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JINGLE, a JCMT legacy survey of dust and gas for galaxy evolution studies: II. SCUBA-2 850 μm data reduction and dust flux density catalogues

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ABSTRACT

We present the SCUBA-2 850 μm component of JINGLE, the new JCMT large survey for dust and gas in nearby galaxies, which with 193 galaxies is the largest targeted survey of nearby galaxies at 850 μm . We provide details of our SCUBA-2 data reduction pipeline, optimised for slightly extended sources, and including a calibration model adjusted to match conventions used in other far-infrared data. We measure total integrated fluxes for the entire JINGLE sample in 10 infrared/submillimetre bands, including all WISE, *Herschel*-PACS, *Herschel*-SPIRE and SCUBA-2 850 μm maps, statistically accounting for the contamination by CO($J=3-2$) in the 850 μm band. Of our initial sample of 193 galaxies, 191 are detected at 250 μm with a $\geq 5\sigma$ significance. In the SCUBA-2 850 μm band we detect 126 galaxies with $\geq 3\sigma$ significance. The distribution of the JINGLE galaxies in far-infrared/sub-millimetre colour-colour plots reveals that the sample is not well fit by single modified-blackbody models that assume a single dust-emissivity index (β). Instead, our new 850 μm data suggest either that a large fraction of our objects require $\beta < 1.5$, or that a model allowing for an excess of sub-mm emission (e.g., a broken dust emissivity law, or a very cold dust component $\lesssim 10\text{K}$) is required. We provide relations to convert far-infrared colours to dust temperature and β for JINGLE-like galaxies. For JINGLE the FIR colours correlate more strongly with star-formation rate surface-density rather than the stellar surface-density, suggesting heating of dust is greater due to younger rather than older stellar-populations, consistent with the low proportion of early-type galaxies in the sample.

Key words: galaxies: ISM – galaxies: photometry – submillimetre: ISM – galaxies: spiral

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1 INTRODUCTION

Studies of the interstellar medium in large, varied galaxy samples are crucial to our understanding of star formation and galaxy evolution. Surveys of both atomic gas (via the HI 21cm line) and molecular gas (often traced by emission lines of the CO molecule) have revealed that there are key scaling relations in the local universe between global galaxy properties and the contents of the interstellar medium (ISM) (e.g. Roberts & Haynes 1994; Young et al. 1995; Catinella et al. 2010; Saintonge et al. 2011; Bothwell et al. 2013; Boselli et al. 2014b). For example the Schmidt-Kennicutt law (Schmidt 1959; Kennicutt 1989) relates the surface density of star-formation to the surface density of gas in the galaxy. Large studies using HI, dust continuum emission in the far-infrared (FIR), CO and other molecular line tracers have revealed that these scaling laws can depend on factors such as morphological type, mass, and environment (e.g. Young et al. 2011; Lisenfeld et al. 2011; Cortese et al. 2011; Smith et al. 2012a; Cortese et al. 2012; Clark et al. 2015; De Vis et al. 2017; Rémy-Ruyer et al. 2014).

Studies of dust in the ISM are important as over the lifetime of the Universe half of all light emitted from stars has been absorbed by dust and then re-emitted in the far-infrared (Lagache et al. 2005). Stars are formed in dense clouds of gas and dust, and so far-infrared observations can be vital for measuring an accurate star-formation rate in galaxies due to absorption of the UV and optical light (Kennicutt 1998; Calzetti 2001). Dust is important for molecules in the ISM as it catalyses reactions as atoms bind to the surface of dust grains (e.g. Gould & Salpeter 1963; Hagen et al. 1979; van Dishoeck 2004). Given the difficulty with directly measuring molecular gas with CO or other tracers, and this especially at high redshifts, dust is also seen as a promising probe of the entire cold ISM (e.g. Guélin et al. 1993, 1995; Israel 1997; Scoville et al. 2014).

While there have been surveys with observations of HI, CO and dust continuum, they are the exceptions. One example is the SINGS sample (Kennicutt et al. 2003) which targeted ~ 70 galaxies with distances < 30 Mpc, obtaining dust continuum data with *Herschel* (KINGFISH, Kennicutt et al. 2011), and exquisite gas measurements in HI (THINGS, Walter et al. 2008) and CO (HERACLES, Leroy et al. 2009). Another example is the *Herschel* Reference Survey (HRS, Boselli et al. 2010a) which targeted 322 *K*-band selected galaxies in a volume-limited sample, and has collected data on all three components of the cold ISM (Ciesla et al. 2012; Cortese et al. 2014; Boselli et al. 2014a). To make sure to sample a full range of galaxy properties, large statistical samples beyond the very local Universe focusing on these different components of the cold ISM are required. Using the James Clark Maxwell Telescope (JCMT), the JCMT dust and gas In Nearby Galaxies Legacy Exploration (JINGLE; Saintonge et al. 2018, hereafter Paper I), aims to address this. JINGLE has observed dust continuum at $850\ \mu\text{m}$ for 193 SDSS-selected galaxies ($M_* > 10^9 M_\odot$), and CO($J=2-1$) line emission for a subset of $\sim 35\%$ of them. The sample was selected in fields with coverage from *Herschel*-ATLAS (hereafter *H*-ATLAS, Eales et al. 2010), which observes between $100\text{--}500\ \mu\text{m}$, the Arecibo ALFALFA HI survey (Giovanelli et al. 2005), and the MaNGA and SAMI optical integral field spectroscopic surveys (Bundy et al. 2015; Bryant

et al. 2015), providing additionally the all important information about dust, atomic gas, ionised gas, and resolved stellar properties.

The SCUBA-2 data (the subject of this paper) of 193 galaxies is the largest targeted survey of nearby galaxies at $850\ \mu\text{m}$. Adding a data point far down the Rayleigh–Jeans tail provides an improvement on obtaining dust measurements, over just using *Herschel* data between $70\text{--}500\ \mu\text{m}$. For example, there can be a degeneracy in fitted parameters between the dust emissivity index (β , see Section 6.3) and dust temperature measured (Smith et al. 2012b). By adding a data-point at $850\ \mu\text{m}$ further down the Rayleigh–Jeans tail of blackbody emission the degeneracy can be reduced by approximately a factor of two. An additional benefit is that the *Herschel* SPIRE calibration is dominated by correlated uncertainties (4% correlated versus 1.5% uncorrelated) between bands due to uncertainties in planet models (Bendo et al. 2013), which can have a significant effect on SED fitting results. With an independent longer wavelength point at $850\ \mu\text{m}$ (as SCUBA-2 uses different calibrators) we can obtain a more accurate measurement of β , which is important as it can be an indication of grain size, composition, or surface mantles. Differences in the assumed numerical value of β can have a significant effect on the dust-mass absorption coefficient (κ_d) and consequently on the dust masses that are inferred from far-infrared observations (e.g. Bianchi 2013; Clark et al. 2016).

In this paper we present our observing strategy (Section 2.2) and our customised data reduction pipeline, which is optimised for the specific requirements of the JINGLE survey (see Section 3). In Section 4 we describe our flux extraction methods for both point and extended sources, and describe our simulations to account for any flux attenuation from the pipeline high-pass filter. Section 5 describes our approach to remove CO($J=3-2$) contamination of our $850\ \mu\text{m}$ fluxes. Finally, in Section 6 we investigate FIR/submm colour ratios, and how they relate to model fits of temperature, dust emissivity, and other physical properties of the galaxy.

2 SAMPLE AND OBSERVATIONS

2.1 JINGLE sample overview

The sources in the JINGLE survey are selected based on detections from the *H*-ATLAS survey (Eales et al. 2010) which observed ~ 600 sq. degrees with *Herschel* (Pilbratt et al. 2010). *H*-ATLAS observed in parallel mode using PACS at 100 and $160\ \mu\text{m}$, and SPIRE at 250 , 350 and $500\ \mu\text{m}$ simultaneously. As JCMT is in the northern hemisphere, JINGLE selected objects in the equatorial GAMA fields ($161\ \text{deg}^2$) and the North Galactic Pole (NGP) field ($180.1\ \text{deg}^2$). For our photometry (Section 4) we use the PACS maps provided in *H*-ATLAS DR1 (GAMA fields, Valiante et al. 2016) and DR2 (NGP, Smith et al. 2017). For SPIRE we use the same timelines used to generate the maps in the DR1 and DR2 releases, except we apply the relative gain corrections and calibration corrections to optimise the maps for extended sources (these maps are also made available in the JINGLE data release).

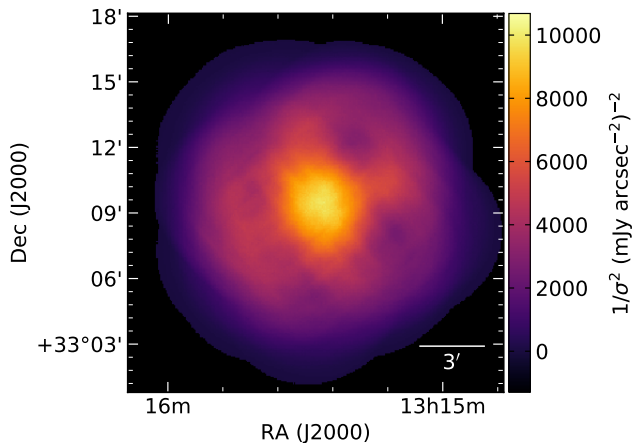


Figure 1. An example weight image for a typical observation, which illustrates how the sensitivity of our observation decreases with radius. The weight image is calculated from the inverse square of the noise.

2.2 SCUBA-2 observations

In this section we outline the strategy used by the JINGLE survey to obtain good quality maps at $850\ \mu\text{m}$ with SCUBA-2 (we also simultaneously obtain $450\ \mu\text{m}$ but as the sensitivity is lower than SPIRE $500\ \mu\text{m}$ we only concentrate on $850\ \mu\text{m}$ in this work). The redshift selection of the JINGLE survey ($0.01 < z < 0.05$) puts our targets at a distance of $\sim 41\text{--}212\ \text{Mpc}$ with a mean D_{25} (the B -band isophotal diameter at a surface brightness of $25\ \text{mag arcsec}^{-2}$) of $1.28'$, with the largest size of $4.6'$. However, from the SPIRE $250\ \mu\text{m}$ maps the largest aperture required to accurately measure the flux of our objects has a diameter of $2.2'$. For all JINGLE targets we therefore observe using the ‘Constant Velocity (CV) Daisy’ mapping mode which is the smallest observing mode available with SCUBA-2, and provides an even coverage in the central $3'$ of the observation. The ‘CV Daisy’ is a circular scanning pattern designed so that the target is always within the field-of-view of the array throughout the integration while moving at a constant $155''/\text{s}$. The observation provides useable coverage out to $\sim 6.0'$ in radius, but beyond $1.5'$ the map sensitivity decreases rapidly. A typical weight map ($1/\sigma^2$) for a JINGLE Daisy observation is shown in Figure 1.

The sensitivity estimates for each JINGLE target were derived by fitting a modified-blackbody model to the *Herschel* data (fluxes were taken from an initial *H-ATLAS* catalogue). As literature values of the dust-emissivity index (β) tend to lie in the region $1.5\text{--}2.0$ (Smith et al. 2012b; Planck Collaboration et al. 2011; Cortese et al. 2014; Planck Collaboration et al. 2014), we use a constant value of $\beta = 2$, which should be a conservative estimate of the $850\ \mu\text{m}$ flux. To account for JINGLE galaxies being extended we use the radius of the galaxy to divide the total flux evenly between $1\text{--}8$ independent elements. The SCUBA-2 exposure time calculator was then used to calculate the total integration time required to achieve a signal-to-noise ratio of 5. The estimates were made assuming matched-filtering which as we discuss in Section 3 is not appropriate for our data, this is slightly

compensated by our decision to use pixel sizes larger than the $1''$ default (when the survey was proposed) in the integration time calculator. The weather band (dependent on the amount of water vapour in the atmosphere) was chosen so the total integration time is always under 2 hours. Overall our observations were generally taken in slightly better conditions than assumed when using the exposure time calculator.

3 DATA REDUCTION

The JINGLE survey was designed specifically so as to minimise problems with emission on large-scales, which has been a problem with ground-based imaging due to atmospheric and instrumental variations often requiring spatial filtering to be applied. To try avoid this issue the minimum redshift of $z > 0.01$ is used to select targets without large angular sizes. However, by having isolated objects which are resolved, but do not have extended structures (like galactic regions or very-nearby galaxies) means JINGLE is between the two extremes of point sources or preserving large-scale emission for which SCUBA-2 pipelines have been developed. We therefore, have worked to customise the data reduction pipeline for our specific case of marginally resolved objects. In this section we describe our customised pipeline.

3.1 SCUBA-2 Data Processing

To reduce the raw SCUBA-2 data and create maps, we use the default SCUBA-2 map-maker MAKEMAP provided as part of the STARLINK software package (Currie et al. 2014)¹. We refer to Chapin et al. (2013) for full details of the MAKEMAP algorithm, but briefly, after an initial ‘cleaning’ stage, which removes bad bolometers and artefacts such as glitches, MAKEMAP begins an iterative procedure to split the bolometer signals into various components. In each iteration a ‘common-mode’ signal predominately from sky noise is identified and subtracted, an extinction correction is applied, and then the bolometer signals are high-pass filtered. From these timelines (time-ordered detector readouts) a map is made to identify astronomical emission (AST model) which is subtracted from the next iteration. This iterative procedure continues until the map converges (see below), or reaches the maximum number of iterations. This technique is highly-customisable, allowing settings at every step to be adjusted to optimise the data for your science target. Below we describe how we optimised the JINGLE data for our case of slightly extended sources.

The standard MAKEMAP implementation processes each observation individually and then combines the individual maps at the end. If the memory requirements are too big for a machine, the observation is split into ‘chunks’, each processed separately (note for all JINGLE processing a machine with enough memory is used so no ‘chunking’ is required). While this is computationally efficient it does not make best use of the data, instead the astronomical model in each iteration should be estimated from all the observations due to improved sensitivity and better resilience to

¹ All maps were created using STARLINK package version 2017A

atmospheric/instrumental variations. The SKYLOOP script provided in STARLINK solves this problem, and is designed to help recover more signal, by using all the data at the end of every iteration. Our script is modified from the standard SKYLOOP as we found that for all but the first observation passed to MAKEMAP, the proportion of data flagged became excessively high. We modified the script to output a map for each observation after every iteration and then mosaic the data, instead of feeding all the data to MAKEMAP in one go. After discussions with the observatory, SKYLOOP in the STARLINK 2017A package has been updated so this problem no longer exists.

The second difference is our method to apply source masking to the data. In SCUBA-2 terminology ‘Masking’ is where regions are selected where the AST model is allowed to vary, in any other area of the observation the AST model is set to zero. This improves the reduction as it helps to reduce any degeneracies between AST and other noise models. Such masks can be static (i.e., fixed for all iterations) or dynamic where the mask is generated based on some threshold in the map at each configuration. For JINGLE due to the high-sensitivity of *Herschel* SPIRE observations we have very good prior knowledge of where we expect to find emission at 450 and 850 μm , and can be confident we can select regions without missing any sub-millimetre emission. This masking is very effective for our targets as our galaxies are isolated nearby sources and we therefore use the surrounding areas as background. Unlike other surveys which apply a two-pass approach, where they perform an initial run to create a fixed mask by reducing the data either without a mask or with an adaptive mask, through the use of our SPIRE data we can use our fixed mask. We generate a 2D-image as an input mask which has the advantage over the standard *ast.zero_circle* parameter in MAKEMAP that we can use elliptical shapes and include multiple sources. Our source ellipse is based on the size of the aperture used for our far-infrared photometry which we describe in detail in Section 4.

The behaviour of MAKEMAP is controlled using a configuration file which specifies all the settings to optimise the map-maker for the science required. The JINGLE DR1 configuration file is made up of the following commands:

- *numiter* = -300 $\&$ *maptol* = 0.001 – Mairs et al. (2015) found for The JCMT Gould Belt Survey (Ward-Thompson et al. 2007) that a stricter convergence tolerance helped recover more low-surface brightness structure in their maps. Whether a map has converged is set by the *maptol* parameter which is the threshold value for the mean change of the map between iterations normalised by the root-mean-square of the pixel variances (i.e., when the map stops changing, note that this is calculated from regions not masked). The default of this parameter value is set to 0.05 and so the value of 0.001 is a much more stringent criterion. Like Mairs et al. (2015) to compensate for this stricter limit we increase the number of allowed iterations from a typical value of ~ 40 to 300 (the minus sign in the parameter allows the map to finish if converged, otherwise all iterations would be performed). We found that with our masking we normally converge reasonably quickly (~ 20 iterations).

- *com.perarray* = 1 – This parameter controls whether the ‘common-mode’ component (i.e., atmosphere or instru-

ment variations) is required to be the same across the whole array or whether each sub-array is treated independently. We choose to set this so each sub-array is treated independently; this can remove emission scales above the size of a sub-array ($\sim 3'$), but as we filter at smaller scales we can safely treat each sub-array independently.

- *flt.filt_edge_largescale* = 150 – This parameter has possibly the largest effect when creating SCUBA-2 maps as it controls the scale (in arcseconds) of the high-pass filtering applied to the data. If set too high, spurious signals from the atmosphere or drifts in the array could dominate the astronomical signal, however, if set too low the filtering could remove a significant fraction of the source’s emission. Figure 2 shows one extreme example of the effect of the filter scale when we were testing our data reduction. When the filter scale was set to 100'', JINGLE 119 was not detected in our pipeline. After experimenting with a test sample of objects we settled on a filter scale of 150'' as the best compromise between sensitivity and recovering the flux. In Section 4.3 we describe our simulations to quantify the exact effect of the filtering on each source.

- *ast.zero_mask* = 1 $\&$ *ast.zero_snr* = 0 – These parameters set MAKEMAP to use a static mask (as described above) using a file specified by the REF parameter, and to not apply a threshold to the pixels within this mask to be included in the AST model. Specifying a file via the REF parameter sets the world coordinate system (WCS) for the output image to match the static mask.

- *ast.zero_freeze* = 0, *com.zero_freeze* = 0 $\&$ *flt.zero_freeze* = 0 – These are the default values which state that we do not freeze the AST, common-mode signal (COM), or low-frequency Fourier component (FLT), models after any iteration. These parameters are included as a previous version of skyloop required these parameters to be explicitly set.

We also investigated whether we should set the FLT mask using the same mask defined for the AST model (using the *flt.zero_mask* = 1. In the regions set by the FLT mask the detector timelines are replaced by a linear interpolation before the high-pass filter is applied, with the aim of avoiding ringing around the source. By default this is only applied for the first two iterations and so the signal can be identified in the AST model. We found that this had a large effect on our maps with our extracted flux densities being on average a factor of two higher than our predicted 850 μm fluxes, derived extrapolating modified blackbody fits to *Herschel* $\leq 500 \mu\text{m}$ flux densities. However, our standard reduction agreed well with our predictions (this agrees with our simulations injecting sources into the map in Section 4.3). For some JINGLE targets setting the FLT mask resulted in an 850 μm flux density higher than the SPIRE 500 μm flux, which for a standard local galaxy (with no radio contamination) would be un-physical. As such we do not set the FLT mask for our reduction, with the exception of JINGLE 70 and 132 which in our standard reduction resulted in very significant negative fluxes (-4.7σ and -7.5σ). Setting the FLT mask in these two cases removed the negative artefact and gave fluxes in line with our predictions.

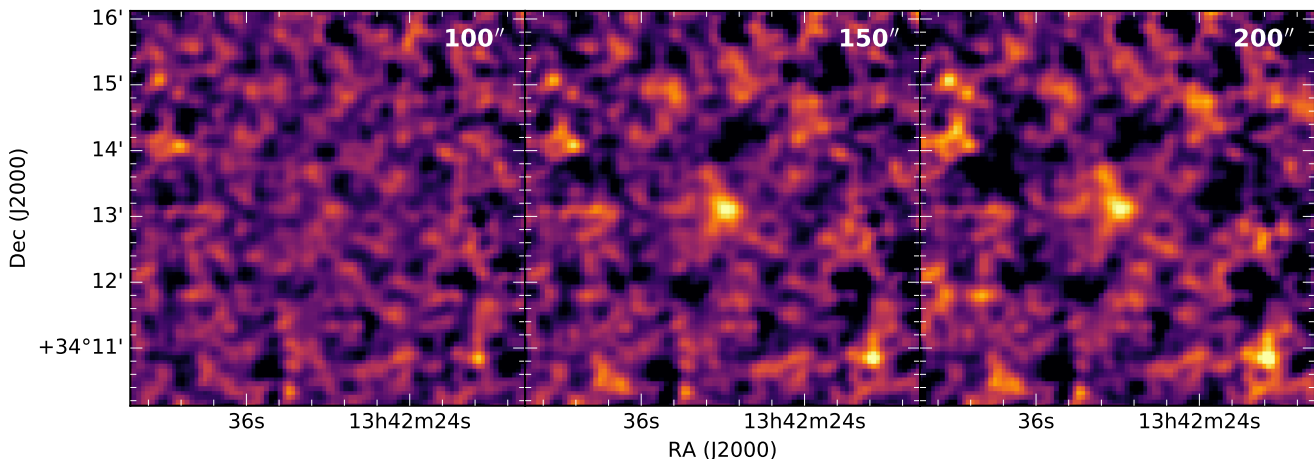


Figure 2. An example of how different filter scales (as set by the *flt.filt_edge_largescale* parameter) affect the resulting SCUBA-2 850 μm map for JINGLE 119. When a 100'' filter (left) is applied our target is filtered out of the map, but increasing the filtering to 200'' (right) leads to increased noise.

3.2 Calibration

The standard procedure at the JCMT is to take a few calibration observations spread over the evening which can be used to see if there are any changes or particular problems over the night. JINGLE calibration observations are based on either Mars, Uranus, CRL618, CRL2688 or Arp220 depending on source visibility. There are good reasons to assume the calibration does vary throughout the night, as the dish cools at the beginning of the night. As the pipeline is very flexible the standard advice is to apply the same procedure to your calibrators as adopted for the target object. For example, the default is to use a 30'' radius aperture for your source and calibrator. However, while this advice is sensible for point sources, it becomes problematic for extended sources where different fractions of the source will be present in apertures, and filtering will have a different effect based on the amount being masked. We instead decide to take an alternative approach, where we optimise the reduction of the calibrators to give the most accurate calibration of our maps (i.e., the calibration that results in the true flux densities of our targets if no attenuation of large angular scales occurs), and then simulate the effect of filtering on our individual targets.

To investigate the calibration we look at calibration observations from the beginning of the project to December 2016, when we had observed just over half the SCUBA-2 sample. When making maps of our calibrators, the aim was to use as similar as possible configuration file as used for our targets, however, some adjustments were required. We discovered when creating maps of our calibrators that, in particular for Mars, the centre of the source was clipped due to the cleaning process. To mitigate this effect we set the *ast.mapspike* parameter to 10, which controls the signal-to-noise of detector samples to remove from a map-pixel, and the *dcthresh* parameter which controls the signal-to-noise ratio of DC step detection to 10000. As the calibration observations are centred on the calibrators, and the sources are all point-like in SCUBA-2 images we set the *ast.zero_circle* and *flt.zero_circle* to 2', which sets the static mask (see Section 3.1) of the AST and FLT models to a

circle in the centre of the image. As with our target map-making in Section 3.1, one of the most important parameters is the filter-scale parameter (*flt.filt_edge_largescale*). Ideally, the filter scale would be set to a very high value so the flux of our calibrator is not attenuated, however, as the calibration observation are short for our dimmer calibrators this would lead to the measurement being dominated by large-scale noise on the map. To test the optimum filter-scale so that our calibrators are not attenuated we take two independent observations of Mars (our brightest calibrator) and create a map for filter-scales from 100'' to 200'' in 10'' intervals and then from 200'' to 400'' in 20'' intervals. Figure 3 shows the fraction of flux measured depending on the filter-scale. From these results we chose to filter our calibrator observations on a scale of 220'' as this is safely on the plateau, but small enough to provide a flat background in the map, and we set *com.perarray* to 0.

To calculate the flux conversion factor (FCF) from our calibrator maps we run the standard PICARD recipe SCUBA2_CHECK_CAL which uses the default settings which matches the method used by Dempsey et al. (2013) who calculated the standard FCF values. As our sources are extended we calibrate our maps in units of mJy arcsec⁻² rather than mJy beam⁻¹. Figure 4 shows the ratio of flux conversion factors to the standard value obtained at 850 μm for our calibrators. The median ratio for our calibrators is 1.002 suggesting our observations are consistent with the standard, although we find the scatter for Arp220 is significantly larger than the other sources, probably due to it being one of the fainter calibrators. The scatter in FCF values result in a calibration uncertainty of 5.7%, similar to the value of 5% found in Dempsey et al. (2013).

The FCF values in Figure 4 are plotted against the time in the day the observation was taken. We see no obvious trend to suggest the FCF is higher at the beginning of the night. However, to check whether using a variable calibration leads to an improvement over assuming a fixed calibration we reduce all maps using both methods. For the variable case we calculate the FCF for each observation based on the linear interpolation between the calibration observation before and after the JINGLE observation (if observations are

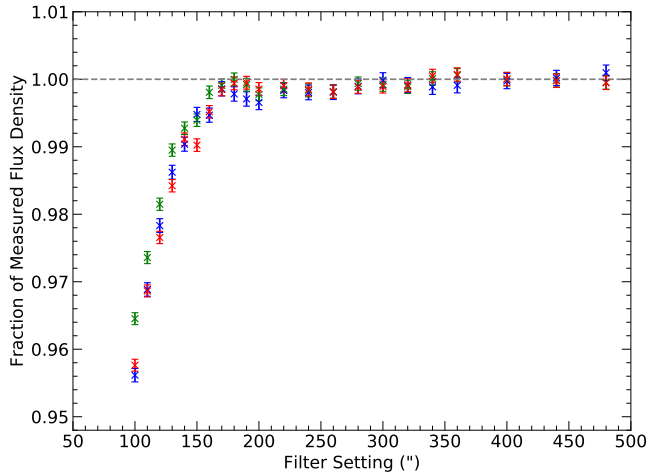


Figure 3. The fraction of flux measured compared to the average value measured in maps where the filter-scale has been set to $>350''$. The blue points correspond to observation id 20151226_00069 with the `com.perarray` set to 0, the green and red points both correspond to observation id 20151227_00067 with the `com.perarray` set to 0 (green) and 1 (red)

not available before and after we use the nearest calibration value). As SKYLOOP takes all the data simultaneously we multiply the raw data for each science observation by the ratio of our variable FCF to standard FCF and then apply the standard FCF to the resultant map from SKYLOOP. To see if there is any improvement in the map we see if there is any reduction in the uncertainty map (accounting for the difference in average FCF value). We did not see any maps with a significant improvement and so have decided to use our fixed values. We will re-visit this in subsequent works where we use the $450\ \mu\text{m}$ data, where the calibration is more likely to vary.

The standard PICARD recipe to calculate the FCF uses a $30''$ radius aperture with a background annulus between radii of $45\text{--}60''$. While these apertures were chosen as a good compromise between encircling enough of the emission and achieving a good signal-to-noise, it systematically overestimates the FCF as it does not account for the fraction of the beam outside the aperture. We decided to modify our FCF value to account for this aperture correction, so our images match the calibration scheme of our other far-infrared datasets. To find this correction we use the measurements that characterise the SCUBA-2 beam from Dempsey et al. (2013), who find it is well described by the sum of two Gaussians, the primary beam with a $13.0''$ FWHM and a secondary with a $48''$ FWHM. By integrating the beam profile from a radius (r) of zero to infinity and within radii of $0\text{--}30''$ and $45\text{--}60''$, we calculate what fraction of the beam is outside the standard aperture and the amount included within the background region. We calculate that the FCF should be multiplied by a factor of 0.910 (this is in reasonable agreement with values measured by, Dempsey et al. 2013) and so we apply an FCF of $2.134\ \text{Jy pW}^{-1}\ \text{arcsec}^{-2}$.

3.3 WCS correction

For a few objects we found cases where there appeared to be an offset between where we expected to see emission at $850\ \mu\text{m}$ from images at other wavelengths (e.g., *Herschel*). After investigating potential cases we found four cases where small offsets seen in our observations were also seen in either the pointing or flux calibration observation performed after the science observations. From these observations we derive offsets of $3.91''$, $4.17''$ and $5.22''$ for JINGLE 35 (J131958.31+281449.3), JINGLE 149 (J125610.97+280947.4) and JINGLE 186 (J132035.40340821.7), respectively. For JINGLE 117, observations on two separate days were found to have offsets of $6.84''$ and $4.43''$, therefore these corrections are applied to the raw timeline data, before the map-making procedure.

3.4 Individual Galaxy Considerations

For the full JINGLE survey there are seven pairs of JINGLE targets and one triple system, where the JINGLE SCUBA-2 observations of each individual target overlap. As the full sensitivity region of Daisy maps is quite small and the only other map modes are significantly bigger and limited to circular observations, we decided to observe these targets fully. These objects are processed together to make full use of the overlapping data, as the greater redundancy is useful for identifying and removing atmospheric and instrumental noise.

For three galaxies (JINGLE 45, 119, 150) we found partially overlapping observations in the archive from M13AN02 (Ivison et al. 2016) and M18AP013 (a JINGLE extension program), targeting other sources. For these galaxies combining the serendipitous observations with the JINGLE data, leads to a ~ 12 , 5, and 26% improvement in the instrumental noise for JINGLE 45, 119, 150, respectively.

We also inspected the SPIRE maps and our initial SCUBA-2 maps to see if there were other bright objects in the field that should be included in the input mask. The serendipitously detected bright high-redshift objects described in Paper I were also added to the mask and the maps re-run to ensure their flux was not suppressed.

To test our choice of filter correction we also ran a reduction of the maps where we varied the filter scales to a value of $175''$ for all galaxies and $200''$ for a selection of objects. For five galaxies we found a significant change in the flux measurement which were identified as cases in our $150''$ reduction where we find strong negative regions usually around a central source, often leading to a negative flux estimation. For JINGLE 45 we therefore use the map filtered at $175''$ and for JINGLE 19, 44, 154 and 185, we use maps filtered at $200''$.

The 193 SCUBA-2 maps described here, as well as the WISE, PACS and SPIRE cut-outs of our targets are available on the JINGLE data release page². We also provide SCUBA-2 maps that have been Gaussian smoothed with

² <http://www.star.ucl.ac.uk/JINGLE/data.html>

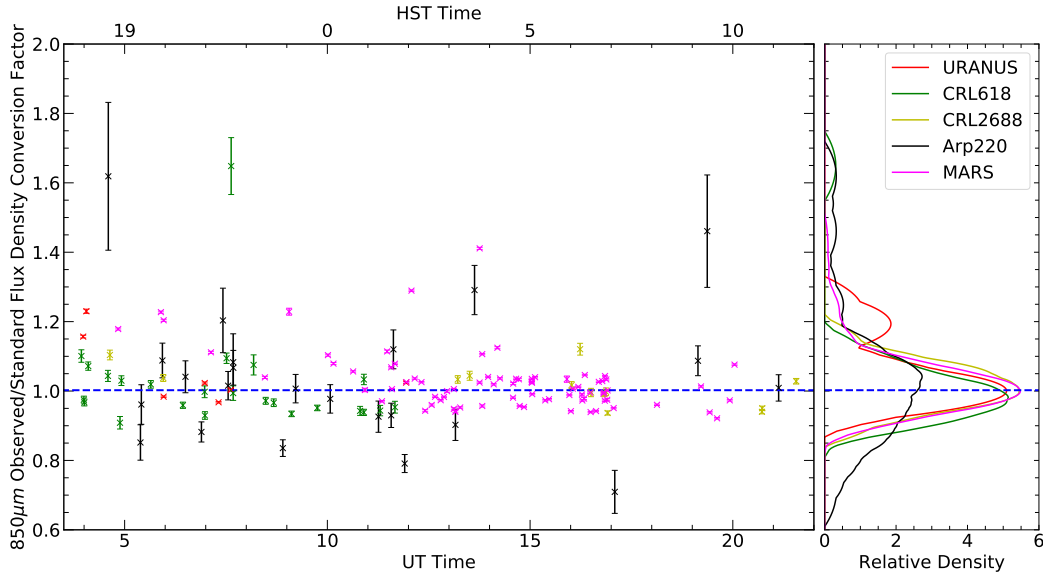


Figure 4. Flux conversion factors (FCFs) measured either before or after JINGLE observations at $850\mu\text{m}$ up to December 2016, calculated from observations of Mars (magenta), Uranus (red), CRL618 (green), CRL2688 (yellow) and Arp220 (black). The blue line shows the median value of all our observed FCF values which is in very good agreement with the standard FCF value ($1.002\times$). From the kernel density estimation (KDE) graph (right) we see that most of the calibration sources agree well with each other, although Arp220 being a dimmer source has a larger scatter. The data points are ordered by the time they were observed; we see no obvious trends.

either a FWHM of 12 and $24''$, and a map generated by applying the PICARD matched-filtered algorithm which creates a map optimised for the detection of point sources. The code used to perform the data reduction is available on GitHub³.

4 FAR-INFRARED/SUB-MM FLUX EXTRACTION

4.1 Photometry

Here we describe our method of deriving flux densities in the far-infrared and sub-mm ($22\text{--}850\mu\text{m}$). To extract flux densities from the *Herschel*, *WISE* and SCUBA-2 maps we perform aperture photometry on the dust images separately from the procedures used at shorter wavelengths (i.e., UV, optical) presented in Paper I as our dust images (particularly in the SCUBA-2 bands) have reduced signal-to-noise, due to the difficulties in observing in the far-infrared/sub-mm. Instead we decide to have a consistent set of apertures optimised on the dust data, rather than the larger apertures that would be derived from optical wavelengths. These flux-densities are therefore ideal for the analysis of the dust properties of galaxies (i.e., far-infrared dust SED fitting).

The size and fluxes from our dust apertures are generated by an automated code, but a large degree of manual customisation is allowed to remove problems that arise (for example other bright objects or map artefacts). To define our apertures we use the SPIRE $250\mu\text{m}$ band, as at this wavelength we have the greatest sensitivity to dust structures. For the first step an automatic mask is created to remove the influence of any other bright nearby galaxy. We

do this by performing a query on the RC3 catalogue (Corwin et al. 1994) and masking all pixels within a radius of $2.0R_{25}$, which will safely mask the vast majority of dust contamination (Smith et al. 2016). If other bright contaminating objects are found on the image their positions and sizes can be manually added to the mask. An initial noise estimate of the image is then calculated using an iterative sigma-clipping technique. This rough noise is then used to find all contiguous pixels above a given signal-to-noise (set to a default of 3) within a region with $R < 0.5R_{25}$. The smallest ellipse which fits these pixels is then used to define the centre, the axes ratio, and the Position Angle (PA) of the ellipse. If no significant pixels exist, or the shape can not be found, we default to the optical centre, the D_{25} axis ratio, and PA, from the RC3 catalogue. Using these parameters, and our rough noise estimates we then create a radial profile assuming a default background radius between 1.1 and $1.4R_{25}$. If the galaxy is above our detection criterion of a signal-to-noise of 3 then the object is preliminarily classed as a detection and the preliminary size of the aperture is determined by the radius at which our radial surface-brightness ratio crosses a signal-to-noise of 2, and multiplied by a default factor of 1.2 (this can be adjusted for special cases, i.e., to avoid contaminating objects) to ensure we enclose most of the emission.

To calculate the final aperture we need to make a better estimate of the noise in the image. To do this we use an estimate based on the method presented in Smith (2013) and Ciesla et al. (2012), where we consider the instrumental noise, confusion noise and background error separately. The instrumental noise is found using the ‘error’ extension provided with the maps, which are typically calculated based on the standard deviation of samples contributing to that pixels. As a full map is available the exact instrumental noise for a particular aperture is calculated by adding the pixels

³ <https://github.com/mwls/SCUBA2-public>

in the noise map in quadrature. As the confusion noise is not independent between pixels (as pixels are smaller than the beam), the confusion noise is calculated based on the square-root of the number of beams within the aperture. The third contribution is from large-scale backgrounds, including cirrus contamination. The method employed to measure the large-scale background by Ciesla et al. (2012) was very conservative and led to an overestimate of the uncertainty (see Smith 2013). We therefore modify this approach and instead use NEBULISER⁴, also used by the *H*-ATLAS team (Valiante et al. 2016), which filters the map based on a threshold radius into small-scale and large-scale. We set this threshold to 90'' in the SPIRE bands. We then use our large-scale map (or effectively cirrus map) and place a series of aperture across the image with the same size, shape and background region as our preliminary aperture on the cirrus map, and obtain the standard deviation of these apertures. Using NEBULISER in this way prevents us counting the same noise components multiple times. Ideally, these apertures would be randomly placed, but as some images are limited in size (especially in the SCUBA-2 bands), and we are primarily interested in local conditions, we use a grid of apertures around the target source (again avoiding other nearby galaxies). We scale this estimate for the relative number of pixels in our final aperture, which should be accurate assuming that the size of our preliminary aperture is similar to our final result.

As the point spread function (PSF) in far-infrared instruments can have significant extended features, we apply aperture corrections. These can be quite large factors, as even though the emission is low-surface brightness, the emission can be spread over a large area, and can contaminate background regions leading to an over-subtraction of the flux. As dust is expected to be distributed roughly in dust discs (Pohlen et al. 2010; Hunt et al. 2015; Smith et al. 2016), we fit an exponential plus constant model using the same PA and axis-ratio as the pixels within the aperture. In the fitting process each trial model is convolved with the radially averaged beam. By using our model within the aperture, we then predict the amount of emission outside the aperture, correcting our flux for the missing emission, and the effect of emission in the background region. For bands with lower resolution we use the the model parameters found from the 250 μm image and convolve the model with the appropriate beam (see Table 1 for whether the model was fit in each band).

While for the vast majority of objects in our sample the automated apertures are valid, a small subset (15) of objects require us to lock the centre to the optical centre. For 16 galaxies we lowered the default aperture expansion-factor of 1.2, although the factor is never set to be less than 1.0. These changes were often made to avoid problems from nearby background sources or other contaminating features.

In bands other than the SPIRE 250 μm , we use a matched aperture to the 250 μm band, except we correct for the difference in beam size by modifying the size of the semi-major and semi-minor axis. This is achieved by taking the 250 μm aperture size and subtracting in quadrature the

250 μm FWHM/2 and adding the current band's FWHM/2. Finally a useful diagnostic graphic is created showing the images with apertures, the resultant fluxes, peak signal-to-noise ratios, aperture corrections, the radial profile and growth curves. An example of this image is shown for JINGLE 122 in both SPIRE (Figure 5) and SCUBA-2 (Figure 6) wavebands.

Confusion measurements for SCUBA-2 are difficult as surveys of galaxies at higher redshifts apply matched-filtering to optimise the detection of point sources. This convolution changes the confusion noise measured on the map. To find the confusion noise we perform a simulation where we create a fake map that just contains confusion noise, and then apply the SCUBA-2 match filter. We vary the level of confusion until the amount of confusion noise in the raw map matches the 0.8 mJy beam⁻¹ measured by the SCUBA-2 Cosmology Legacy Survey (Geach et al. 2017). As part of our test we also check how the confusion noise scales with aperture size, we found that for SCUBA-2 assuming the contribution of confusion varying as $\sigma_{conf}\sqrt{N_{beam}}$ did not represent the noise adequately, instead we use $\sigma_{conf}\sqrt{N_{beam}-const}$ where σ_{conf} is 1.22 mJy beam⁻¹ and *const* is -0.61. Table 1 lists the beam area, confusion noise, and nebuliser filter scale assumed for the flux density measurement in each band.

Both JINGLE 57 and 60 at 850 μm were found to have significant contamination within the aperture from a point source presumably at a higher redshift. To remove the contamination we fit the point source using the same procedure outlined in Section 4.2, and then subtract the source from the image. We then run the aperture code as normal on the subtracted image. A more complicated case is JINGLE 29 where the more compact early-type object overlaps with a more flocculent late-type galaxy. To separate these objects we use the typical size of dust-disks based on the D_{25} from Smith et al. (2012a) to choose a region that likely encompasses the extent of emission from JINGLE 29, this also agreed with the region that would be chosen visually. Then all pixels within that region contaminated by the late-type galaxy were replaced with a sigma-clipped mean of the map, and the resulting image then used to perform the usual aperture flux extraction method.

4.2 Point Source Extraction

Several of our sources are compact enough they are well described by a point source in the far-infrared wavelengths from PACS to SCUBA-2. By identifying point sources we can use extraction techniques optimised to recover flux densities using the PSF as a prior, rather than relying on aperture photometry. This improves both the flux estimate and the significance of the detection (particularly in the SCUBA-2 bands). For SPIRE to fit a point source the images are firstly converted into units of Jy beam⁻¹ and then NEBULISER is run over the maps to remove any emission on scales greater than 3'. We then fit a radially averaged PSF in each band, using a model grid based on five times smaller pixels than the original pixel size to minimise the effect of pixelisation. For the SPIRE 250 μm we allow the central position to vary by up to 6'' compared to the optical centre, while for other bands the central position is locked to the SPIRE 250 μm position. The largest offset was found to be 3.3'' with

⁴ <http://casu.ast.cam.ac.uk/surveys-projects/software-release/background-filtering>

JINGLE122

J131313.45+335903.9

Detected

Peak S/N: 41.8

(350 μ m:21.8, 500 μ m:7.8)

Flux Densities:

250 μ m
0.897 +/- 0.033 Jy

350 μ m
0.399 +/- 0.023 Jy

500 μ m
0.149 +/- 0.020 Jy

Aperture Correction

Factors:

PSW: 19%

PMW: 22%

PLW: 36%

- Vertical Lines
- S/N Rad Threshold
- Aperture Radius
- Background Region

- Image Lines
- Optical Centre
- Background Region
- Exclusion Regions
- Beam

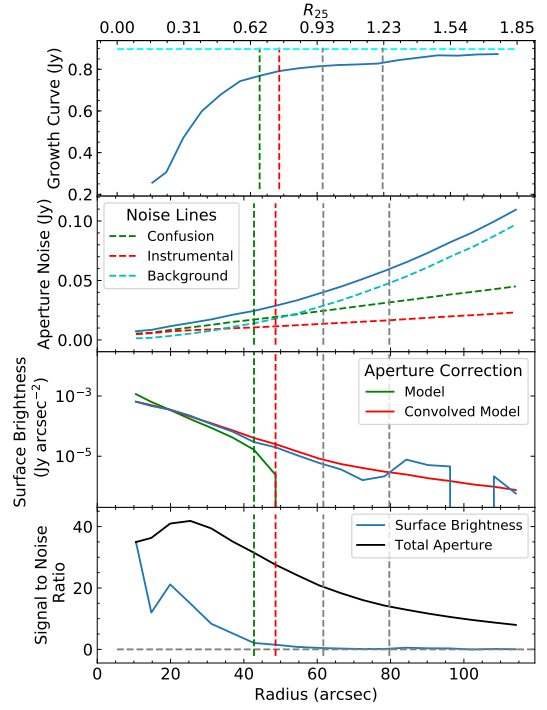
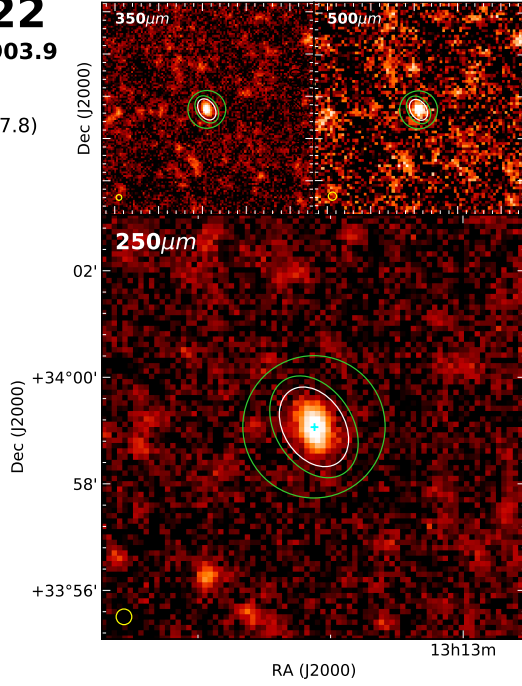


Figure 5. The diagnostic image produced by the photometry pipeline for JINGLE 122 for the SPIRE wavelengths. The images in the centre show the aperture (white) and the background region used (green) on the images, as well as the optical centre (cyan cross) and FWHM beam size in yellow. If objects were to be excluded from the analysis they would be shown in blue. On the right of the graphic we have various curves which vary with radius; the top plot shows the growth curve (horizontal cyan line is the flux measurement), the second from top plot shows how the aperture noise grows as function of the different noise sources, third from top is the surface brightness profile (blue) with the exponential aperture correction models shown (green model, red convolved with beam), and the bottom plot shows the signal-to-noise of the surface brightness profile (blue) and total aperture (black). On all the radial plots the vertical green line shows where the surface brightness profile crosses a signal-to-noise of 2 and the red line is the radius of the aperture. The vertical grey lines show the radii of the background region. Finally on the left the peak signal-to-noise, fluxes, and aperture corrections are listed.

JINGLE122

J131313.45+335903.9

SPIRE Detection

SCUBA Detection

Peak S/N 850 μ m:7.2

SPIRE Matched Fluxes:

850 μ m
0.027 +/- 0.008 Jy

450 μ m
No Data Available

Ap-Correction Factors

850: 7%

450: No Data

Filter Correction Factors

850: 5.6 +/- 0.3%

450: No Data

- Vertical Lines
- S/N Rad Threshold
- Aperture Radius
- Background Region

- Image Lines
- Optical Centre
- Background Region
- Exclusion Regions
- Beam

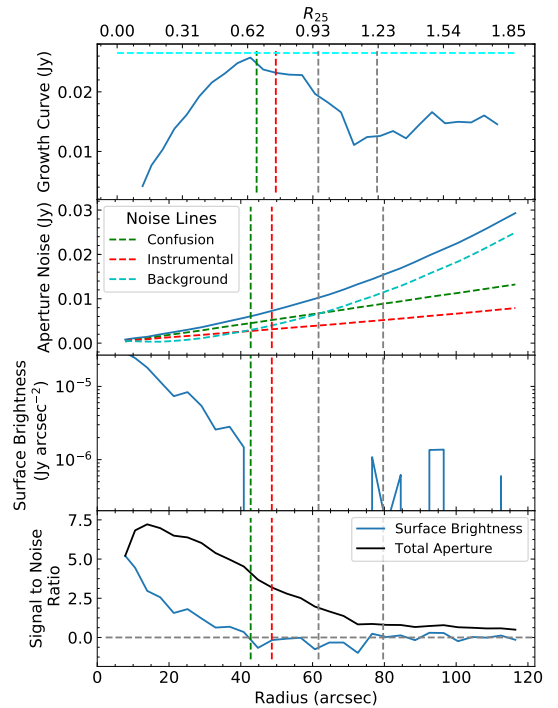
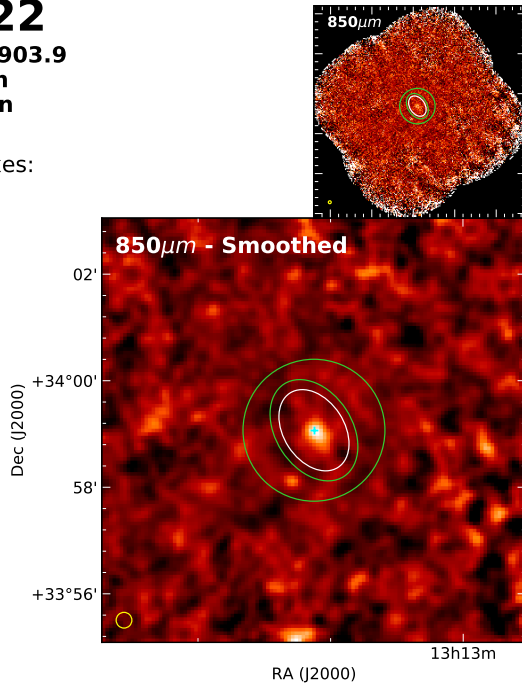


Figure 6. Same as Figure 5, except now applied to SCUBA-2. Unlike the SPIRE 250 μ m figure the aperture is matched to the SPIRE band (rather than derived from the image).

Table 1. Individual Band Parameters

Instrument	Band	FWHM ($''$)	Beam Area (arcsecond ²)	Confusion Noise (mJy beam ⁻¹)	Nebuliser Median Filter ($''$)
SPIRE	250 μm	17.6	469.7	5.8	90
	350 μm	23.9	831.7	6.3	90
	500 μm	35.2	1793.5	6.8	90
PACS	100 μm	11.4	147.2 ^a	0.27	60
	160 μm	13.7	212.7 ^a	0.92	60
	3.4 μm	6.1	42.2 ^a	-	60
WISE	4.6 μm	6.4	46.4 ^a	-	60
	12 μm	6.5	47.9 ^a	-	60
	22 μm	12.0	163.2 ^a	-	60
SCUBA-2	850 μm	13.0	229.5	($m = 1.22$ $c = -0.61$) ^b	60

Notes. For the WISE wavelengths we assume the contribution of confusion noise is negligible and so set to zero in the code. Confusion noise estimates for SPIRE were taken from [Nguyen et al. \(2010\)](#), with the PACS FWHM, beam areas and confusion noise taken from [Smith et al. \(2017\)](#).

^a These beam areas assume a Gaussian beam profile calculated from the FWHM. These instruments are not calibrated in Jy beam⁻¹ (or similar), but the conversion is used just to scale the confusion noise (where appropriate).

^b See Section 4.1 for details on the SCUBA-2 confusion noise.

most offsets substantially smaller. For the SPIRE bands the uncertainty in the flux is measured using the uncertainties measured in the fit combined with the confusion noise (as given in Table 1). An example of the point source extraction applied to a SPIRE image is shown in Figure 7.

To decide whether our candidate galaxies can be reasonably approximated by point sources as well as a visual inspection, we run the residual images through our aperture photometry code as described in Section 4.1. If the residual signal in the aperture at 250 μm has a signal-to-noise < 2.5 then we use the point source procedure. From our sample of 193 galaxies, 42 meet our criteria to be considered point-like.

For the SCUBA-2 images we first apply the ‘matched-filter’ ORAC-DR recipe which optimises SCUBA-2 maps for point source extraction. Effectively the recipe subtracts a Gaussian smoothed version of the map and convolves the map with the SCUBA-2 PSF. We then fit the matched-filter PSF characterised by [Geach et al. \(2017\)](#), with the source centre fixed to the SPIRE position. However, as the matched-filtering convolves the image the noise of each pixel can no longer be considered independent. We therefore perform a Monte Carlo simulation where for every pixel in the raw map we scale a normally distributed random number to account for the instrumental noise in the pixel and add it to the map. The same matched-filtering and point source extraction is performed to the new map and the whole process repeated 100 times. The uncertainty in the flux is then estimated based on the distribution of extracted fluxes. Note, that despite the slightly higher angular resolution of the JCMT at 850 μm than SPIRE at 250 μm , with the smoothing action of the matched-filter these sources should be well represented by a point source).

Three candidate point-galaxies that narrowly failed our criteria (JINGLE 162, 182 and 192) at SCUBA-2 850 μm , were found to have higher flux measurements when using the point source measurement, compared to the aperture photometry method, due to noise in the aperture. For these three galaxies we therefore report the point source flux (this

again is helped by the smoothing action of the matched-filter).

The *H*-ATLAS team recommends for PACS data that apertures of radius 11.4 and 13.7 $''$ are used for the extraction of point sources ([Valiante et al. 2016](#); [Smith et al. 2017](#)), due to variations of the beam shape. As the WISE data is of similar or better resolution we use the same 11.4 $''$ radius for the W1–W3 bands and 13.7 $''$ radius aperture for the W4 band. Whether the point source extraction method is used in either the SPIRE or the SCUBA-2 850 μm bands is specified in the flux-density catalogue table available in the supplementary materials or on the JINGLE data release page².

4.3 Simulations

As mentioned in Section 3.1, the SCUBA-2 pipeline relies on high-pass filtering, and while we have tried to minimise the effect when creating a map there could be some residual flux attenuation. To try to correct for this effect for each object we run a simulation where we inject a model of the source into the raw instrumental data of JINGLE observations that appear to have no signal (although to avoid any bias we avoid the location of the target galaxy). From our observations observed up to December 2016 we identified 12 target regions that could be used for this purpose. The raw data with the injected source is then processed with the same map-making procedure as for the real data (including same mask size and configuration). To minimise the effect of the noise on our correction, we also re-run the map-making but with no model source injected using the same mask. This blank map is then subtracted from the map with the injected source and estimates of both the aperture and filter corrections are calculated.

To make the model that is injected into the raw data we take the exponential models used to find the aperture corrections in the SPIRE data (see Section 4.1), these are then convolved with the SCUBA-2 beam and the flux adjusted to

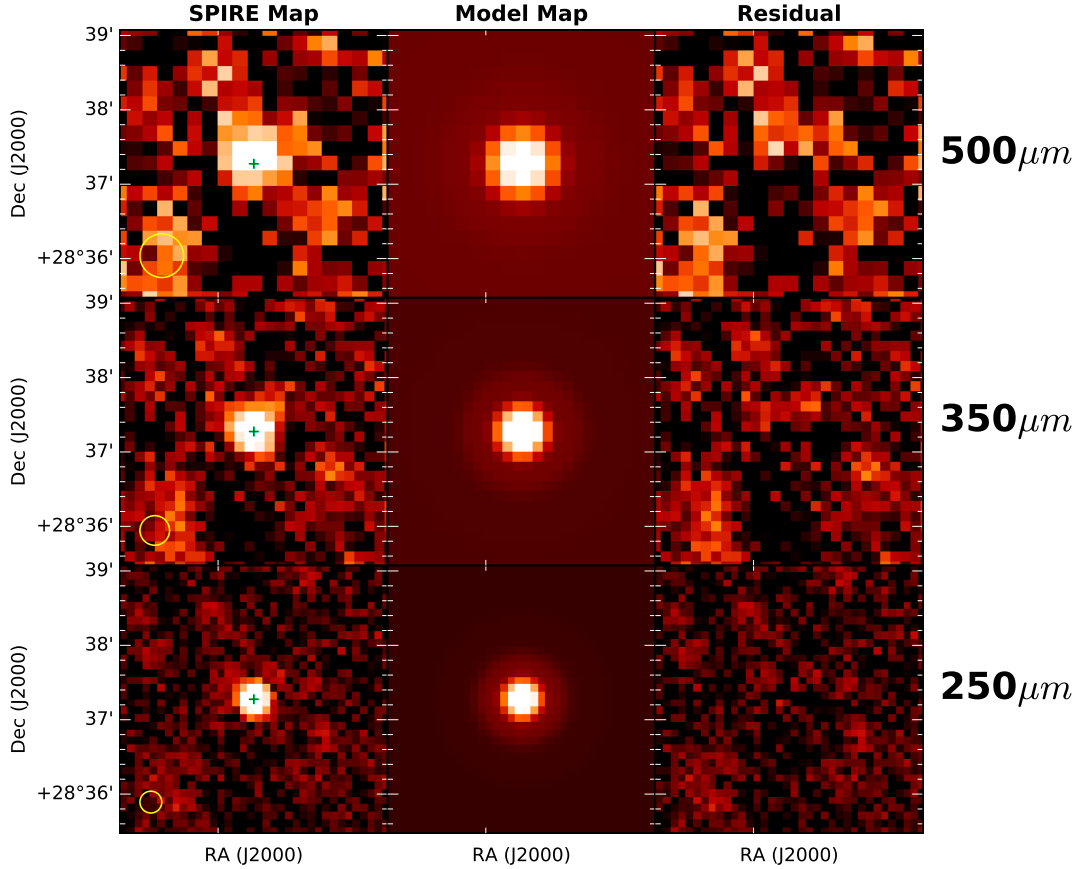


Figure 7. An example of the point source fitting method for JINGLE 33 in the SPIRE wavelengths. The grid of images show the SPIRE maps (left column), the point source model images (middle column), and the residual images after the point source is subtracted (right column) for the 250, 350 and 500 μm bands. On the SPIRE maps the optical centre is shown by the cyan cross, the fitted centre by the green cross (in this example the two overlap so only the green cross is visible) and the beam FWHM in yellow.

match our predicted 850 μm flux. The model source is then converted into pW so it matches the units of the instrumental timelines. For each JINGLE target we inject the galaxy into three different blank fields chosen to have the closest sensitivity to our target observation (estimated in the central region of the instrumental uncertainty map). The model source to be injected is adjusted so the signal-to-noise ratio is the same for the injected source as expected in the real map (from the predicted flux-density). The centre and orientation of the injected source is chosen randomly, but the distance from the centre of the image is limited to be within 2' so we are not affected by the noisy regions at the edges of the map. The random location is not allowed to overlap with the intended target of the ‘blank’ data. The size of the mask used is identical to that applied to the reduction of the target observation.

For galaxies that we assume are point-like the correction factor (this correction also accounts for any systematic from the matched-filtering and flux extraction process we have not accounted for) has an average value of 1.135 with a narrow distribution (minimum factor 1.121 and maximum 1.158). For extended galaxies the average filter correction is lower at 1.045, but has a much larger range from 1.006 to 1.139. These corrections have been applied to all fluxes provided in the catalogue.

4.4 Catalogue Statistics

At 250 μm 191 of the 193 galaxies in the JINGLE sample are easily detected ($> 5\sigma$) by SPIRE with most objects having peak signal-to-noise ratios between 10–50 (the high fraction is not surprising as galaxies were included in our sample based on a preliminary *H*-ATLAS catalogue). The two notable exceptions are JINGLE 62 and JINGLE 130, neither of which are listed in the release version of the *H*-ATLAS DR2 catalogue (Maddox et al. 2018). From our measurements JINGLE 130 is identified with a peak signal-to-noise ratio of 4.5 and is identified as a point-like source. We therefore follow the same procedures as galaxies detected with signal-to-noise ratios > 5 . JINGLE 62 is fainter than JINGLE 130 and has a significance far below 3.0σ , we therefore just report an upper-limit to the flux for this source across all the wavebands, using an elliptical aperture with semi-major axis equal to R_{25} (this size was chosen to match findings of the HRS, Ciesla et al. 2012; Smith et al. 2012a).

Table 2 provides an indication of our detection rates across the different bands, by giving the number of objects with signal-to-noise ratios greater than 5 or 3. The table is split by peak signal-to-noise ratio, which is the aperture that gives the highest signal-to-noise, while the total signal-to-noise is the measurement on the total aperture. The peak

Table 2. Sample Detection Statistics

Instrument	Band	Peak S/N		Total S/N	
		$\geq 5\sigma$	$\geq 3\sigma$	$\geq 5\sigma$	$\geq 3\sigma$
SPIRE	250 μm	191/193	192/193	191/193	192/193
	350 μm	186/193	190/193	185/193	190/193
	500 μm	93/193	138/193	79/193	132/193
PACS	100 μm	177/190	181/190	157/190	174/190
	160 μm	182/190	187/190	177/190	185/190
WISE	3.4 μm	192/193	192/193	189/193	189/193
	4.6 μm	192/193	192/193	189/193	189/193
	12 μm	192/193	192/193	192/193	192/193
	22 μm	188/193	192/193	186/193	188/193
SCUBA-2	850 μm	62/193	126/193	30/193	83/193

Notes. The total number of objects in the PACS sample is 190 instead of the 193 for SPIRE, WISE and SCUBA-2, as 3 galaxies were not covered by PACS imaging due to the 22' offset between the PACS and SPIRE cameras on the *Herschel* focal plane.

measurement is more analogous to a detection statistic while the total aperture shows the ability to measure the total flux of an object (for point sources the peak and total S/N are the same). As outlined in Section 4 we use the SPIRE 250 μm data to define our apertures, so that we get aperture matched fluxes (or point source estimates) across all wavebands even if the signal-to-noise ratio is significantly lower in other bands.

Table 2 shows that we have secure WISE and SPIRE 250/350 μm detections for almost every object in our sample and over 82% with PACS coverage (with total S/N > 5). At SPIRE 500 μm as we move down the Rayleigh-Jeans tail the number of galaxies detected drops to 48% at 5σ , and 41% have total fluxes with signal-to-noise ratios greater than 5. The JINGLE SCUBA-2 850 μm of course is more challenging due to the intrinsically fainter emission of local galaxies in this waveband and the effects of the atmosphere, but we do detect 126 of our 193 (65%) at a level of $> 3\sigma$, and have 83 (43%) galaxies where the total signal-to-noise of the flux measurement is greater than 3. However, galaxies with lower signal-to-noise can still be useful to constrain the dust properties of an object. The distribution of the peak and total aperture signal-to-noise measurements for our SCUBA-2 sample is shown in Figure 8.

The complete table of flux-densities in the WISE, PACS, SPIRE and the SCUBA-2 850 μm bands is available in the supplementary materials and on the JINGLE data release page². The python scripts used to perform the photometry are available online⁵.

5 CO(3-2) CONTAMINATION

The CO($J=3-2$) emission line at 867.6 μm (354.796 GHz) falls within the bandpass filter for the 850 μm SCUBA-

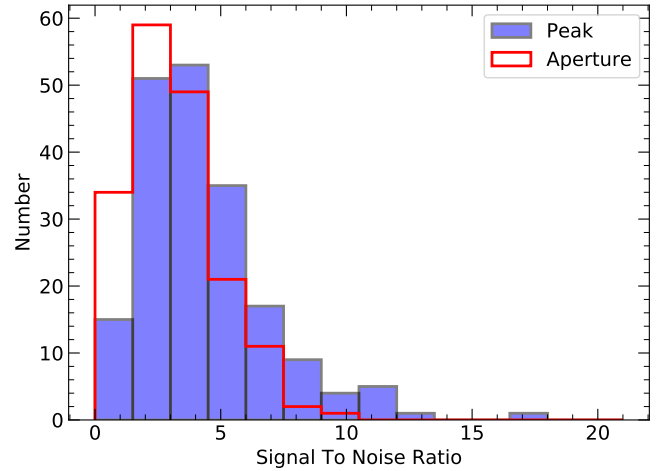


Figure 8. A histogram of the signal-to-noise ratio values for the flux measurements of JINGLE galaxies at the SCUBA-2 850 μm wavelength. The blue filled histogram is the peak signal-to-noise (see text) while the open red histogram is the signal-to-noise of the SPIRE matched aperture.

2 band for all the galaxies in the JINGLE sample ($z < 0.05$); leading to potentially significant contamination of the continuum flux (Drabek et al. 2012). In order to correct for this contamination, we need to estimate the strength of the CO($J=3-2$) emission line. For the JINGLE galaxies we do not have CO($J=3-2$) observations, but we have predictions of the CO($J=1-0$) luminosity from Xiao et al. in prep (Paper III). In this paper, they show that these predictions are in good agreement with the CO line luminosities measured for a sub-sample of JINGLE galaxies. The CO($J=3-2$) line luminosity can be estimated from the CO($J=1-0$) transition, assuming an excitation line ratio $r_{31} = L'_{\text{CO}(J=3-2)} / L'_{\text{CO}(J=1-0)}$.

To estimate the ratio r_{31} for the JINGLE galaxies, we studied a sample of ~ 20 galaxies from the COLD GASS survey (Saintonge et al. 2011) which allowed us to measure r_{31} in a redshift range similar to that of the JINGLE survey. The galaxies of this sample also have stellar masses, specific star-formation rates (SSFR) and star-formation efficiencies (SFE = SFR / $M(\text{H}_2)$) comparable to the JINGLE sample. We found a mean value of $r_{31} = 0.53 \pm 0.05$. This value is consistent with observations of low-redshift galaxies (Yao et al. 2003; Mao et al. 2010; Papadopoulos et al. 2012, average $r_{31} = 0.6-0.7$) and intermediate redshift ($z \sim 0.3$) star forming galaxies (Bauermeister et al. 2013, average $r_{31} = 0.46 \pm 0.07$). Therefore we decided to use the constant value $r_{31} = 0.5$.

We converted the obtained line luminosity $L'_{\text{CO}(J=3-2)}$ into line intensity $I_{\text{CO}(J=3-2)} = \int T_{\text{MB}} dv$ in $[\text{K km s}^{-1}]$ using eq.(2) from Solomon et al. (1997):

$$I_{\text{CO}(J=3-2)} = \frac{L'_{\text{CO}(J=3-2)}}{23.5 \cdot \Omega_b D_L^2 (1+z)^{-3}} \quad (1)$$

where $L'_{\text{CO}(J=3-2)}$ is in units of $[\text{K km s}^{-1} \text{ pc}^2]$, Ω_b is the telescope beam area in arcsec^2 and D_L is the luminosity distance in Mpc. We assume that the sizes of our galaxies

⁵ <https://github.com/mwls/Public-Scripts>

are comparable to the size of the SCUBA-2 beam ($13.5''$ FWHM), and use $\Omega_b = \pi(13.5/2)^2 \text{ arcsec}^2$.

To convert the line intensity from line units [K km s^{-1}] to SCUBA-2 $850 \mu\text{m}$ continuum flux, we used the conversion factor C defined by Drabek et al. (2012) as:

$$C = \frac{F_\nu}{\int T_{\text{MB}} d\nu} = \frac{2k\nu^3}{c^3} \frac{g(\nu)}{\int g(\nu) d\nu} \Omega_B \quad (2)$$

where C has the units $\text{mJy beam}^{-1} \text{ per K km s}^{-1}$. F_ν is the line flux, $\int T_{\text{MB}} d\nu$ is the integrated main-beam temperature, k is the Boltzmann constant, $g(\nu)$ is the transmission at the frequency ν of the $850 \mu\text{m}$ filter multiplied by the atmospheric transmission, and Ω_B is the telescope beam area.

The C factor that converts the line intensity into SCUBA-2 continuum flux depends on the atmospheric transmission. We used the updated prescription for the C conversion factor from Parsons et al. (2018), which provides the C factor as a function of the precipitable water vapour (PWV). The main differences from the prescription by Drabek et al. (2012) are that they used an updated main-beam FWHM of $13.0''$ with a relative amplitude of 0.98, and that they include a correction for the secondary beam component, which has a relative amplitude of 0.02 (Dempsey et al. 2013). The C factor as a function of PWV is defined as:

$$C = 0.574 + 0.1151 \cdot \text{PWV} - 0.0485 \cdot \text{PWV}^2 + 0.0109 \cdot \text{PWV}^3 - 0.000856 \cdot \text{PWV}^4 \text{ mJy beam}^{-1} / \text{K km s}^{-1}, \quad (3)$$

where the PWV is related to the sky opacity at 225 GHz ($\tau_{225 \text{ GHz}}$) as $\text{PWV} = (\tau_{225 \text{ GHz}} - 0.017)/0.04 \text{ mm}$. The C factor for our sample varies in the range $0.63 - 0.75 \text{ [mJy beam}^{-1} \text{ per K km s}^{-1}]$. The C factors are calculated at the rest-frame frequency of $\text{CO}(J=3-2)$. To account for the different transmission at the observed frequency of $\text{CO}(J=3-2)$, we applied the following correction:

$$C_{\text{obs}} = \frac{1}{(1+z)^3} \cdot \frac{g(\nu_{\text{obs}})}{g(\nu_{\text{rest}})} \cdot C \quad (4)$$

where ν_{rest} is the $\text{CO}(J=3-2)$ rest-frame frequency and $\nu_{\text{obs}} = \nu_{\text{rest}}/(1+z)$ is the $\text{CO}(J=3-2)$ observed frequency. For each JINGLE observation we average the $\tau_{225 \text{ GHz}}$ values for each observation and then calculate the correction.

Thus, the $\text{CO}(J=3-2)$ flux contamination in [mJy beam^{-1}] is:

$$F_{\text{CO}(3-2)} = \int T_{\text{MB}} d\nu \cdot C_{\text{obs}} \quad (5)$$

The $\text{CO}(J=3-2)$ flux contamination is in the range $0.7-41.2\%$ of our predicted $850 \mu\text{m}$ values with a mean contamination of 10.1% . For the vast majority of sources (78%) the correction is less than 15% of the predicted $850 \mu\text{m}$ flux density, and 91% of sources have corrections less than 20% of the predicted $850 \mu\text{m}$ flux density. The SCUBA-2 flux estimated contamination is provided in the main catalogue table available in the Supplementary Materials. After correcting for the $\text{CO}(J=3-2)$ flux contamination 73 of our targets have aperture fluxes with a signal-to-noise greater than 3.

6 RESULTS & DISCUSSION

6.1 Comparison of the Predicted and Observed SCUBA-2 fluxes

In this section, we compare the observed SCUBA-2 $850 \mu\text{m}$ fluxes with predictions from fitting the far-infrared SED using our *Herschel* data in the wavelength range $100-500 \mu\text{m}$. We used a single modified black-body model (SMBB) defined as (Hildebrand 1983):

$$F_\lambda = \frac{M_{\text{dust}}}{D^2} \kappa_\lambda B_\lambda(T_{\text{dust}}) \quad (6)$$

where M_{dust} is the dust mass in the galaxy, D is the distance of the galaxy and $B_\lambda(T_{\text{dust}})$ is the Planck function for the emission of a black-body with a dust temperature T_{dust} . The dust mass absorption coefficient κ varies as a function of wavelength:

$$\kappa_\lambda = \kappa_0 \left(\frac{\lambda_0}{\lambda} \right)^\beta \quad (7)$$

where κ_0 is the reference dust mass absorption coefficient. We use a constant value $\kappa_0 = 0.051 \text{ m}^2 \text{ kg}^{-1}$ at $500 \mu\text{m}$ from Clark et al. (2016). We assume a fixed emissivity index $\beta = 2$ to have conservative estimates of the flux, since β is typically in the range $1.5-2.0$ (e.g. Boselli et al. 2012; Galametz et al. 2012; Smith et al. 2012b; Clemens et al. 2013; Cortese et al. 2014). We fit the model to the *Herschel* fluxes ($100, 160, 250, 350$ and $500 \mu\text{m}$), using a Bayesian non-hierarchical method implemented using the MCMC code *emcee* (Goodman & Weare 2010; Foreman-Mackey et al. 2013), and assuming that the noise is normally distributed. For details about the fitting procedure see Lamperti et al, in prep. The predicted spectra in the $850 \mu\text{m}$ band were convolved with the filter transmission curve of the SCUBA-2 band, before comparing them with the observed fluxes.

Figure 9 shows the comparison between the predicted and observed $850 \mu\text{m}$ fluxes. As the JINGLE observations were individually tuned to reach a fixed detection threshold (rather than a fixed sensitivity), galaxies with a higher measured flux-density than predicted tend to be detections, while those with lower than predicted flux density do not reach our detection criteria. While the kernel-density estimator (KDE) graph (right panel Figure 9), peaks in agreement with our predictions, generally there is a large scatter. For the majority of objects ($\sim 75\%$) the scatter could be explained simply due to the uncertainty in the flux measurement. However, the distribution is wider than you would expect from the uncertainties (both above and below our predictions). More galaxies (110 versus 82) show an excess of the observed fluxes (i.e., they are higher than the predictions), meaning that for those galaxies the slope of the SED in the sub-millimeter is shallower. The most likely explanation is that their dust SED should be modelled with a lower β value, either due to the intrinsic dust properties or from multiple dust temperature components along the line of sight. Several studies have found β varying between $0.5-2.5$ on global scales (e.g., Boselli et al. 2012; Galametz et al. 2012; Clemens et al. 2013; Cortese et al. 2014) and within a galaxy (Smith et al. 2012b; Tabatabaei et al. 2014; Kirkpatrick et al. 2014). We will investigate this further in subsequent sections.

There are also galaxies for which the observed flux is

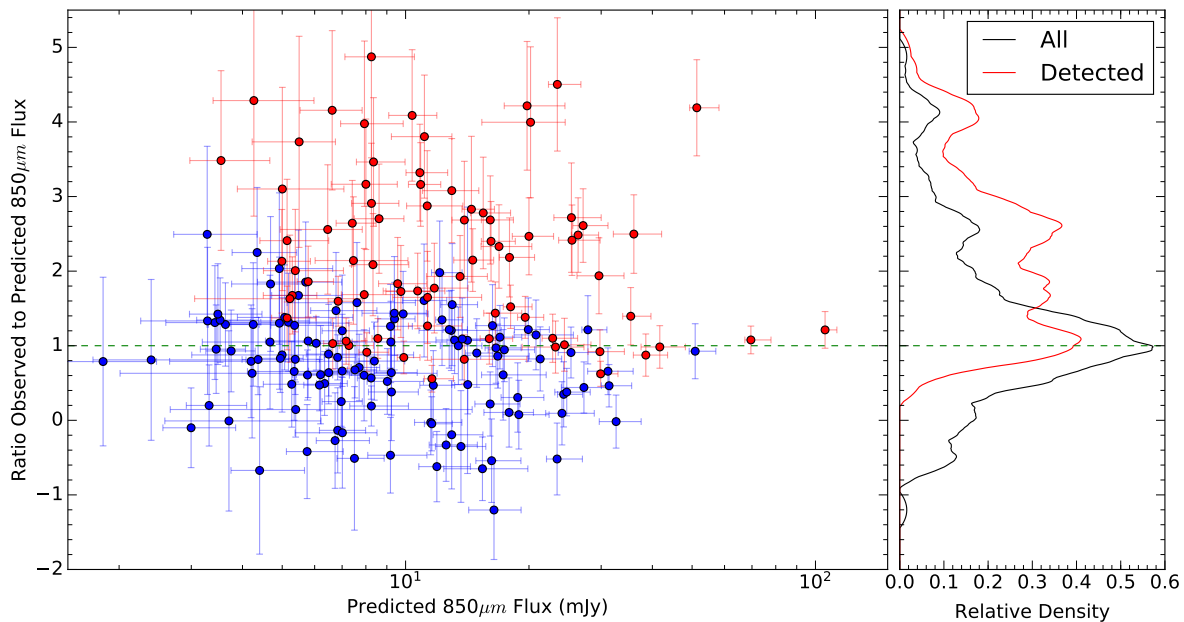


Figure 9. The left panel shows the ratio of the observed flux to the predicted $850\ \mu\text{m}$ flux versus the predicted flux. The $850\ \mu\text{m}$ flux predictions are based on a fit of a single modified black-body to the *Herschel* bands with a fixed emissivity index $\beta = 2$. Objects that have measured fluxes with signal-to-noise ratios greater than 3 are shown by the red points, and all other objects in blue. The horizontal green dashed line is where the ratio of observed to predicted $850\ \mu\text{m}$ flux is 1. The right panel shows the KDE of the ratio of observed to predicted $850\ \mu\text{m}$ flux for all objects (black line) and those with measured flux ratios with $S/N > 3$ (red line).

lower than the predicted flux. Our predictions could overestimate the true $850\ \mu\text{m}$ flux density if the dust emissivity index could be greater than 2 (e.g., Lis et al. 1998; Meny et al. 2007; Smith et al. 2012b), or if we have overestimated the $\text{CO}(J=3-2)$ contribution for these galaxies. For JINGLE we estimated the $\text{CO}(J=3-2)$ contribution from the predicted $\text{CO}(J=1-0)$ luminosity, using a constant line excitation correction factor $r_{31} = L_{\text{CO}(3-2)}/L_{\text{CO}(1-0)} = 0.5$, but this factor is known to vary in the range 0.1–1.9 (Mauersberger et al. 1999; Yao et al. 2003; Mao et al. 2010). Assuming an extreme value of $r_{31} = 0.1$, the $\text{CO}(J=3-2)$ contribution will be a factor of five lower. The estimated CO contribution is lower than 30% of the observed flux for 95% of our sample. Thus an overestimation of r_{31} can only account for a deficit of $\sim 25\%$ of the observed flux with respect to the predictions. As shown in Figure 9, the observed deficit can be as large as the predicted flux density, therefore the $\text{CO}(J=3-2)$ contribution is not sufficient to explain the deficit.

6.2 Far-infrared Colours

In this section we investigate how far-infrared/sub-mm colours vary with each other and compare with the predictions of modified blackbody models. We define a FIR/sub-mm colour as the ratio of two flux-densities at wavelengths in the range $100\text{--}850\ \mu\text{m}$. We investigate a selection of colour ratios that sample different regions of the far-infrared spectrum. For example the ratio at smaller wavelengths like F_{100}/F_{160} , and F_{100}/F_{250} should be more sensitive to temperature variations of the dust. Colour-ratios including flux densities at longer wavelengths F_{160}/F_{500} , F_{250}/F_{500} , and

F_{160}/F_{850} should be more sensitive to changes in the dust emissivity index (β).

Figure 10 shows a grid of the FIR/sub-mm colour ratios (similar to those derived for the HRS in Boselli et al. 2012 and Cortese et al. 2014), the best correlations between FIR colours is seen for indicators that trace the dust emissivity index (i.e., longer wavelengths). The highest correlation is seen between the F_{250}/F_{850} and F_{160}/F_{850} , with the Spearman rank coefficient (ρ) equal to 0.96. This shows the advantage of having the longer $850\ \mu\text{m}$ SCUBA-2 data to constrain the Rayleigh-Jeans tail, as using only the SPIRE wavelengths $\rho = 0.80$, for the F_{250}/F_{500} and F_{250}/F_{350} colour ratio. The two shorter colour-ratios (F_{100}/F_{250} and F_{100}/F_{160}) also show a very good correlation ($\rho = 0.77$), probably as both are indicative of dust temperature.

Similar correlations to those shown in Figure 10 were found by Boselli et al. (2012), for the HRS (Boselli et al. 2010a), with their best correlation ($\rho = 0.98$) between F_{250}/F_{500} and F_{250}/F_{350} . Overall, the *Herschel* Reference Survey predominately had stronger correlations than we find for JINGLE, although this could be explained due to the HRS targeting local galaxies with significantly higher global flux estimates, and deeper observations than *H-ATLAS*.

In Figure 11 we show the four colour-ratios from Figure 10 with the highest Spearman’s rank correlation, including the uncertainties for each data point. For the colour-ratios not including $850\ \mu\text{m}$ we over-plot objects from the HRS in grey (Boselli et al. 2012; Cortese et al. 2014). Given the large difference in selection (the HRS is a local *K*-band selected sample) compared to JINGLE, the distribution of the detected HRS galaxies in the colour ratio plots appears very similar to that of JINGLE, suggesting the dust properties of both samples is broadly similar. This is possibly due

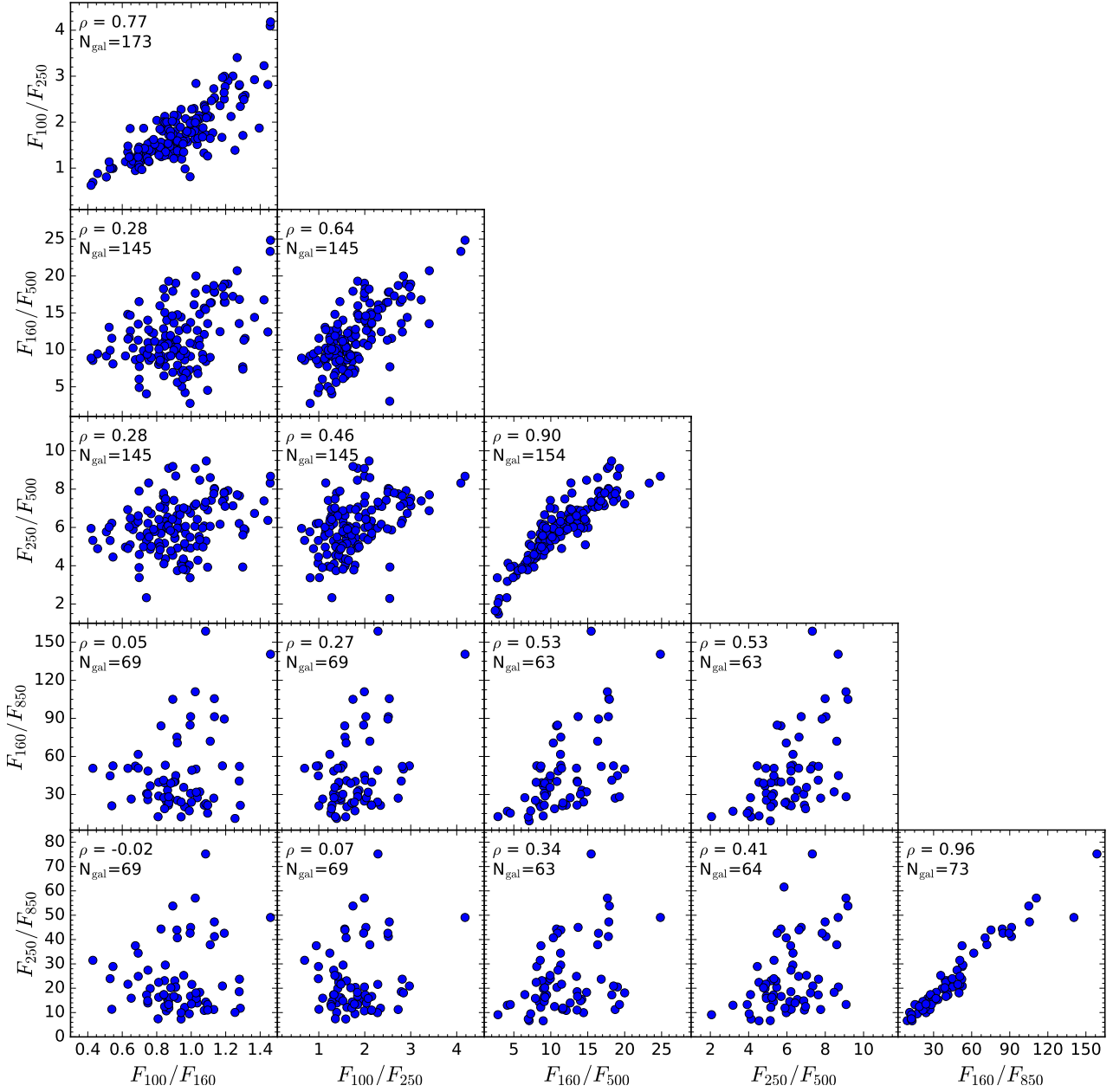


Figure 10. Far-infrared/sub-mm colour-colour plots for selected colours, where each colour is the ratio of two flux-densities in the wavelength range 100–850 μm . Galaxies are only included if each flux-density in the panel has a signal-to-noise ratio greater than 3. The Spearman’s rank correlation coefficient (ρ), and the number of galaxies in each panel is given in the top left corner of each panel.

to the similar stellar masses of the samples, although the specific SFR is quite different (see JINGLE Paper IV, de Looze et al. *in prep*). The red, green and orange lines in Figure 11 show the line produced for a single modified blackbody with a $\beta = 2$, $\beta = 1.5$, and $\beta = 1$ over a temperature range of 15–30 K respectively. The distribution of points in the top-left panel is dominated by the temperature of the dust as shown by the similarity of the two blackbody lines. The blackbody lines show that a temperature range of 15–30 K is sufficient to explain the vast majority of JINGLE and HRS objects (in agreement with JINGLE Paper V, Lamperti et al. *in prep*). The top-right and bottom left panels show the difficulty in making estimates of the dust emissivity index from

Herschel data alone where both modified blackbody models with $\beta = 1.5$ and 2 lie within the scatter of the data for the sample. The bottom-right panel which shows the F_{250}/F_{850} versus F_{160}/F_{850} shows the extra discrimination between models that the longer wavelength data provides, with the objects at high F_{250}/F_{850} and F_{160}/F_{850} appearing to lie on the $\beta = 2$ line, while surprisingly the majority of objects appear to fall below the $\beta = 1$ line.

To test the significance of our colour-plots to distinguish between the dust models, for each data point we calculated the smallest ‘normalised’ distance (i.e., the smallest distance between the model divided by the uncertainties). If the model is an accurate description of the data, the distri-

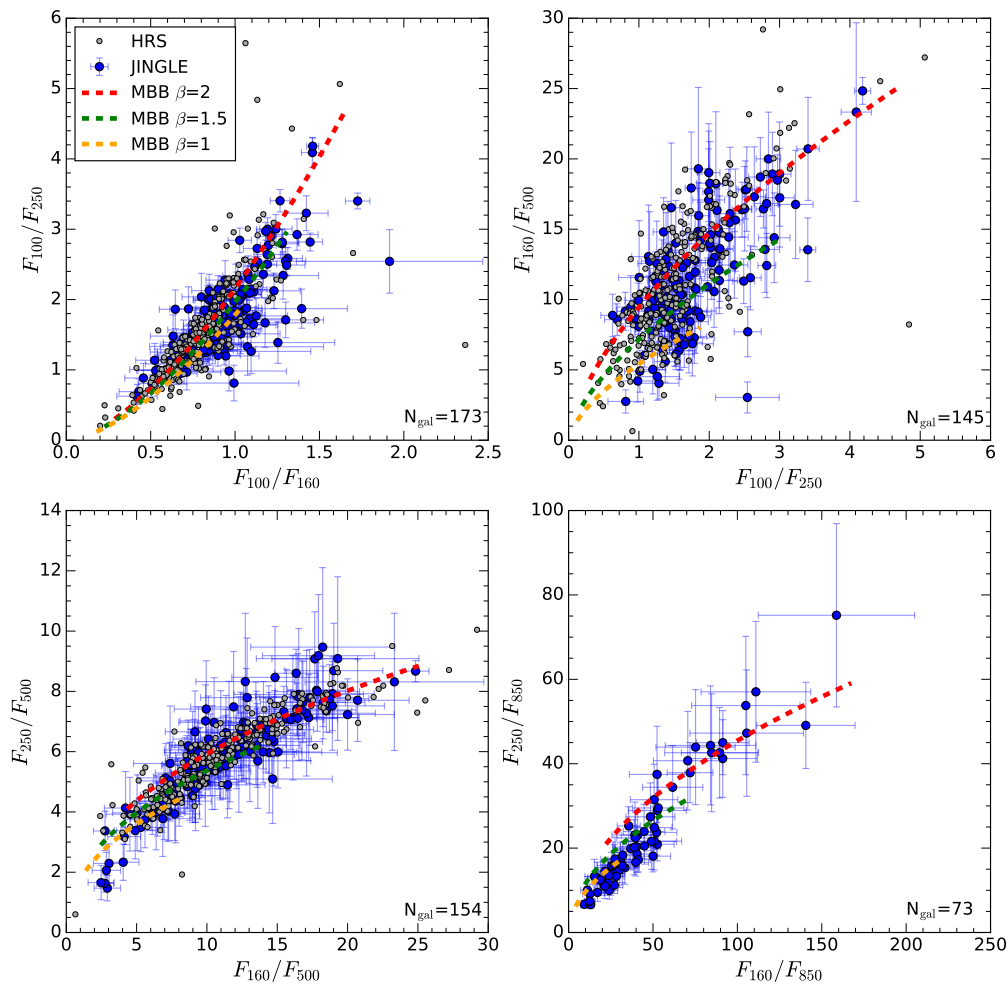


Figure 11. The four far-infrared/sub-mm colour-colour plots from Figure 10 with the strongest correlations. The blue points with error bars are points with a signal-to-noise ratio greater than 3. The grey dots show the position of detected galaxies from the HRS with no photometry flags in the catalogue. The red, green, and orange lines are the flux-density ratios of a single-modified blackbody between a temperature of 15 and 30 K with a β of 2, 1.5 and 1, respectively.

tribution of the normalised distances should be a normal distribution. We then use the Anderson-Darling test⁶ to test the null hypothesis that our distribution is drawn from the normal distribution. The F_{160}/F_{500} versus F_{100}/F_{250} reject the $\beta = 1.5$ and $\beta = 2.0$ models as the test statistic of 1.11 and 4.64 is above the critical level of 1.06 and so the probability of our data being drawn from the distribution is $< 1\%$. The F_{250}/F_{850} versus F_{160}/F_{850} also rejects the the $\beta = 2.0$ model with a probability $< 1\%$, but is not as significant for $\beta = 1.5$ with $\sim 5\%$ probability the data is drawn from the distribution. However, the statistics with F_{850} are likely to be an under-estimate of the model rejection, as a plot of the residuals (i.e., data - model), shows a strong correlation ($\rho > 0.6$) and significant (p-values $< 5 \times 10^{-9}$) as measured by the Spearman rank coefficient.

Our results confirms the findings of Boselli et al. (2012) that a single modified blackbody model with a constant β

cannot explain the distribution of points in both the *Herschel* wavelengths (i.e., the top-right panel in Figure 11), and the SCUBA-2 850 μm data. While both single β models are rejected by the *Herschel* data, values between 1.5–2.0 would explain the majority of data points. The bottom-left panel of Figure 11, shows many colour-ratios lie below the $\beta = 1.5$ line. It is possible to explain some of the systematic shift in the bottom-right panel by having an extreme calibration uncertainty for PACS and SPIRE, or by requiring a large colour-correction. However, a large calibration shift of the *Herschel* data would lead to worse agreement in the other panels, and can only partly resolve the offset. The colour corrections for both PACS and SPIRE are small ($\approx 3\%$) for a typical galaxy spectrum, and cannot resolve the offset. A potential solution is that a single dust component may not be a good model for a galaxy’s spectrum and multiple temperatures along the line-of-sight are required. Clark et al. (2015) performed an investigation of a blind dust survey of nearby galaxies ($15 < D < 46$ Mpc) using *H-ATLAS* (the HAPLESS sample), and so has similar detection criteria to that of JINGLE. For their sample they found that they required a two

⁶ We use the scipy implementation of the Anderson-Darling test (Jones et al. 01).

temperature component distribution to adequately describe their far-infrared fluxes, with many galaxies in their sample having a $\sim 9\text{--}15\text{ K}$ cold component. Using their best fitted parameters, a couple of the objects overlap with the extreme bottom left of the JINGLE distribution in a plot of F_{250}/F_{850} versus F_{160}/F_{850} , the vast majority of HAPLESS galaxies lie between our single modified blackbody with $\beta = 1.5$ and $\beta = 2.0$. However, the HAPLESS SED fits were performed assuming a $\beta = 2$ in the $60\text{--}500\ \mu\text{m}$ wavelength range, and so extrapolating their models may not give the true HAPLESS flux.

A broken emissivity modified-blackbody model where the emissivity law changes β value at a break wavelength (where $\lambda_{\text{break}} > 175\ \mu\text{m}$, see Li & Draine 2001; Gordon et al. 2014, for more details of the model), can be used to provide a better fit. In particular the cluster of objects with low F_{250}/F_{850} versus F_{160}/F_{850} values in Figure 11 could be explained, although the break wavelength is at longer wavelengths (e.g., $\approx 500\ \mu\text{m}$). Even if such a model could explain the excess at $850\ \mu\text{m}$, it is clear one dust model cannot explain the entire distribution of JINGLE galaxies.

Another possibility is that our signal-to-noise cut of 3 is artificially selecting galaxies that have been boosted to high $850\ \mu\text{m}$ flux densities. However, lowering the signal-to-noise cut to 2 results in the same observed discrepancy.

In the next section we avoid this issue as for each galaxy we use the results of fitting a modified blackbody model to each galaxy in JINGLE, as we have performed aperture matched photometry across all bands, even a low signal-to-noise $850\ \mu\text{m}$ measurement can be useful to constrain the fitted parameters.

6.3 Far-infrared Colours as indicators of dust temperature and emissivity

We investigate the effect of including the $850\ \mu\text{m}$ flux on the measurement of the effective emissivity index β . We perform the SED fitting with β as a free parameter, and we compared β measured from the fit with and without the $850\ \mu\text{m}$ flux (Figure 12). The points are scattered predominantly around the one-to-one relation. The largest discrepancies are for galaxies that have β values < 0 when obtained using *Herschel* data only, presumably due to low signal-to-noise in the longer SPIRE wavelengths. When our $850\ \mu\text{m}$ measurements are included all but one of these objects have values above $\beta = 0$ (values of β below 0 are not physically possible).

The KDE panels in Figure 12 show the distribution of β 's measured for the majority of objects is similar before and after the $850\ \mu\text{m}$ is included. However, 48% of the sample have estimates of β below 1.5, which may be an indication of a submm excess, (i.e., an excess emission at wavelengths $\geq 500\ \mu\text{m}$ with respect to what models would predict from the $100\text{--}350\ \mu\text{m}$ wavelength range). Such an excess has been observed in dwarf galaxies (Lisenfeld et al. 2002; Galliano et al. 2003), in late-type galaxies (Dumke et al. 2004; Bendo et al. 2006; Galametz et al. 2009; Relaño et al. 2018), as well as in the Magellanic Clouds (Israel et al. 2010; Bot et al. 2010; Planck Collaboration et al. 2011; Gordon et al. 2014). Several explanations have been proposed including a very cold dust component ($< 10\text{ K}$), spinning dust (Anderson & Watson 1993; Draine & Lazarian 1998a,b), the Two-Level-

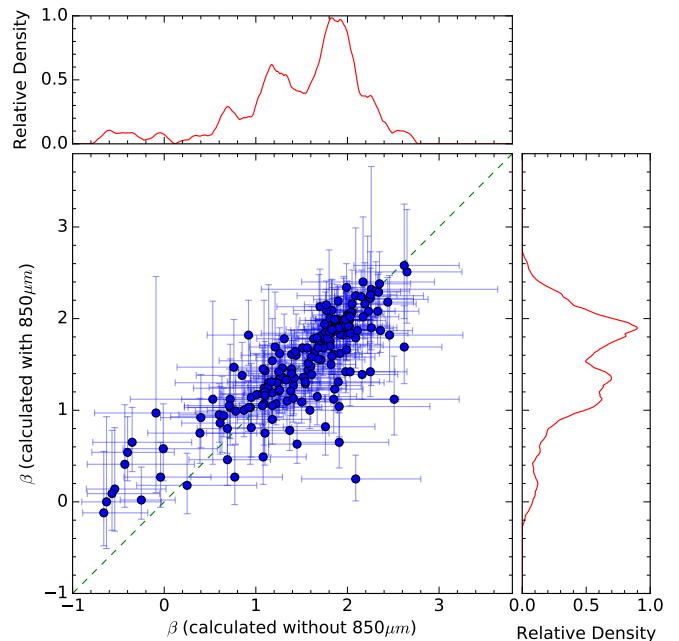


Figure 12. Comparison of the emissivity index β measured from the *Herschel* bands and measured with the additional SCUBA-2 $850\ \mu\text{m}$ flux point. The green dashed line shows the one-to-one relation (i.e., when the two estimates agree). The top and right panels show the KDE for each set of β points, which clearly show how the $850\ \mu\text{m}$ flux changes the distribution of values measured.

System of amorphous dust (Meny et al. 2007), or a broken dust emissivity law model (where the emissivity changes at a ‘break wavelength’ Li & Draine 2001). Explanations invoking very cold dust are problematic as they lead to very large estimated dust masses. In the LMC Gordon et al. (2014) found the broken emissivity law model had the lowest residuals. These models will be investigated in detail in Paper V (Lamperti et al. *in prep*).

The FIR colours that we have been studying can be used as an indicator of the cold dust temperature T_{dust} and emissivity index β (e.g. Boselli et al. 2010b; Boselli et al. 2012; Dale et al. 2012; Bendo et al. 2012; Galametz et al. 2010; Rémy-Ruyer et al. 2013; Cortese et al. 2014). We investigate which infra-red continuum flux ratios have the strongest correlation with T_{dust} and β , measured from a modified blackbody fit including the $850\ \mu\text{m}$ fluxes (see Figure 13). This can be useful for surveys that do not have the wavelength coverage to do full SED fitting.

Figure 13, shows the same FIR/sub-mm colour ratios versus our calculated dust emissivity index and dust temperature. In each panel we compute the Pearson correlation coefficient (ρ), for all objects with a signal-to-noise ratio greater than 3. The flux ratios F_{100}/F_{160} and F_{100}/F_{250} have the strongest correlations with the dust temperature ($\rho = 0.68$ and 0.58 , respectively). These flux bands are sampling the peak of the SED, which for a typical dust temperature between $10\text{--}30\text{ K}$ is in the wavelength range $90\text{--}250\ \mu\text{m}$. There are also fairly strong negative correlations with $\rho = -0.53$ and -0.54 , between dust temperature and the F_{160}/F_{850} or F_{250}/F_{850} , respectively. We do not investigate these further, but these may arise due to a temperature- β

degeneracy (for more discussion see Paper V Lamperti et al. *in prep*).

The flux ratio F_{160}/F_{850} has the strongest correlation with the emissivity index β ($\rho = 0.82$), closely followed by the F_{250}/F_{850} with $\rho = 0.76$. These colour-ratios are sampling the Rayleigh-Jeans tail of the sub-millimeter SED, and therefore are a good proxy of β . We fit the distribution of FIR/submm colour (with signal-noise ratios greater than 3) to either dust temperature or β , assuming this model:

$$F_i/F_j = \gamma 10^{\alpha X} \quad (8)$$

where F_i and F_j are the fluxes in the FIR/submm bands, α and γ are the model parameters and X is either dust temperature or β . We fit this model rather than a straight line to $\log_{10}(F_i/F_j)$ as the flux uncertainties are assumed Gaussian in the linear regime. We perform the fits using the pyMC3 package (Salvatier et al. 2016), and incorporate uncertainties in both FIR/submm colour and dust model parameters. Table 3 shows the correlation coefficients, and the fitted model parameters to relate the FIR/submm colours with β or dust temperature. These fits can be used to estimate T_{dust} and β if only two flux points are available.

For comparison, in Table 3 we also include the correlation coefficients and fits possible with only *Herschel* data. For estimates of β we note the correlation is better if longer wavelength information is available (e.g., $\rho = 0.54$ for F_{250}/F_{500} versus $\rho = 0.79$ for F_{250}/F_{850}).

6.4 Far-infrared Colours Versus Galaxy Physical Parameters

In this final section we investigate if there are any relationships between the FIR/sub-mm colours and a few physical parameters of the galaxy from the quantities calculated in Paper I. We look for correlations with star-formation rate (SFR), stellar mass (M_*), specific star-formation rate (SFR/M_*), the surface density of SFR ($\Sigma(\text{SFR})$) and the surface density of stellar mass (Σ_*). The two surface densities are calculated based on taking the integrated quantity (i.e., SFR or M_*) and dividing by the elliptical area of the galaxy using the Petrosian radius (in kpc) and the axis ratio of each galaxy.

Figure 14 shows the FIR colours versus the various physical properties defined above. The best correlation with a Spearman rank coefficient of $\rho = 0.57$ is between $\Sigma(\text{SFR})$ and the F_{160}/F_{500} (closely followed by the F_{100}/F_{250} and F_{250}/F_{500}), and similar, but reduced correlations are also seen with the SFR. Correlations with $\Sigma(\text{SFR})$ and SFR would be expected due to the strength of the interstellar radiation field leading to an increase of the dust temperature, or equivalently far-infrared luminosity is used as a star formation tracer (e.g., Kennicutt 1998). Surprisingly, from Section 6.3 both the F_{160}/F_{500} and F_{250}/F_{500} colours are better tracers of β , rather than dust temperature. This will be further investigated in Lamperti et al, in prep. There are also some weak correlations with stellar mass and stellar mass surface density which again may be the effect of increasing the interstellar radiation field and therefore increasing the dust temperature, however, our results suggest for the majority of galaxies in JINGLE this is less significant than the heating from star formation.

Boselli et al. (2012) also compared their FIR colour-ratios for the HRS sample to a similar set of physical galaxy properties. For both samples a correlation is found between FIR colours at longer wavelength and the Σ_* (or equivalently in the HRS study H -band surface-brightness). Surprisingly, Boselli et al. (2012) did not find any correlations with SFR (although correlations with $H\alpha$ surface-density were found), this maybe due to the different methods of measuring SFR (dust extinction-corrected $H\alpha$ or FUV emission versus panchromatic SED fitting for JINGLE), or the greater fraction of early-type galaxies in the HRS.

7 CONCLUSIONS

In this paper we have presented new SCUBA-2 data of the JINGLE sample which consists of 193 galaxies in the range 41-212 Mpc. We described our data reduction tests and techniques which attempt to optimise the SCUBA-2 data for the case of slightly extended galaxies, going beyond what is performed in the standard SCUBA-2 pipeline. We investigate the optimum parameters, calibration and offset corrections for our data. By incorporating data from *Herschel* (both PACS and SPIRE) and WISE, we measure aperture-matched global fluxes across the dust SED. Our measurements attempt to account for aperture corrections, effects from the pipeline filtering in the SCUBA-2 wavelengths, all sources of noise in the FIR wavelengths (e.g., instrumental, confusion, cirrus, etc.), and the contamination from the CO($J=3-2$) line.

Using the *Herschel* data we find that the FIR/submm colours are similar to the HRS, and we find similar results to Boselli et al. (2012), that the distribution on colour-colour plots appears incompatible with a single modified blackbody with a constant β , but the majority of galaxies could be described with β in the standard range of 1.5–2.0. By adding our 850 μm fluxes, we find that approximately half of the JINGLE objects require a value of β significantly lower than 1.5. Such low values are hard to reproduce with lower temperatures and so possibly indicate a more complicated dust model like those that have a broken-emissivity law is required. The distribution of JINGLE galaxies confirms one dust model cannot explain the entire sample. These models applied to individual objects will be further investigated by Lamperti et al, in prep. who will use a Bayesian hierarchical fitting approach.

We found that the F_{160}/F_{850} and the F_{250}/F_{850} colours have the strongest correlation ($\rho = 0.84$ and 0.77 , respectively) with β estimated from a single modified blackbody model. The dust temperature is better correlated with shorter wavelength colours from *Herschel* data, with the highest correlation found for the F_{100}/F_{160} colour ($\rho = 0.71$). We provide the fits to these plots to find an estimate of β and temperature from a FIR/submm colour.

Finally, we investigate how the FIR/submm colours vary with different physical parameters presented in Paper I. The best correlation is between the FIR/submm colour and the surface-density of star-formation rate (closely followed by the total star-formation rate), indicating that for JINGLE galaxies dust heating is predominantly due to young-stellar populations rather than older stellar populations. However, there is a significant but reduced correlation with

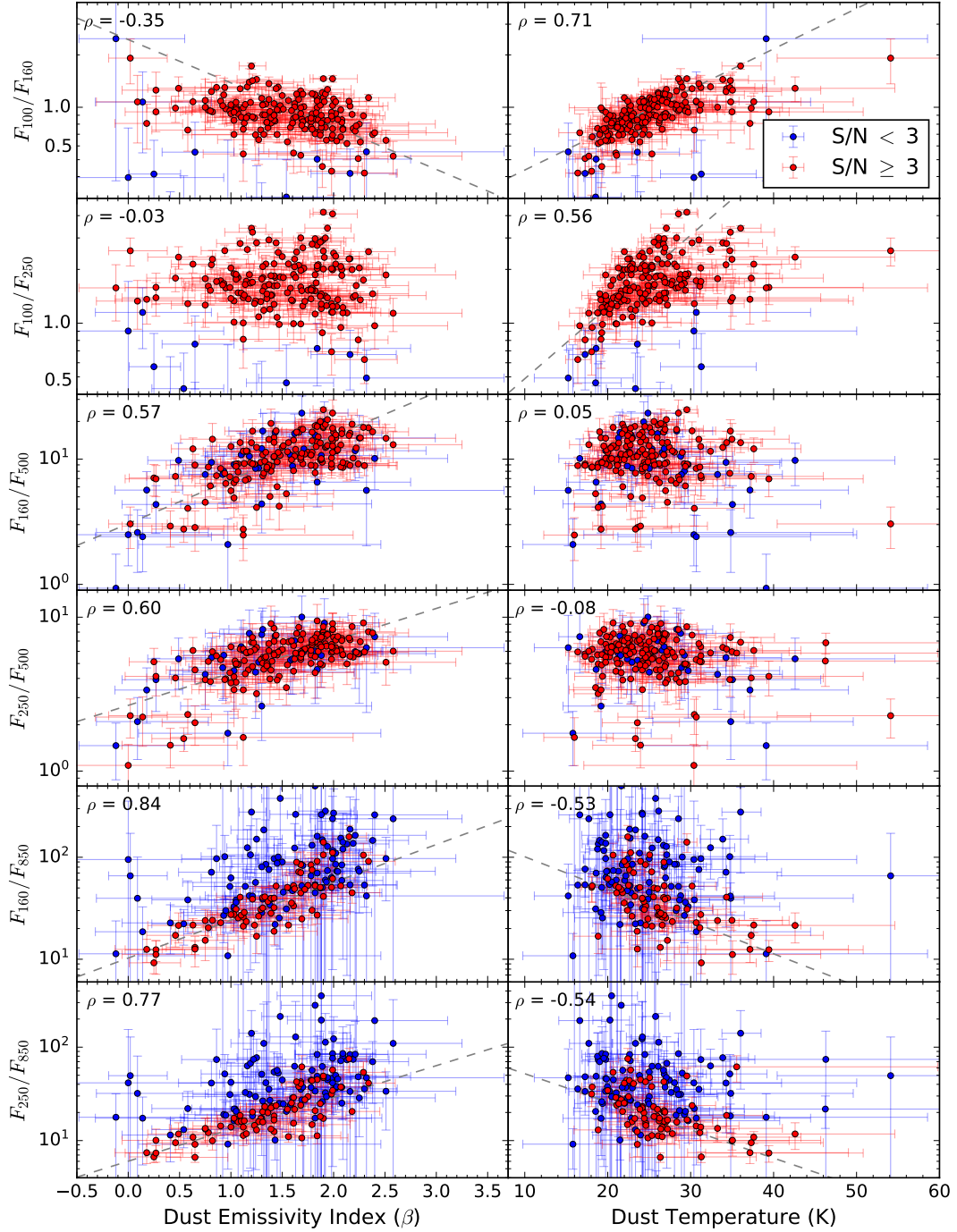


Figure 13. Correlation between flux ratios and the dust emissivity β (left) and dust temperature (right). In blue are shown points for which one of the flux measurements has high uncertainties ($S/N < 3$). The Pearson correlation coefficients ρ for the objects with a signal-to-noise greater than 3 is shown in the top-left of each panel. The flux ratios F_{100}/F_{160} and F_{100}/F_{250} have the strongest correlation with the dust temperature ($\rho = 0.68$ and $\rho = 0.58$, respectively), although there are also negative correlations with the F_{160}/F_{850} and F_{250}/F_{850} . The flux ratio F_{160}/F_{850} and F_{250}/F_{850} show the best correlation with β (with $\rho = 0.82$ and 0.76 , respectively). The grey lines show the best fit models specified in Table 3. Due to the large scatter of the lower signal-to-noise data points, a few objects may lie outside the plotted range.

Table 3. Correlation between flux ratios and dust properties (temperature and β).

Flux Ratio	Dust Emissivity Index (β)			Dust Temperature		
	ρ	α	γ	ρ	α (K ⁻¹)	γ
F_{100}/F_{160}	-0.35	-0.248 ± 0.016	2.45 ± 0.15	0.71	0.0232 ± 0.0013	0.256 ± 0.019
F_{100}/F_{250}	-0.03	-	-	0.56	0.0402 ± 0.0190	0.194 ± 0.020
F_{160}/F_{500}	0.57	0.349 ± 0.019	3.06 ± 0.22	0.05	-	-
F_{250}/F_{500}	0.60	0.210 ± 0.013	2.67 ± 0.14	-0.08	-	-
F_{160}/F_{850}	0.84	0.371 ± 0.031	10.3 ± 1.01	-0.53	-0.0317 ± 0.0036	210.0 ± 52.6
F_{250}/F_{850}	0.77	0.305 ± 0.030	6.02 ± 0.56	-0.54	-0.0301 ± 0.0035	104.0 ± 25.4

Notes. The table provides the Spearman rank coefficient (ρ), and model parameters for our distributions of FIR/submm colours versus fitted dust parameters. For the model parameters described in Equation 8 the median of the posterior distribution is given, and its uncertainty estimated from the 16th and 84th percentile (the uncertainties are to a good approximation symmetrical). We only provide model fits for the parameters for distributions with $|\rho| > 0.3$.

stellar surface-density, suggesting some dust heating from older stellar populations occurs in JINGLE galaxies.

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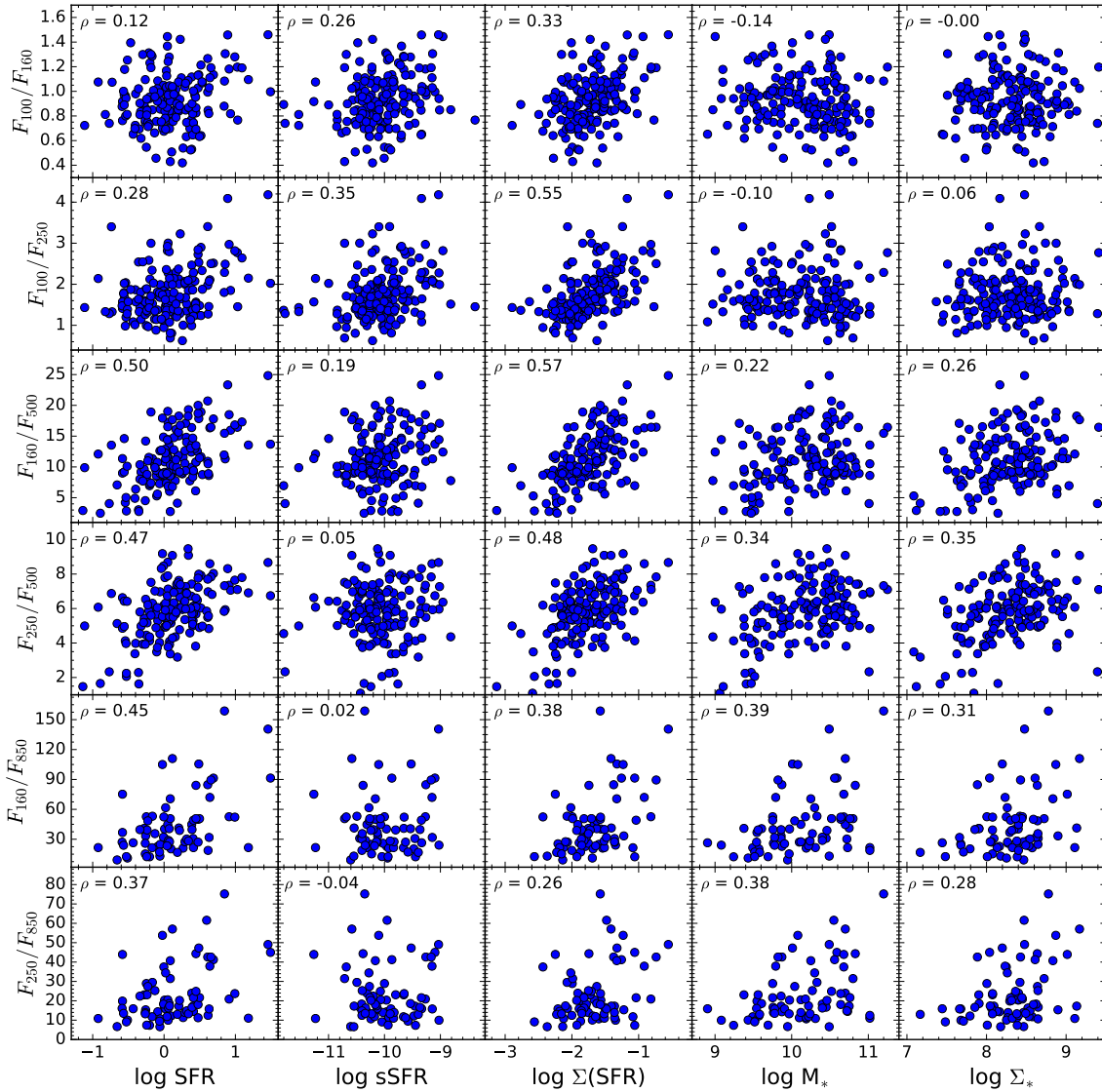


Figure 14. FIR colours versus the SFR, sSFR, $\Sigma(\text{SFR})$, M_* and Σ_* , as defined in Section 6.4. The Spearman rank correlation for each panel is shown in the top left hand corner. Data points are included if the flux densities used to create each FIR/sub-mm colour have a signal-to-noise of at least 3.

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