⁸⁷ Gd	¹⁹¹ Er	²⁰¹ W	¹⁷⁹ Pb
⁴⁹ Sc	⁸⁸ Ni	¹⁰⁴ Ge	¹⁰⁵ Kr	¹¹⁵ Zr	¹⁰⁷ Ru	¹⁴⁹ Pd	¹⁴⁷ Sn	¹⁶⁸ Te	¹³⁴ Ba	¹⁵⁹ Ce	¹⁸⁶ Nd	¹⁸⁹ Gd	¹⁹³ Er	²⁰² W	¹⁸⁰ Pb
⁴⁴ Ti	⁸⁹ Ni	¹⁰⁵ Ge	¹⁰⁸ Kr	¹¹⁶ Zr	¹⁰⁹ Ru	¹⁵⁰ Pd	¹⁴⁸ Sn	¹⁶⁹ Te	¹³⁷ Ba	¹⁶¹ Ce	¹⁸⁷ Nd	¹⁹⁰ Gd	²¹⁷ Er	²³² W	¹⁸¹ Pb
⁴⁶ Ti	⁹¹ Ni	¹⁰⁶ Ge	¹⁰⁹ Kr	¹¹⁷ Zr	¹¹⁴ Ru	¹⁵¹ Pd	¹⁵⁰ Sn	¹⁷⁰ Te	¹⁴¹ Ba	¹⁶² Ce	¹⁸⁸ Nd	¹⁹² Gd	²¹⁸ Er	²⁴¹ W	¹⁸² Pb
⁴⁸ Ti	⁹² Ni	¹⁰⁸ Ge	¹¹³ Kr	¹¹⁸ Zr	¹¹⁵ Ru	¹⁵² Pd	¹⁵¹ Sn	¹⁷¹ Te	¹⁴⁶ Ba	¹⁶³ Ce	¹⁹⁴ Nd	¹⁹⁸ Gd	²¹⁹ Er	¹⁶⁰ Re	¹⁸⁵ Pb
⁴⁹ Ti	⁹³ Ni	¹⁰⁹ Ge	¹¹⁴ Kr	¹²¹ Zr	¹¹⁷ Ru	¹⁵³ Pd	¹⁵² Sn	¹⁷³ Te	¹⁴⁷ Ba	¹⁶⁴ Ce	¹⁹⁶ Nd	²⁰² Gd	²²⁶ Er	¹⁸⁴ Re	¹⁸⁷ Pb
⁵⁰ Ti	⁹⁵ Ni	¹¹¹ Ge	¹¹⁵ Kr	¹²² Zr	¹¹⁹ Ru	⁹⁷ Ag	¹⁵³ Sn	¹⁷⁶ Te	¹⁴⁸ Ba	¹⁶⁵ Ce	¹³¹ Pm	²⁰⁸ Gd	²²⁷ Er	¹⁹⁴ Re	¹⁸⁹ Pb
⁵¹ Ti	⁹⁷ Ni	¹¹² Ge	¹¹⁷ Kr	¹²³ Zr	¹²¹ Ru	⁹⁹ Ag	¹⁵⁴ Sn	¹³⁶ I	¹⁴⁹ Ba	¹⁶⁶ Ce	¹⁵⁸ Pm	²⁰⁹ Gd	¹⁴⁵ <u>Tm</u>	¹⁶⁵ Os	¹⁹¹ Pb
⁶² Ti	⁹⁹ Ni	⁶⁸ As	¹¹⁹ Kr	¹²⁵ Zr	¹²⁵ Ru	¹²³ Ag	¹⁵⁵ Sn	¹⁴⁰ I	¹⁵⁰ Ba	¹⁶⁸ Ce	²¹¹ Gd	¹⁴⁷ <u>Tm</u>	¹⁵¹ Yb	¹⁶⁸ Os	¹⁹⁵ Pb
⁶³ Ti	⁵⁸ Cu	⁷⁴ As	¹²³ Kr	¹²⁶ Zr	¹²⁶ Ru	¹²⁵ Ag	¹⁵⁶ Sn	¹⁴² I	¹⁵¹ Ba	¹⁷⁰ Ce	¹²⁸ Sm	²¹³ Gd	¹⁵¹ Yb	¹⁷⁰ Os	¹⁹⁷ Pb
⁶⁴ Ti	⁵⁹ Cu	⁷⁸ As	¹²⁴ Kr	¹²⁷ Zr	¹²⁷ Ru	¹⁴⁰ Ag	¹⁵⁷ Sn	¹⁵⁴ I	¹⁵² Ba	¹⁷¹ Ce	¹³⁰ <u>Sm</u>	¹⁶⁶ Tb	¹⁵³ Yb	¹⁷³ Os	¹⁹⁹ Pb
⁶⁶ Ti	⁶¹ Cu	⁹⁰ As	⁸⁰ Rb	¹²⁸ Zr	¹²⁹ Ru	⁹⁷ Cd	¹⁵⁸ Sn	¹⁶⁸ I	¹⁵³ Ba	¹⁷² Ce	¹³¹ Sm	¹⁶⁸ Tb	¹⁵⁴ Yb	¹⁸⁰ Os	²⁰¹ Pb
⁶⁸ Ti	⁶³ Cu	⁶⁶ Se	¹²⁴ Rb	¹²⁹ Zr	¹³⁰ Ru	¹³⁰ Cd	¹⁶⁰ Sn	¹⁷² I	¹⁵⁴ Ba	¹⁷³ Ce	¹³² Sm	¹³⁷ Dy	¹⁵⁵ Yb	¹⁸² Os	²¹³ Pb
⁶⁹ Ti	⁶⁵ Cu	⁶⁸ Se	⁸⁰ Sr	¹³⁰ Zr	¹³² Ru	¹³¹ Cd	¹⁶¹ Sn	¹¹⁰ Xe	¹⁵⁶ Ba	¹⁷⁴ Ce	¹³³ Sm	¹³⁹ Dy	¹⁵⁷ Yb	¹⁹⁵ Os	²¹⁴ Pb
⁵¹ V	⁶⁷ Cu	⁷⁰ Se	⁸² Sr	¹³¹ Zr	¹³³ Ru	¹³³ Cd	¹⁶² Sn	¹¹² <u>Xe</u>	¹⁵⁷ Ba	¹⁷⁵ Ce	¹⁵⁷ Sm	¹⁴⁰ Dy	¹⁶⁰ Yb	¹⁹⁶ Os	²¹⁶ Pb
⁵⁰ Cr	⁶⁹ Cu	⁷² Se	⁸⁴ Sr	¹³² Zr	¹³⁴ Ru	¹³⁵ Cd	¹⁶³ Sn	¹¹⁵ <u>Xe</u>	¹⁵⁸ Ba	¹⁷⁶ Ce	¹⁵⁹ Sm	¹⁴¹ Dy	¹⁶¹ Yb	¹⁹⁸ Os	²⁴⁶ Pb
⁵¹ Cr	⁷¹ Cu	⁷⁴ Se	⁸⁵ Sr	¹³³ Zr	¹³⁵ Ru	¹³⁶ Cd	¹⁶⁵ Sn	¹²⁶ <u>Xe</u>	¹⁵⁹ Ba	¹⁷⁷ Ce	¹⁶⁰ Sm	¹⁴³ Dy	¹⁸⁵ Yb	²⁰¹ Os	²⁵² Pb
⁵² Cr	⁵⁸ Zn	⁸⁵ Se	⁸⁶ Sr	¹²⁶ Nb	¹³⁶ Ru	¹³⁷ Cd	¹⁶⁷ Sn	¹³¹ <u>Xe</u>	¹⁶¹ Ba	¹⁷⁸ Ce	¹⁶³ Sm	¹⁴⁵ Dy	¹⁸⁷ Yb	²⁰³ Os	²⁵³ Pb
⁵³ Cr	⁵⁹ Zn	⁸⁷ Se	⁸⁷ Sr	¹²⁸ Nb	¹³⁷ Ru	¹³⁸ Cd	¹⁶⁸ Sn	¹³³ <u>Xe</u>	¹⁶² Ba	¹⁷⁹ Ce	¹⁶⁴ Sm	¹⁴⁶ Dy	¹⁸⁸ Yb	²⁰⁶ Os	²⁵⁴ Pb
⁶⁵ Cr	⁶⁰ Zn	⁸⁸ Se	⁸⁸ Sr	¹³⁸ Ru	¹⁴⁰ Cd	¹⁶⁹ Sn	¹³⁵ Xe	¹³⁵ <u>Xe</u>	¹⁶³ Ba	¹⁸³ Ce	¹⁶⁵ Sm	¹⁵¹ Dy	¹⁹⁰ Yb	²⁵⁰ Os	²⁵⁷ Pb
⁶⁶ Cr	⁶² Zn	⁸⁹ Se	⁸⁹ Sr	⁸⁶ Mo											

TABLE IX. Same as Table I but for (p, α).

³¹ Si	⁴⁸ Ar	⁶² Ar	⁷⁰ Ca	⁷⁰ Se	⁸⁶ Zr	⁹² Mo
³⁴ Si	⁵⁶ Ar	⁶³ Ar	⁵¹ Ti	¹¹⁶ Se	⁸⁷ Zr	⁸⁶ Tc
⁴⁶ Si	⁵⁸ Ar	⁴³ Ca	⁷² Ti	⁷⁰ Br	⁸⁰ Mo	⁹¹ Ru
⁴⁶ S	⁵⁹ Ar	⁴⁷ Ca	⁴⁸ Mn	⁷⁴ Kr	⁸⁹ Mo	⁹³ Ru
⁴⁴ Ar	⁶⁰ Ar	⁶⁸ Ca	⁵⁷ Zn	⁷⁴ Rb	⁹⁰ Mo	⁹⁵ Cd
⁴⁷ Ar	⁶¹ Ar	⁶⁹ Ca	⁶¹ Ge	⁸⁶ Sr	⁹¹ Mo	⁹⁹ Cd